ΛCDM: Triumphs and Tribulations

Joel R. Primack
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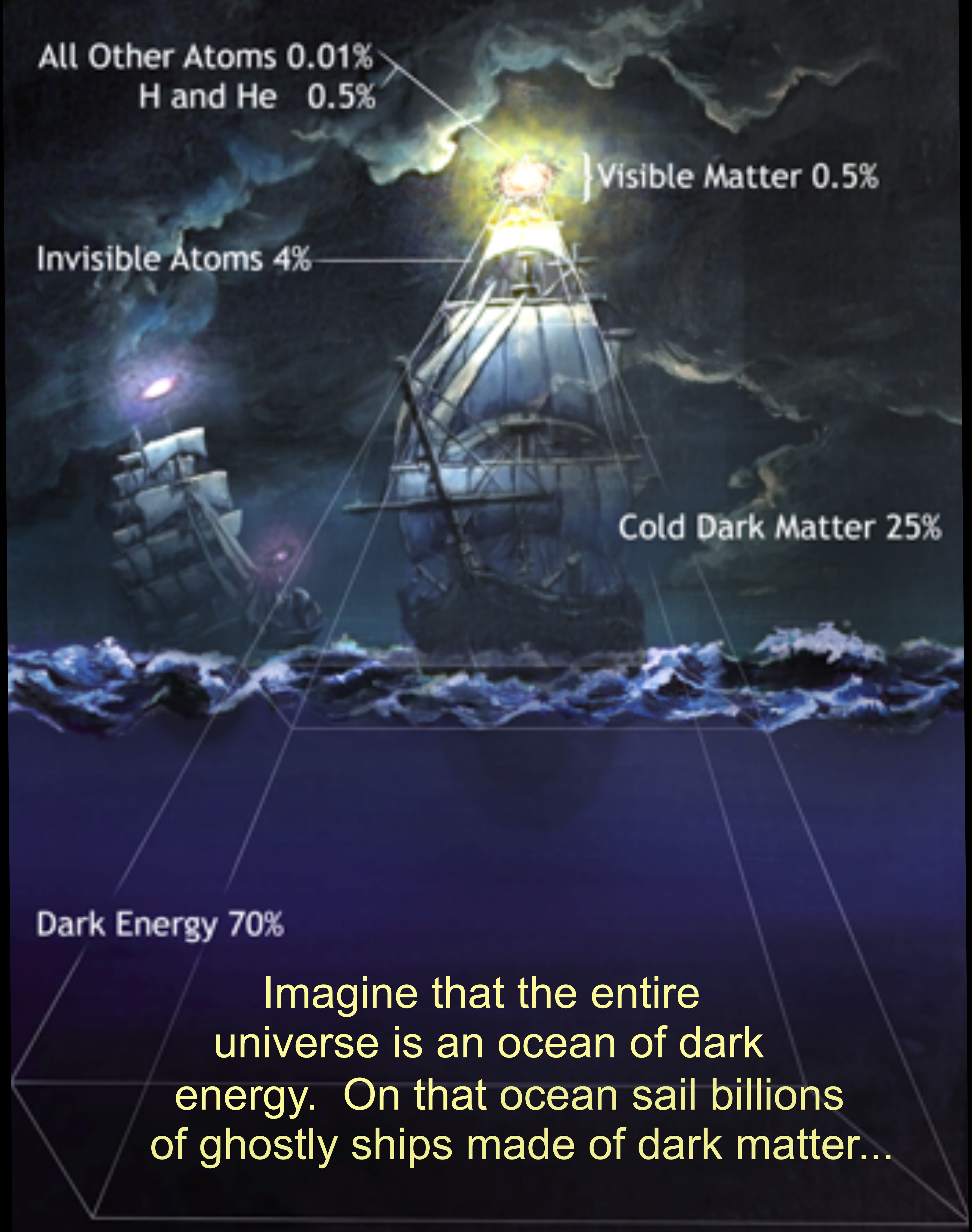
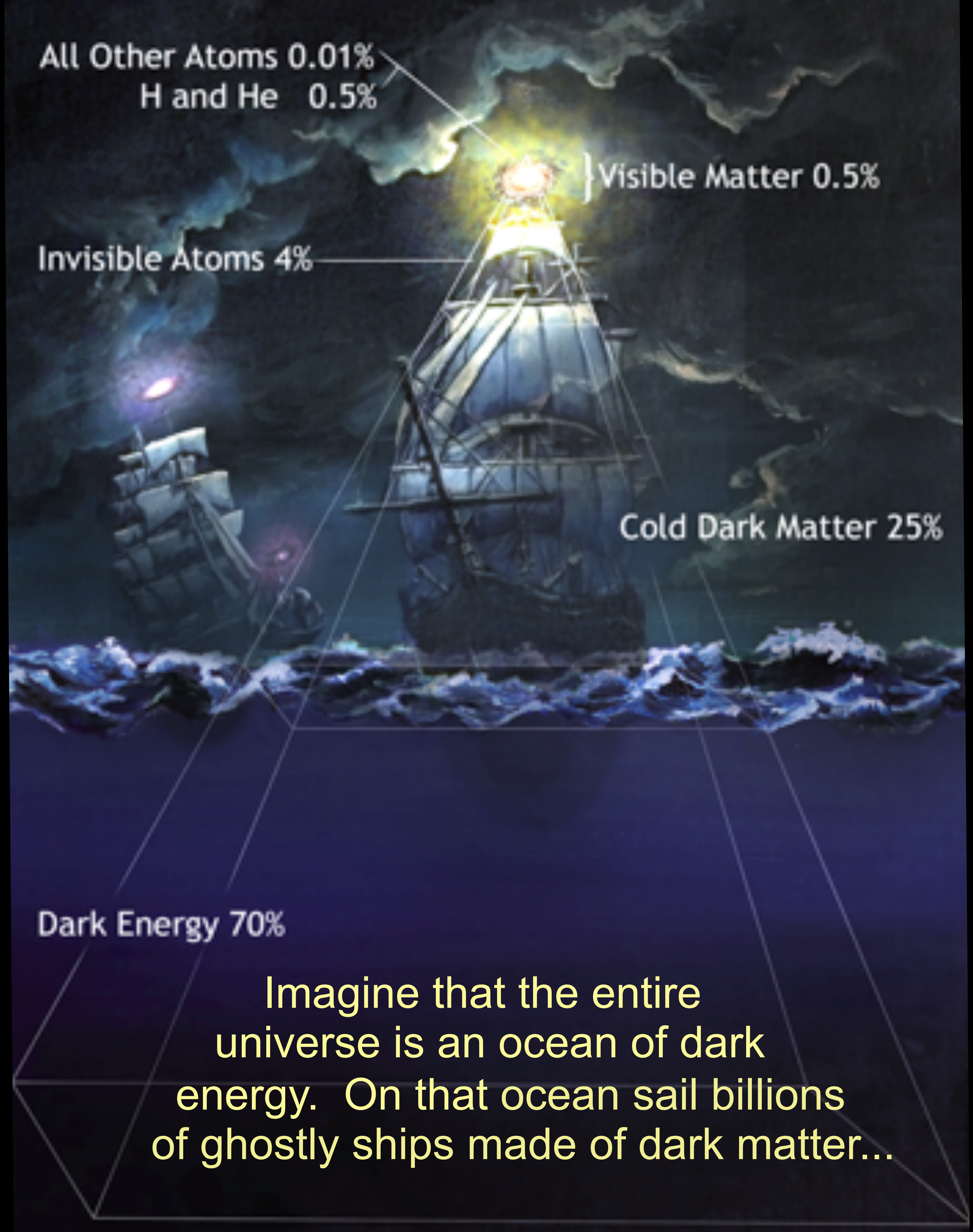
Matter and Energy Content of the Universe



Dark Matter Ships

on a

Dark Energy Ocean



Matter and Energy Content of the Universe

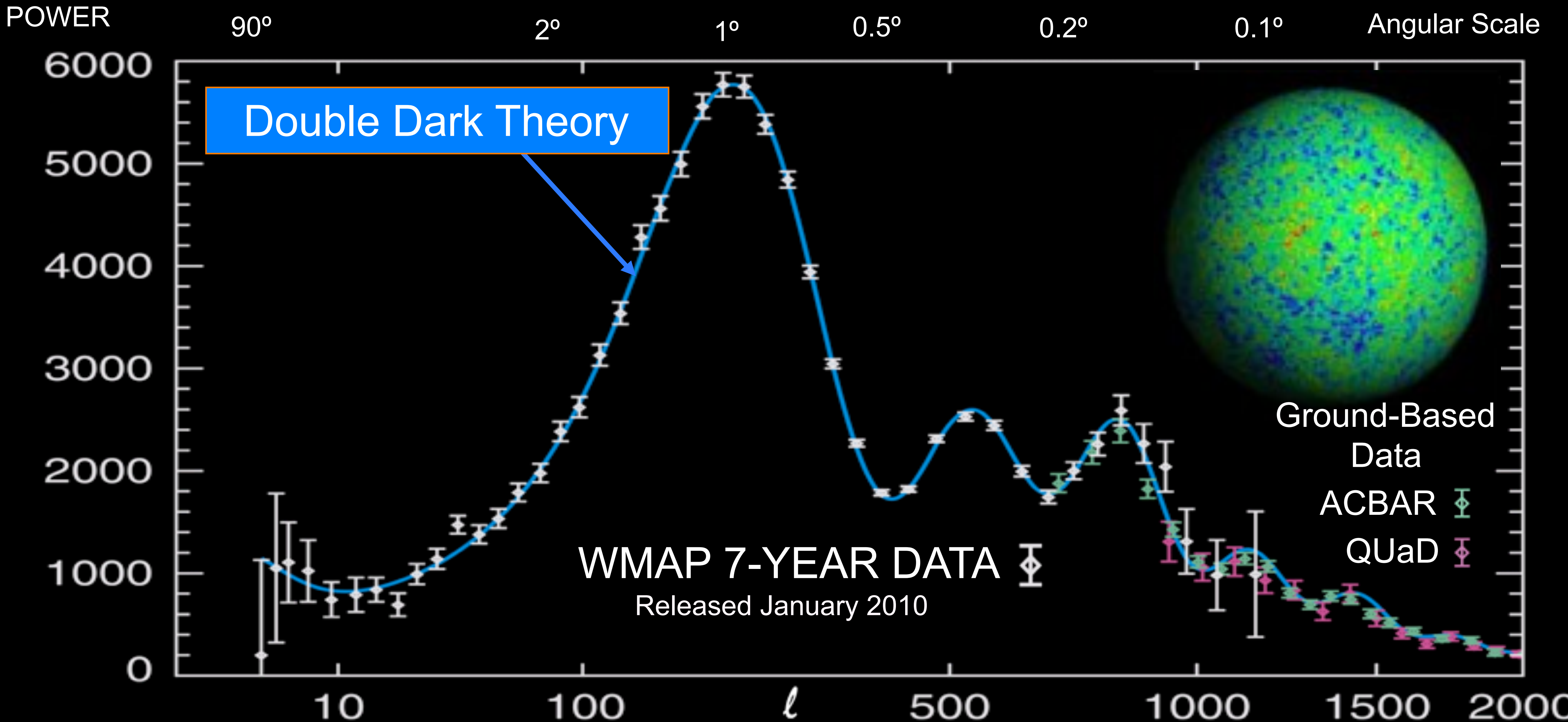
ΛCDM

Double Dark Theory

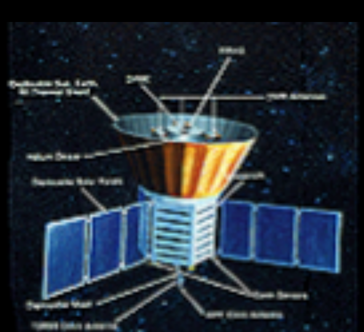
DARK MATTER + DARK ENERGY = DOUBLE DARK THEORY

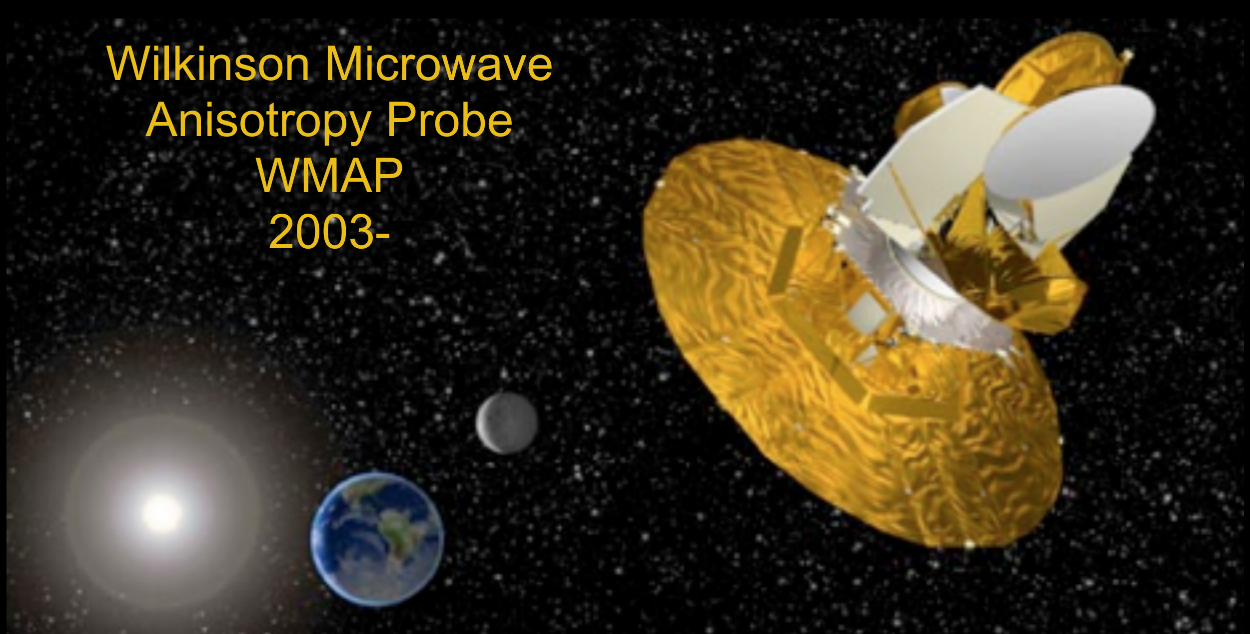
Technical Name:
Lambda Cold Dark Matter (ΛCDM)

Big Bang Data Agrees with Double Dark Theory!



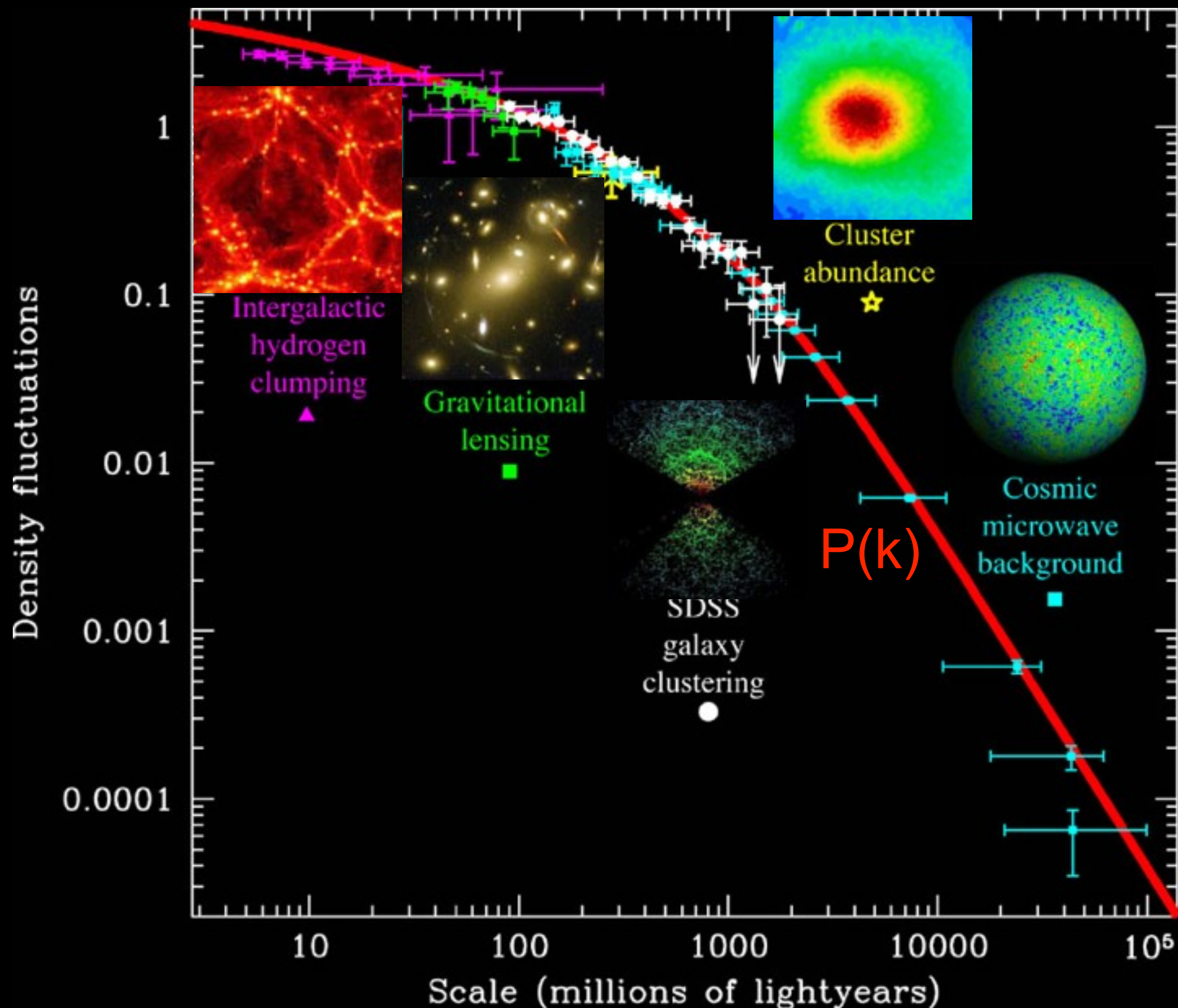
Cosmic Background Explorer COBE 1992



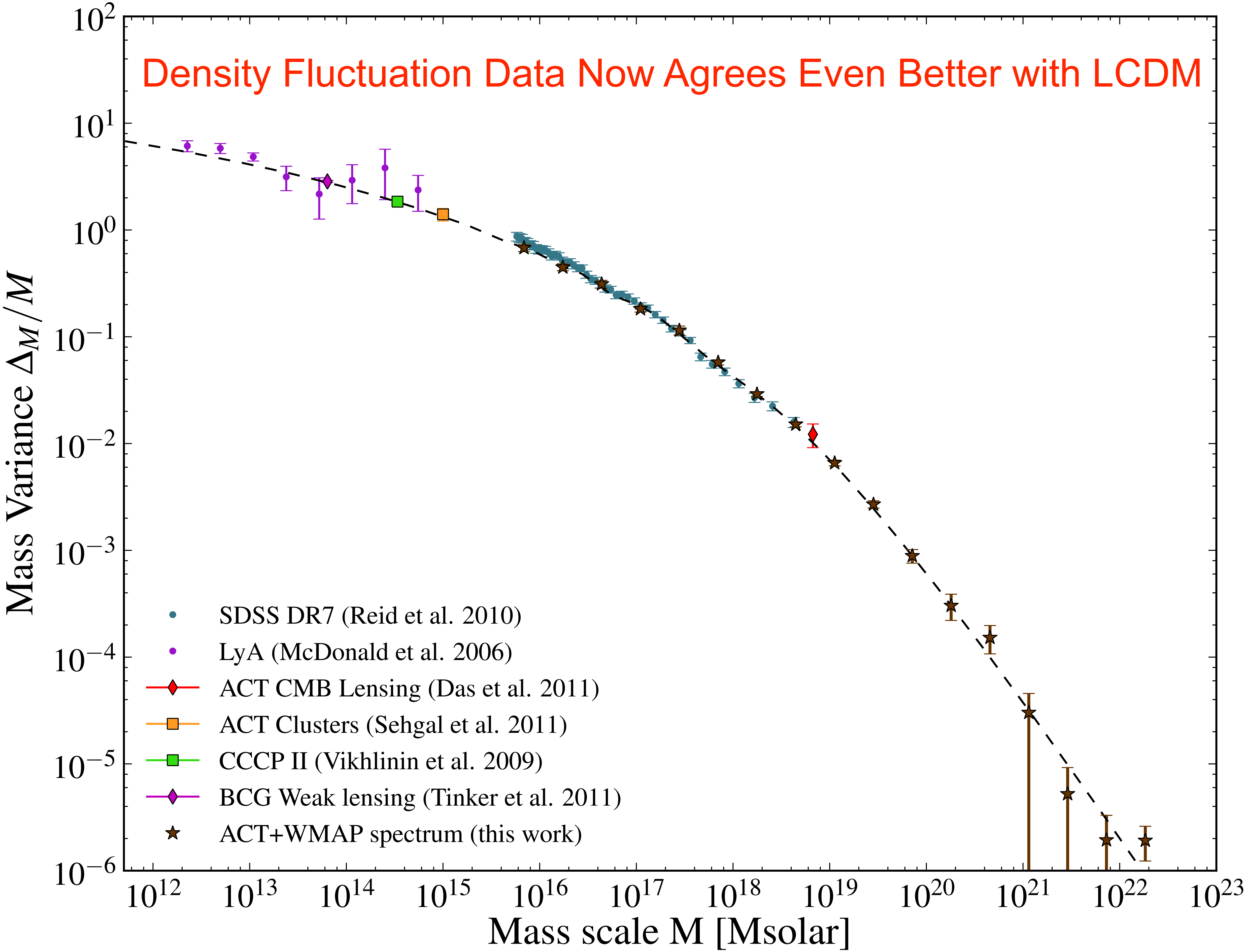


Distribution of Matter

Also Agrees with Double Dark Theory!



Max Tegmark
2003



R. Hlozek et al 2011 The Atacama Cosmology Telescope: a measurement of the primordial power spectrum arXiv:1105:4487

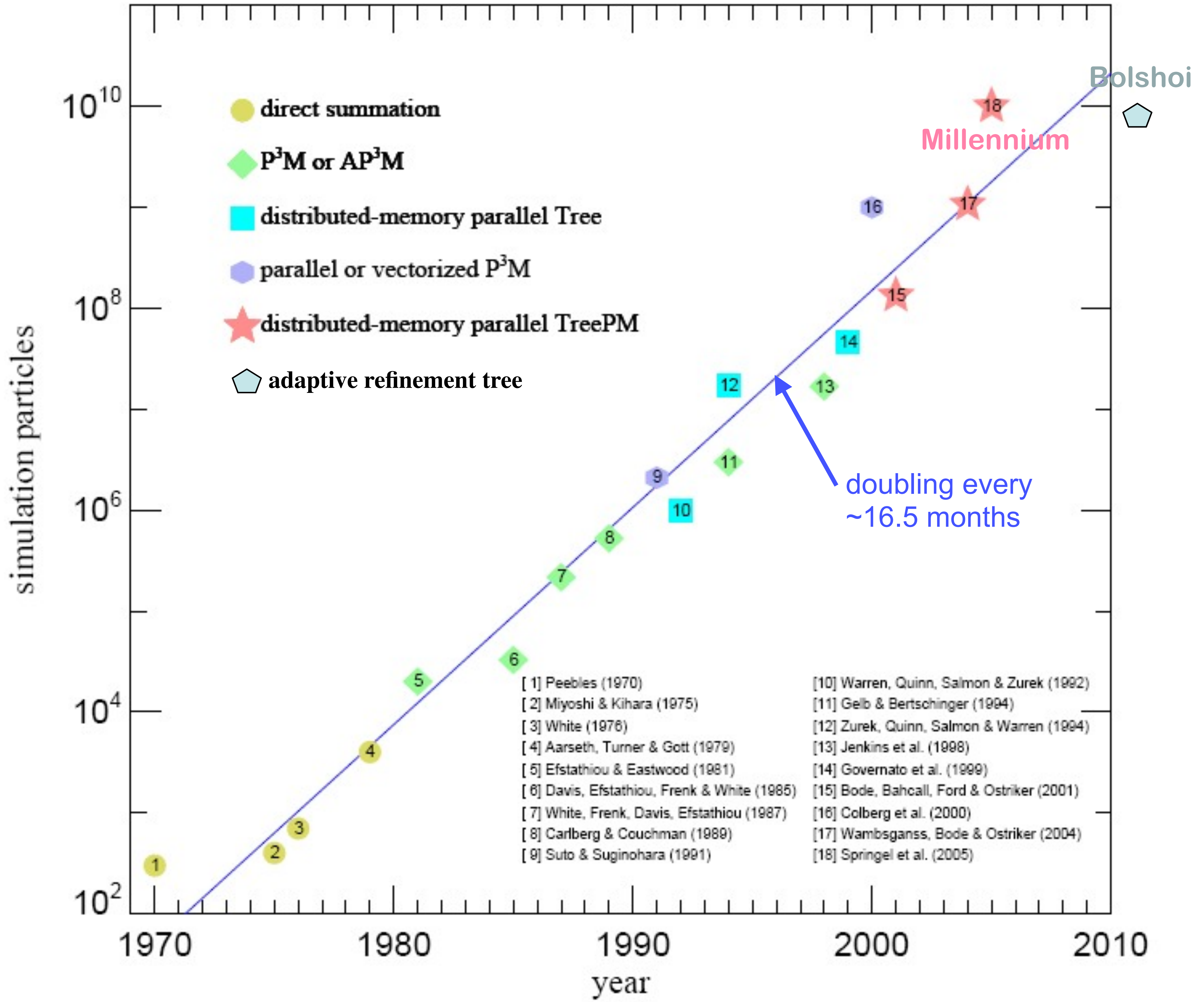
Cosmological Simulations

Astronomical observations represent snapshots of moments in time. It is the role of astrophysical theory to produce movies -- both metaphorical and actual -- that link these snapshots together into a coherent physical theory.

Cosmological dark matter simulations show large scale structure, growth of structure, and dark matter halo properties

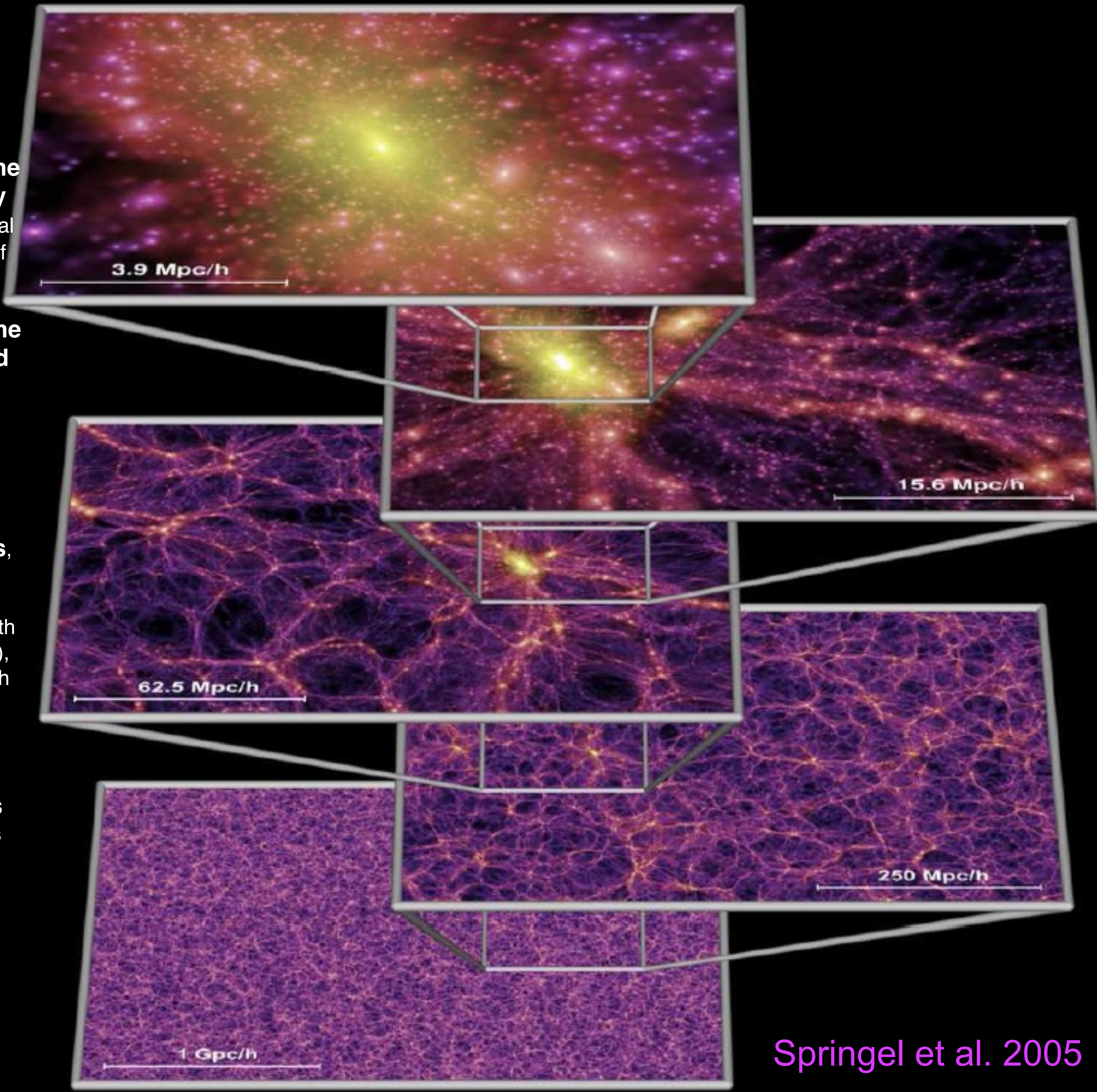
Hydrodynamic galaxy formation simulations: evolution of galaxies, formation of galactic spheroids via mergers, galaxy images in all wavebands including stellar evolution and dust

Particle number in N-body simulations vs. publication date



The Millennium Run

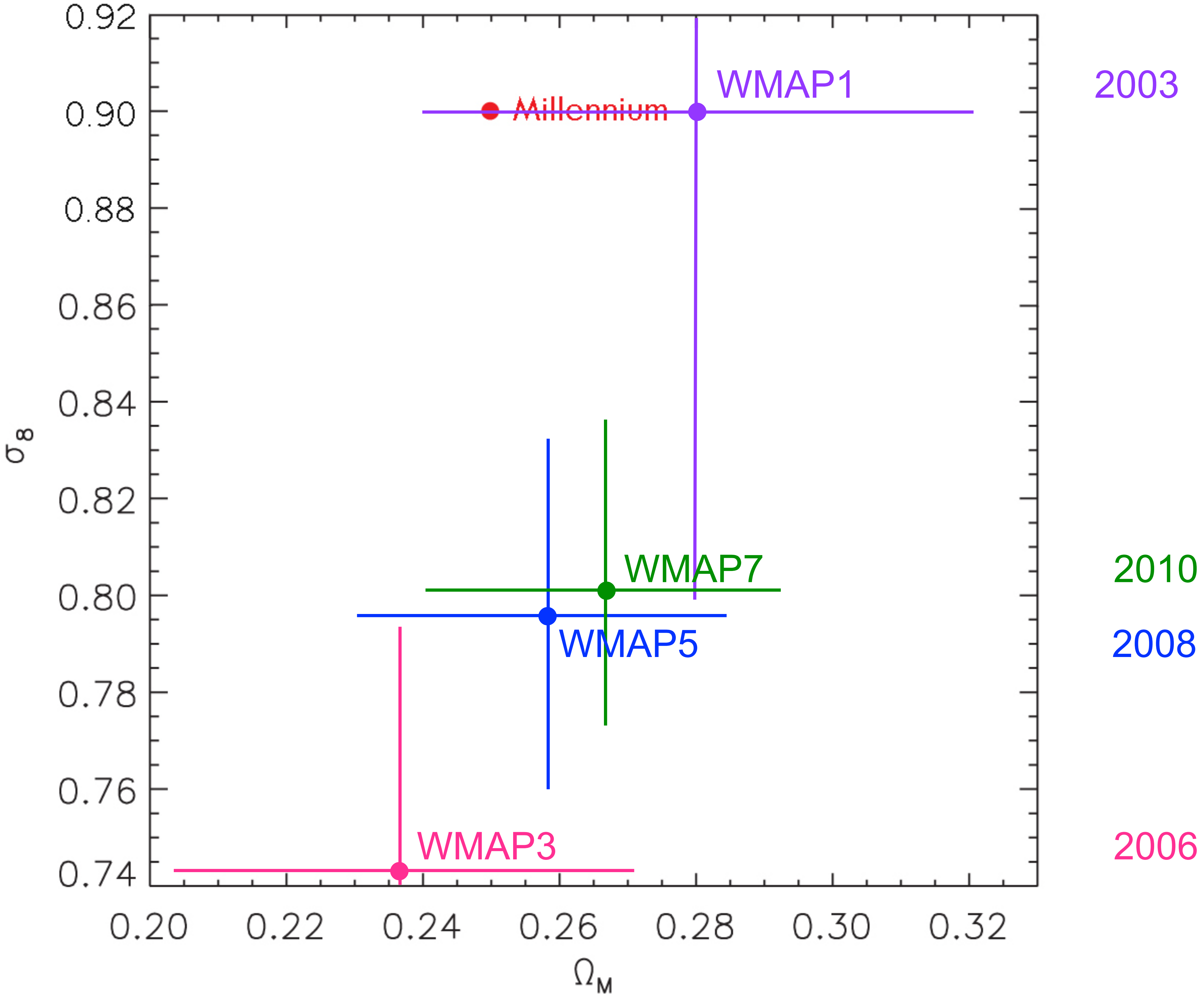
- **properties of halos** (radial profile, concentration, shapes)
- **evolution of the number density of halos**, essential for normalization of Press-Schechter-type models
- **evolution of the distribution and clustering of halos** in real and redshift space, for comparison with observations
- **accretion history of halos**, assembly bias (variation of large-scale clustering with as- sembly history), and correlation with halo properties including angular momenta and shapes
- **halo statistics** including the mass and velocity functions, angular momentum and shapes, subhalo numbers and distribution, and correlation with environment



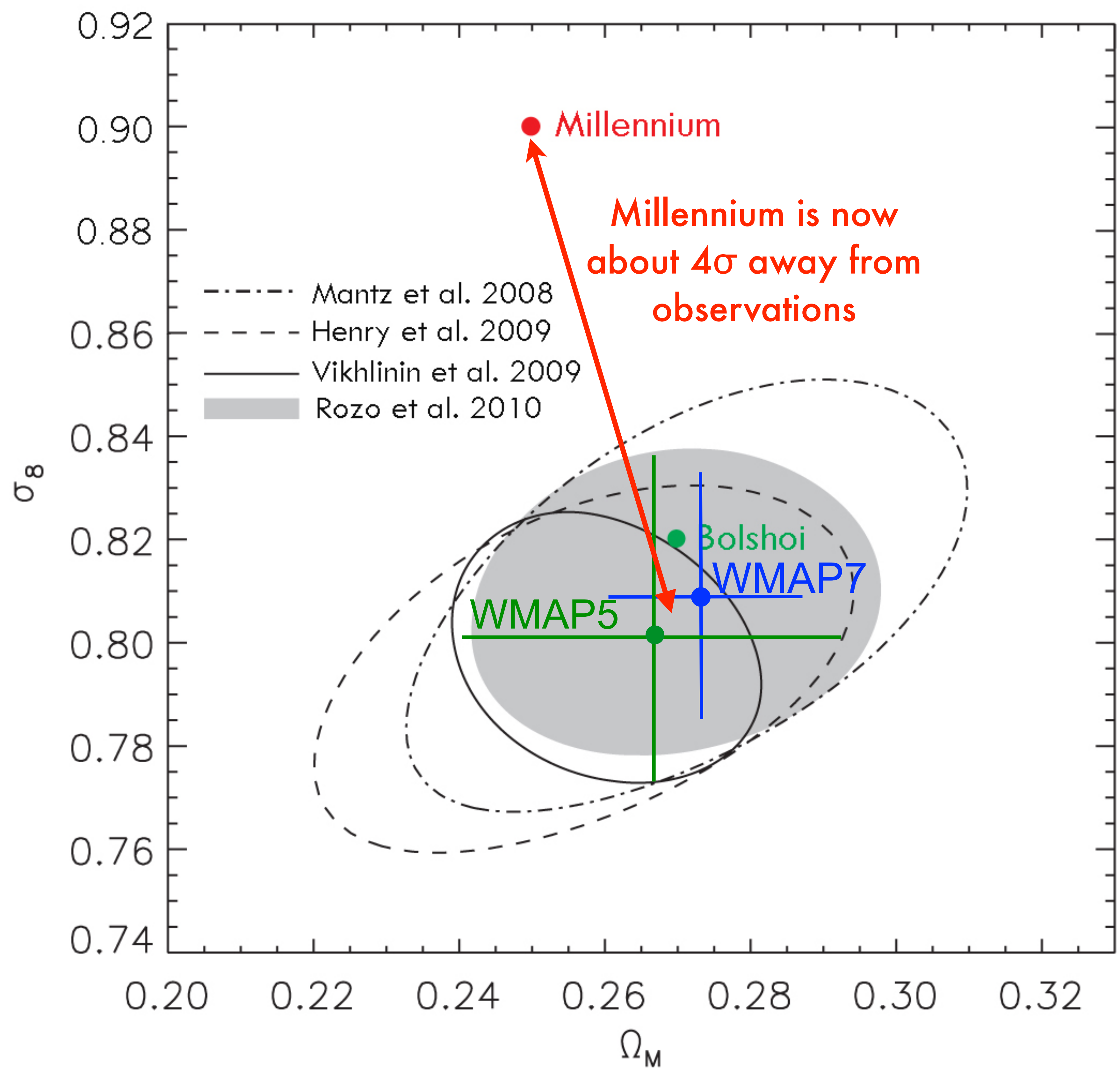
Springel et al. 2005

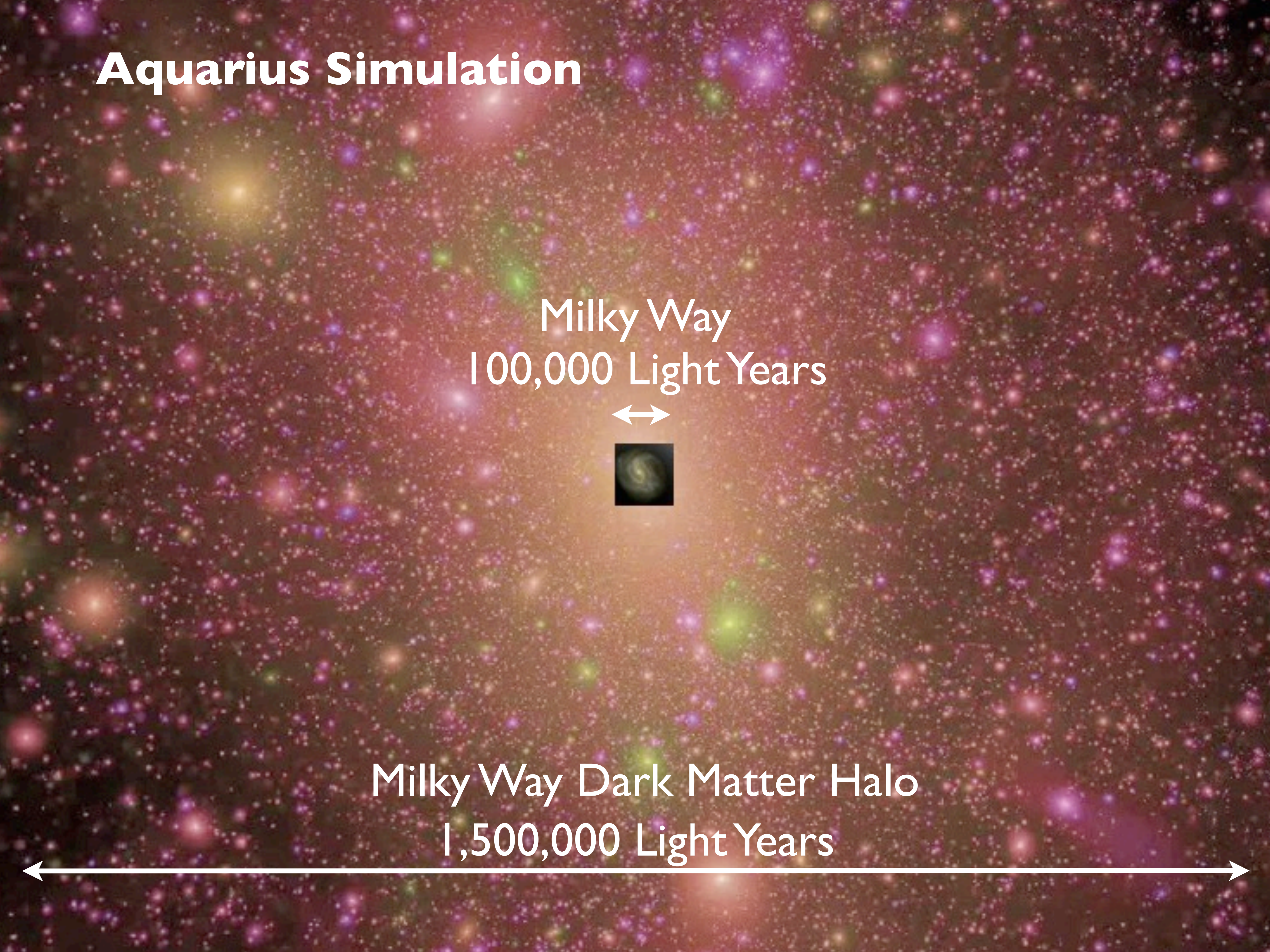
- **void statistics**, including sizes and shapes and their evolution, and the orientation of halo spins around voids
- quantitative descriptions of the evolving **cosmic web**, including applications to weak gravitational lensing
- preparation of **mock catalogs**, essential for analyzing SDSS and other survey data, and for preparing for new large surveys for dark energy etc.
- **merger trees**, essential for **semi-analytic modeling** of the evolving galaxy population, including models for the galaxy merger rate, the history of star formation and galaxy colors and morphology, the evolving AGN luminosity function, stellar and AGN feedback, recycling of gas and metals, etc.

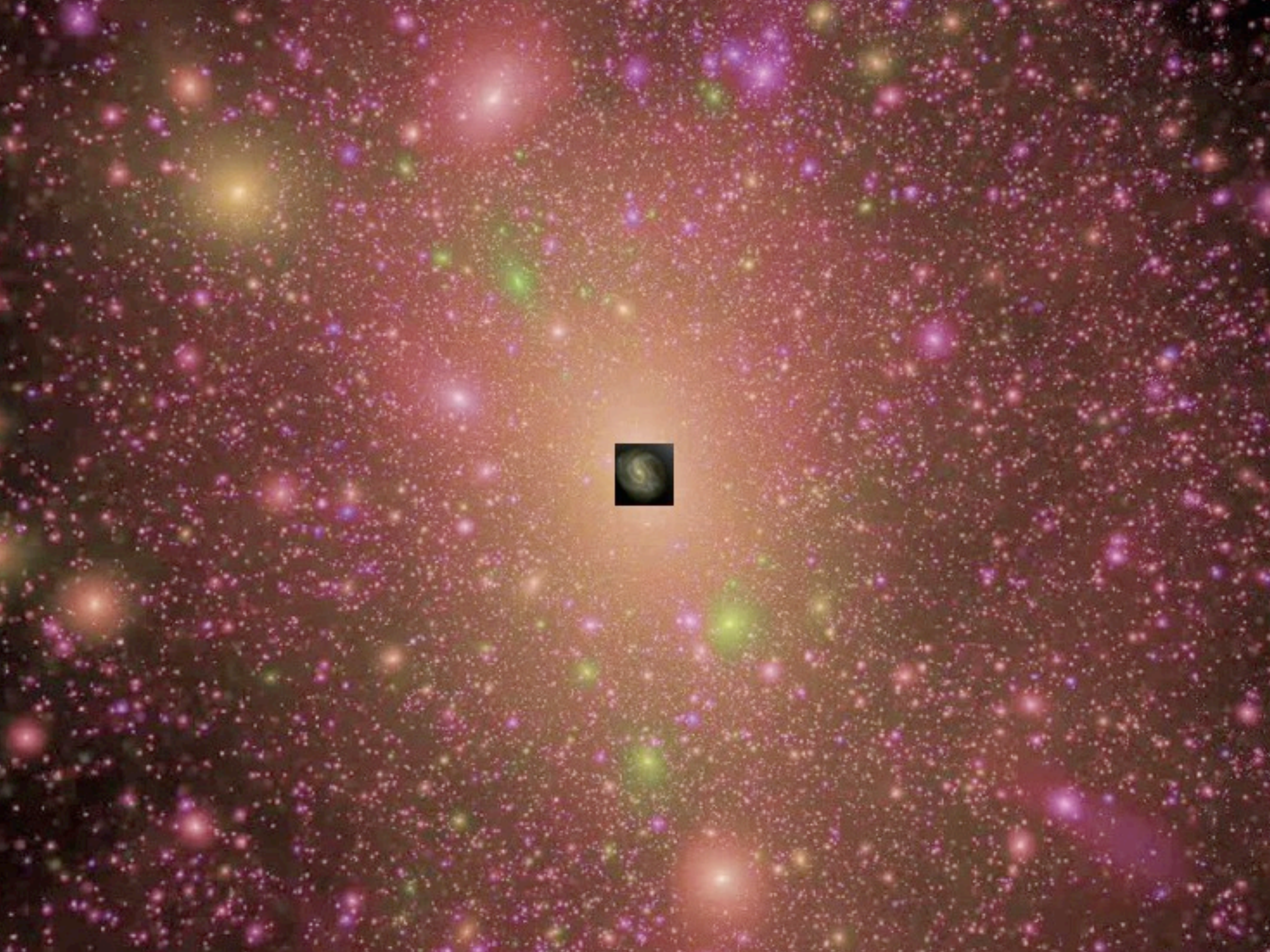
WMAP-only Determination of σ_8 and Ω_M

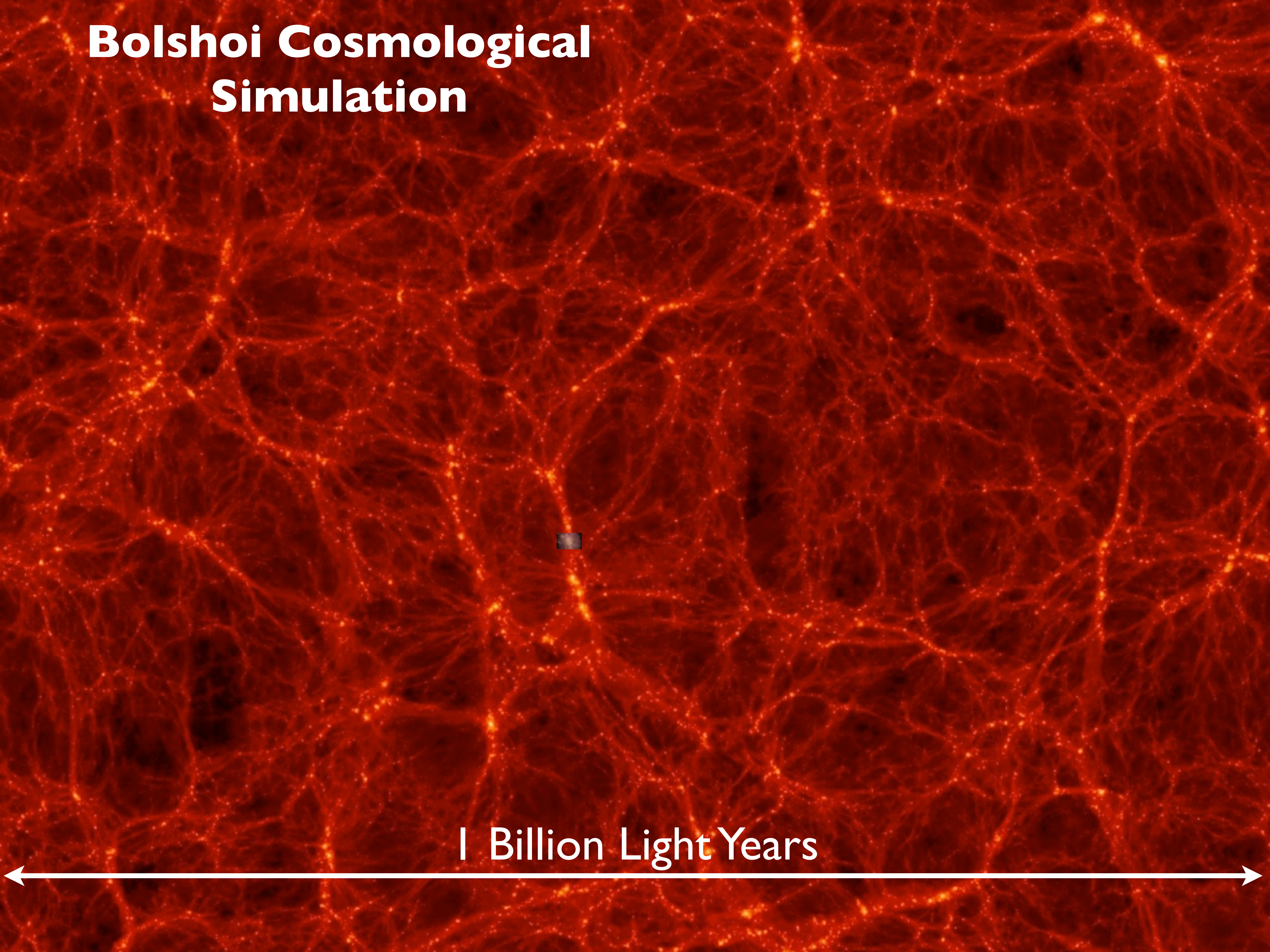


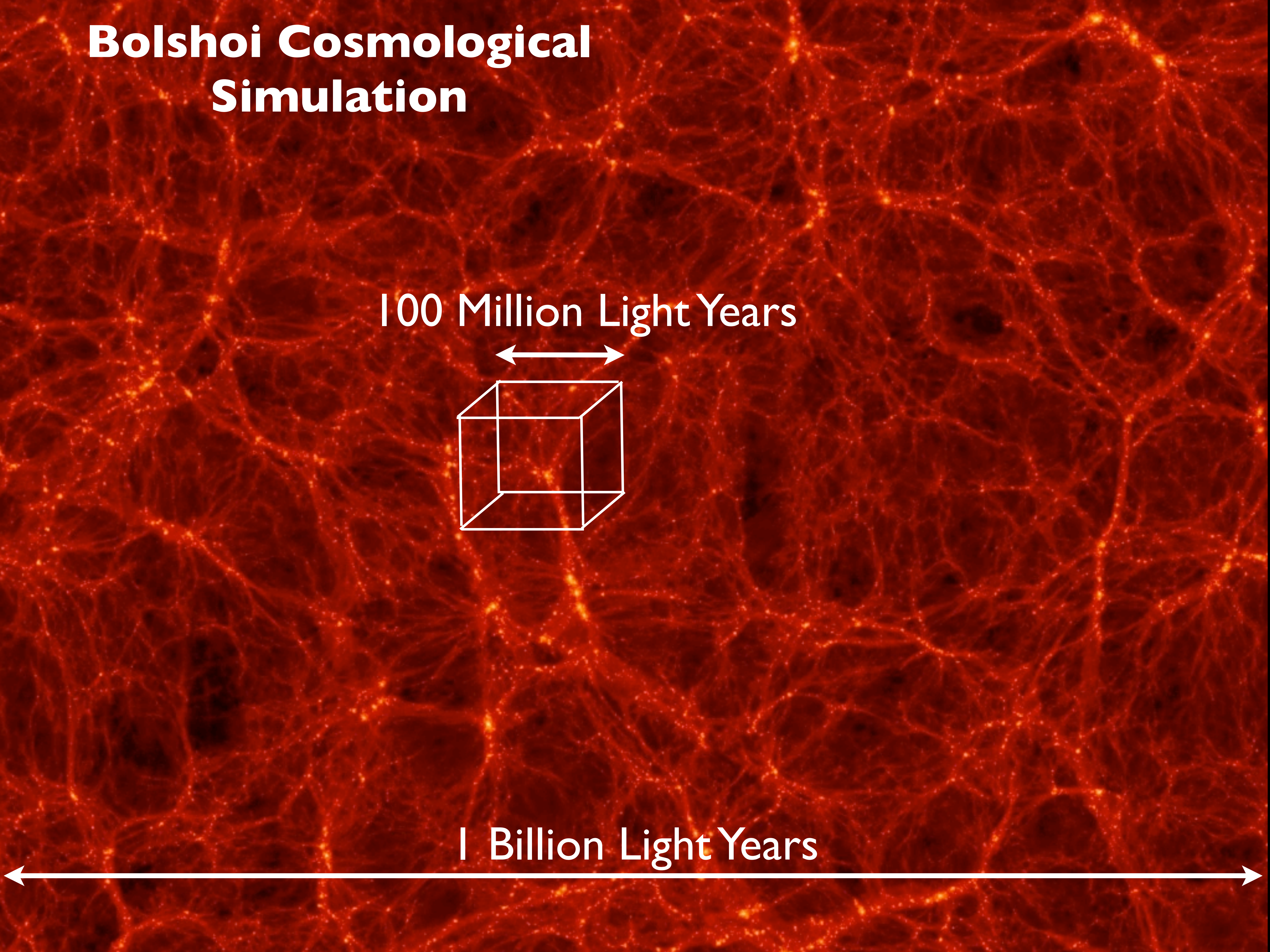
WMAP+SN+Clusters Determination of σ_8 and Ω_M




Aquarius Simulation
Milky Way
100,000 Light Years
Milky Way Dark Matter Halo
1,500,000 Light Years




Bolshoi Cosmological Simulation
1 Billion Light Years


Bolshoi Cosmological Simulation
100 Million Light Years
1 Billion Light Years

Bolshoi Cosmological Simulation

100 Million Light Years

Bjork “Dark Matter”

Biophilia

The Bolshoi simulation

ART code

250Mpc/h Box
LCDM
σ_8 = 0.82
h = 0.70

8G particles
1kpc/h force resolution
1e8 Msun/h mass res

dynamical range 262,000
time-steps = 400,000

NASA AMES
supercomputing center
Pleiades computer
13824 cores
12TB RAM
75TB disk storage
6M cpu hrs
18 days wall-clock time

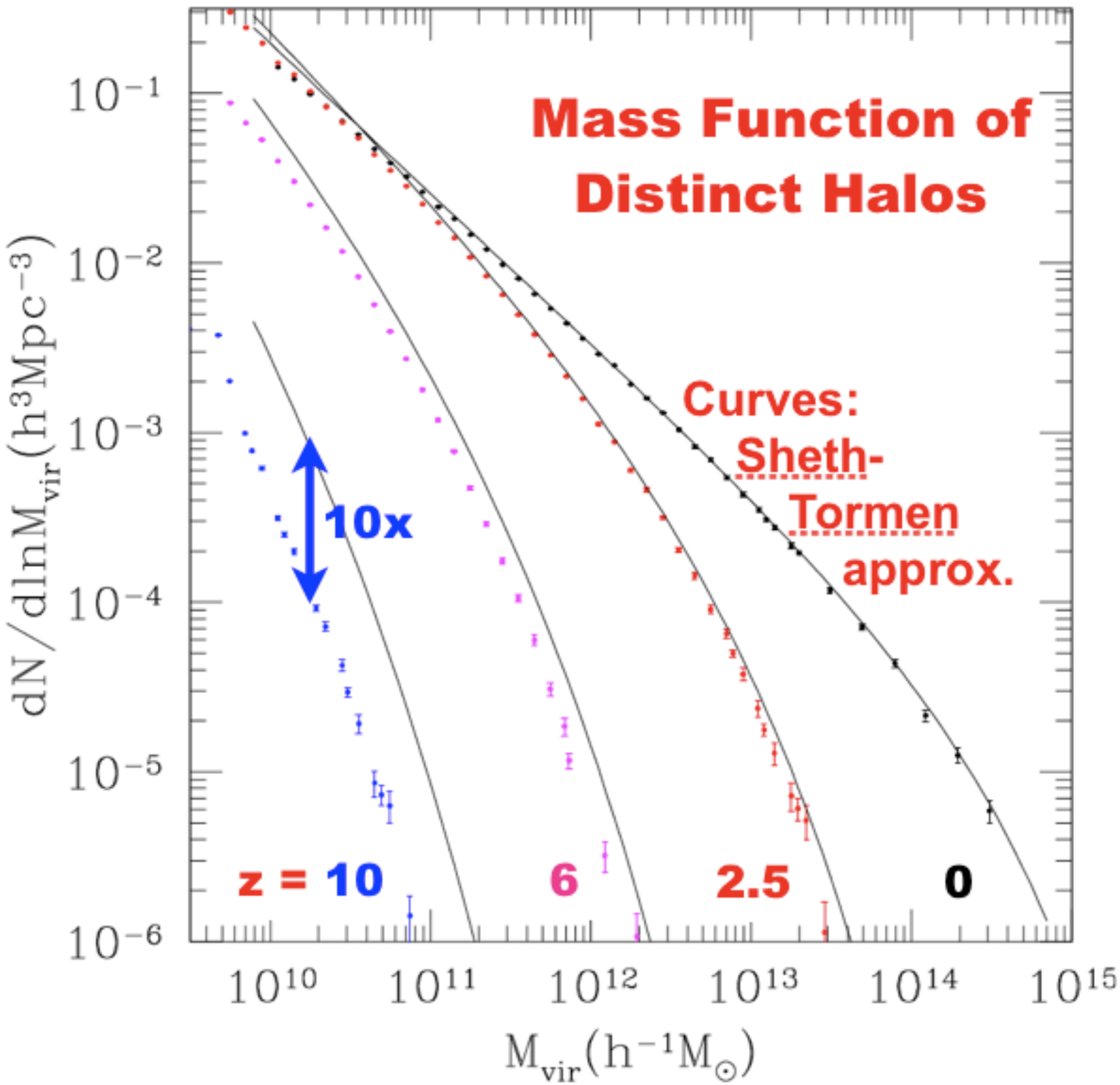
Cosmological parameters are consistent with the latest observations

Force and Mass Resolution are nearly an order of magnitude better than Millennium-I

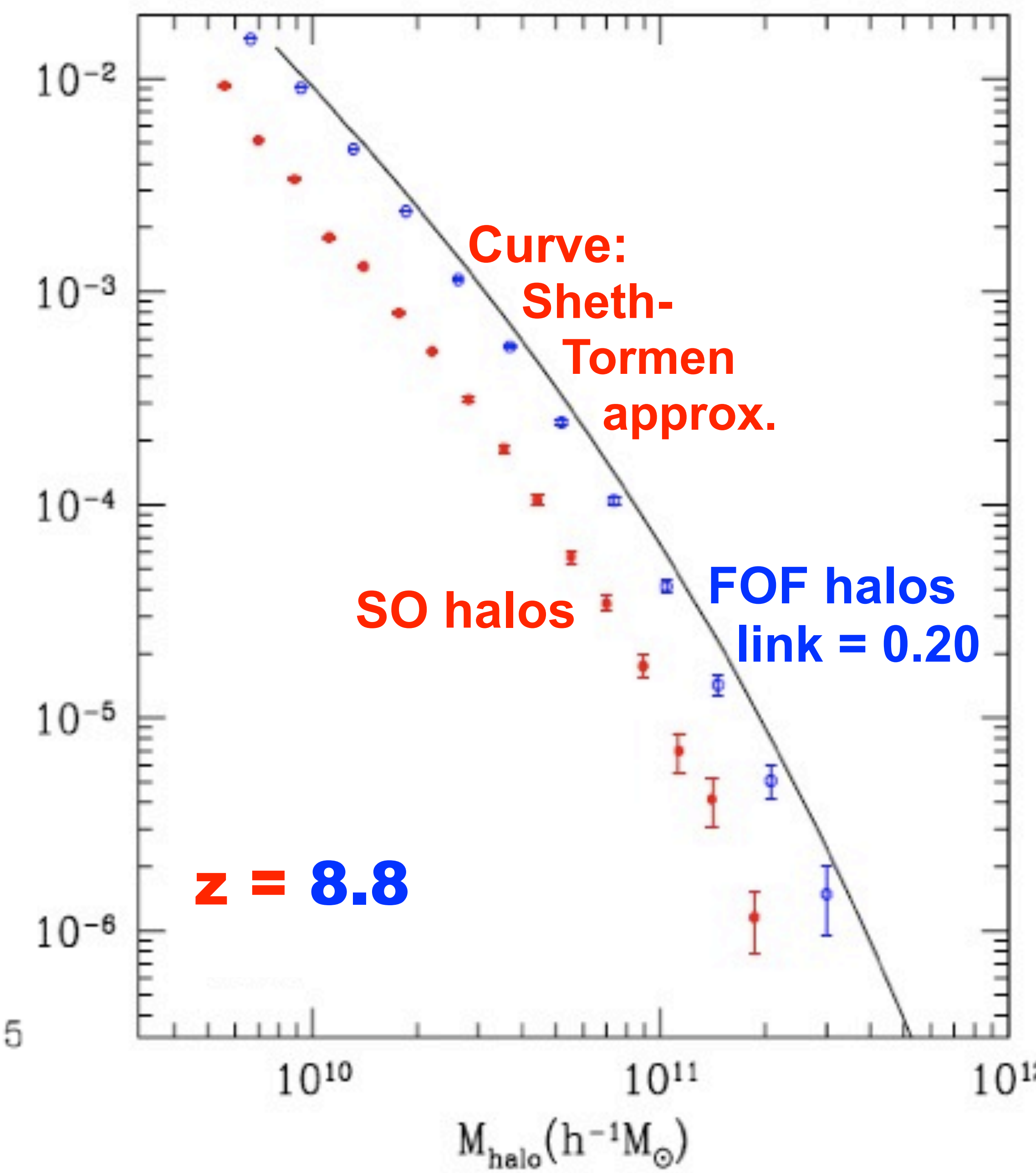
Force resolution is the same as Millennium-II, in a volume 16x larger

Halo finding is complete to V_{circ} > 50 km/s, using both BDM and ROCKSTAR halo finders

Bolshoi and MultiDark halo catalogs were released in September 2011 at Astro Inst Potsdam; Merger Trees will soon be available

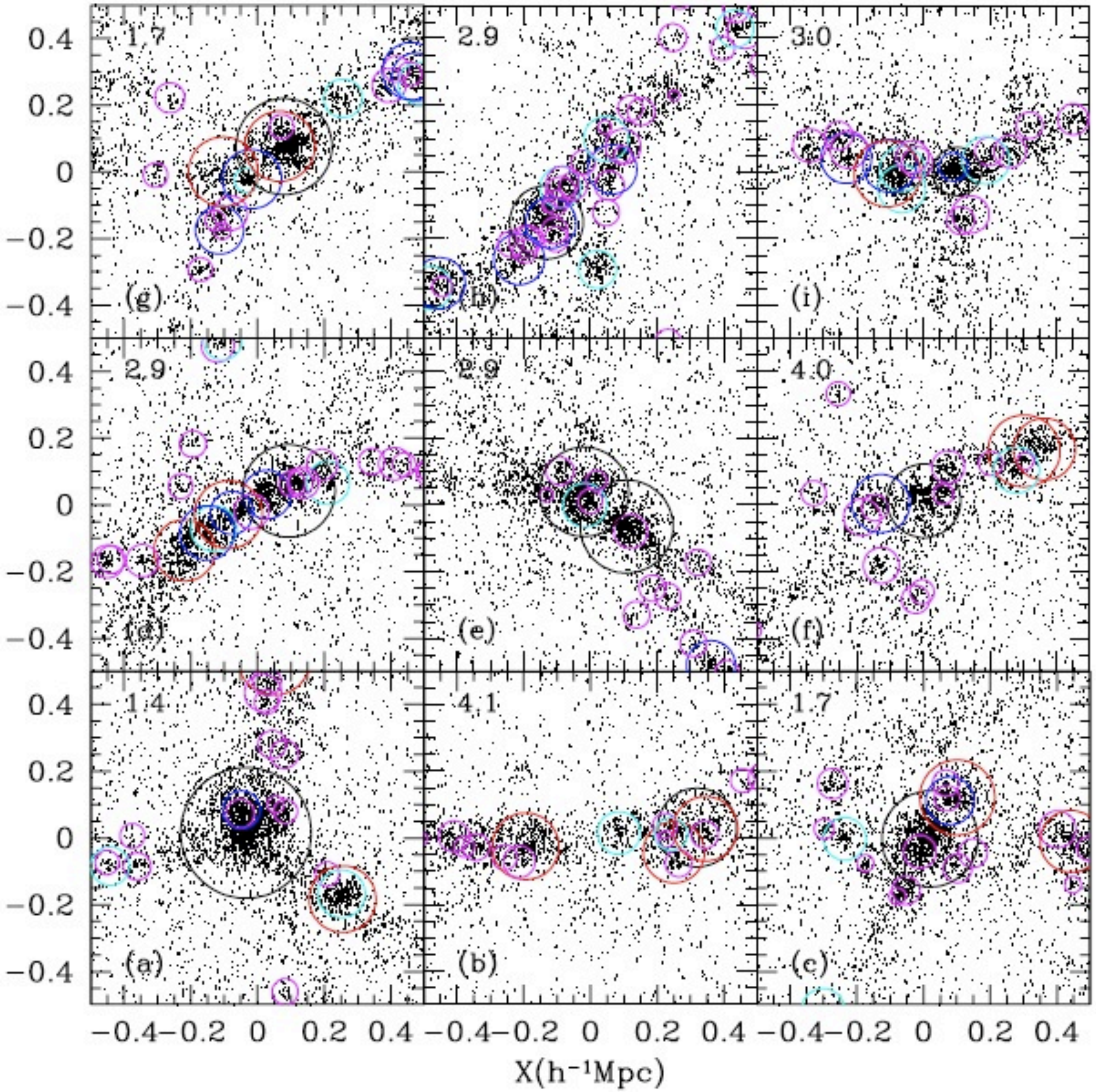


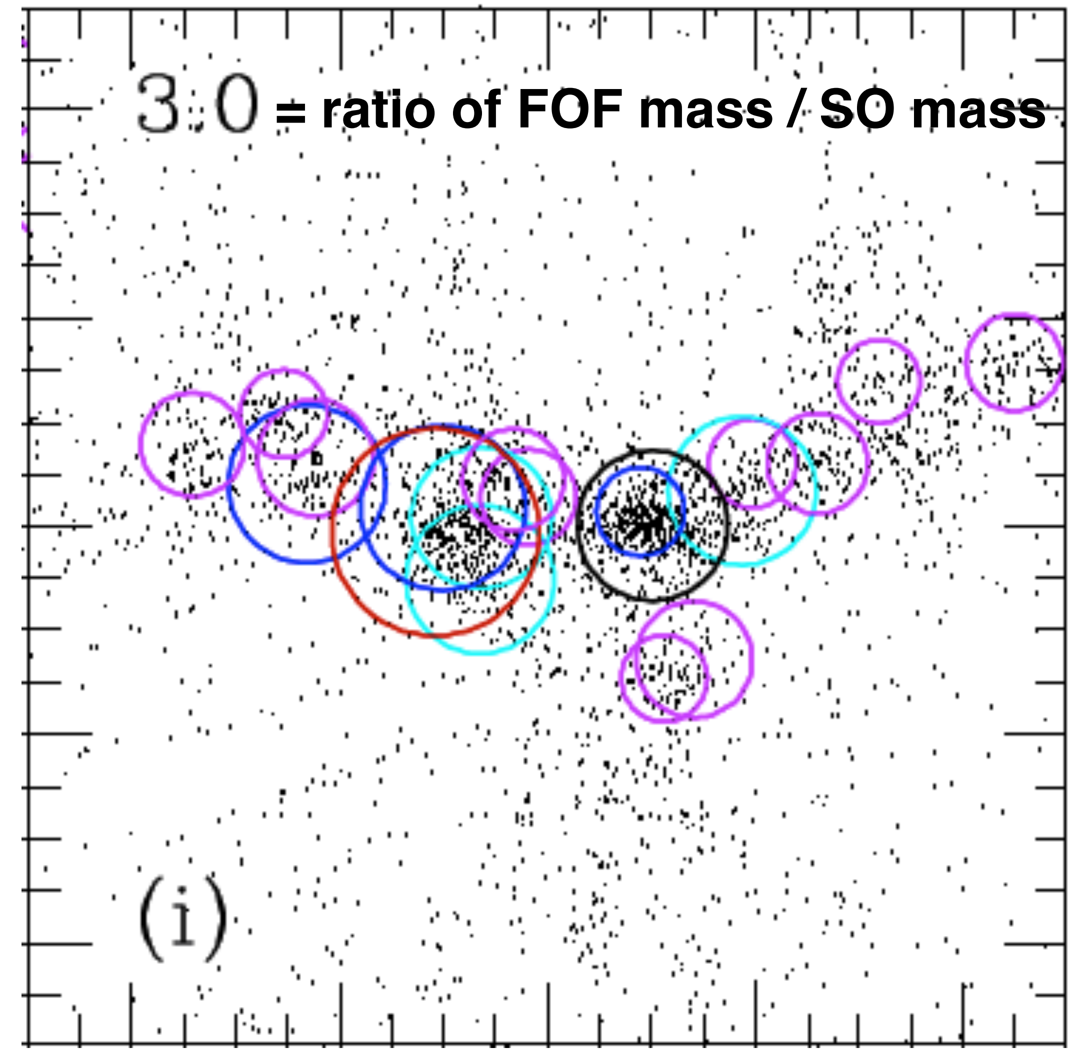
Sheth-Tormen Fails at High Redshifts



The Sheth-Tormen approximation with the same WMAP5 parameters used for the Bolshoi simulation very accurately agrees with abundance of halos at low redshifts, but increasingly overpredicts bound spherical overdensity halo abundance at higher redshifts. ST agrees well with FOF halo abundances, but FOF halos have unrealistically large masses at high *z*.

Klypin, Trujillo-Gomez, & Primack, 2011 ApJ

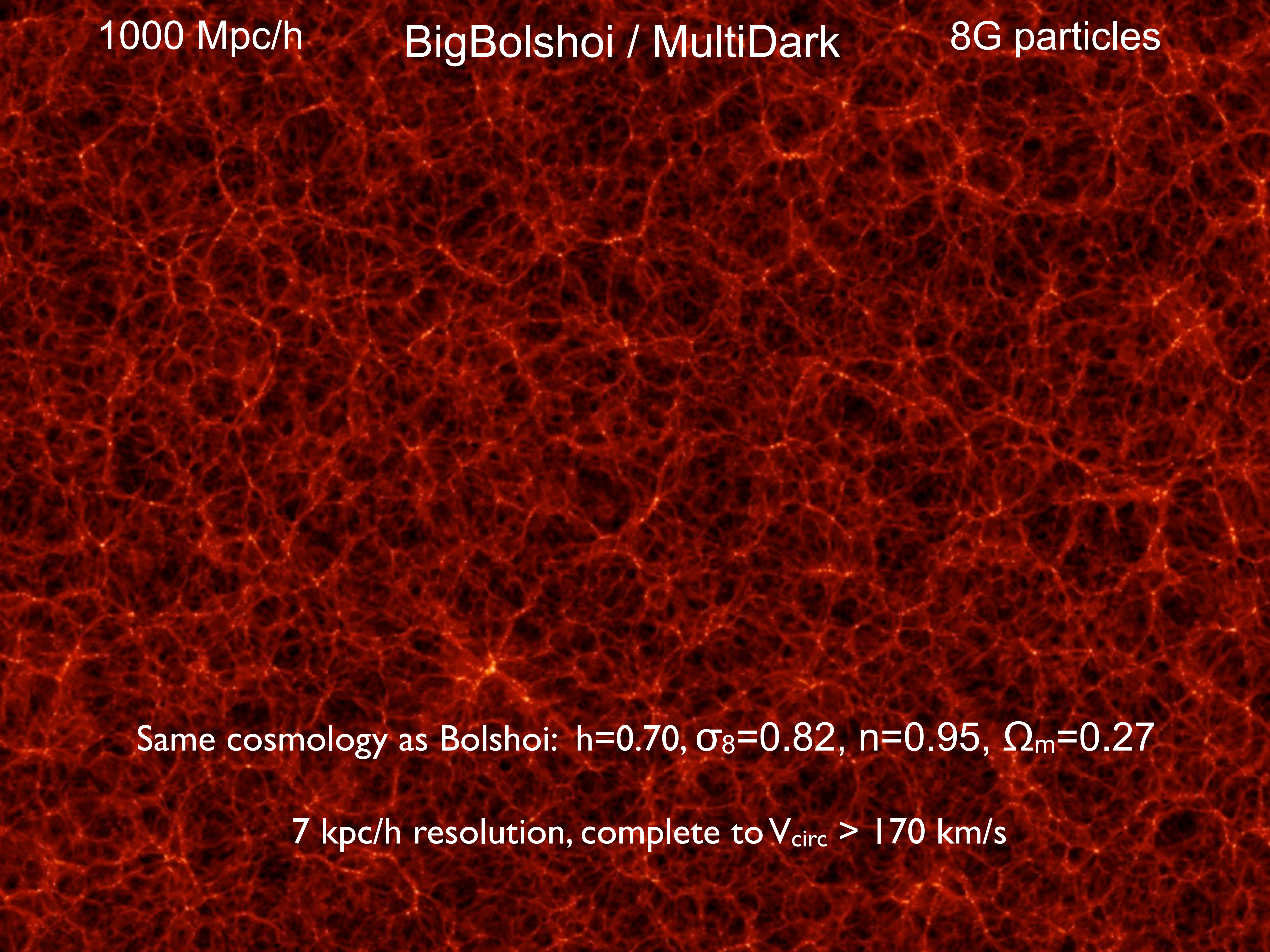




FOF linked together a chain of halos that formed in long and dense filaments (also in panels b, d, f, h; e = major merger)

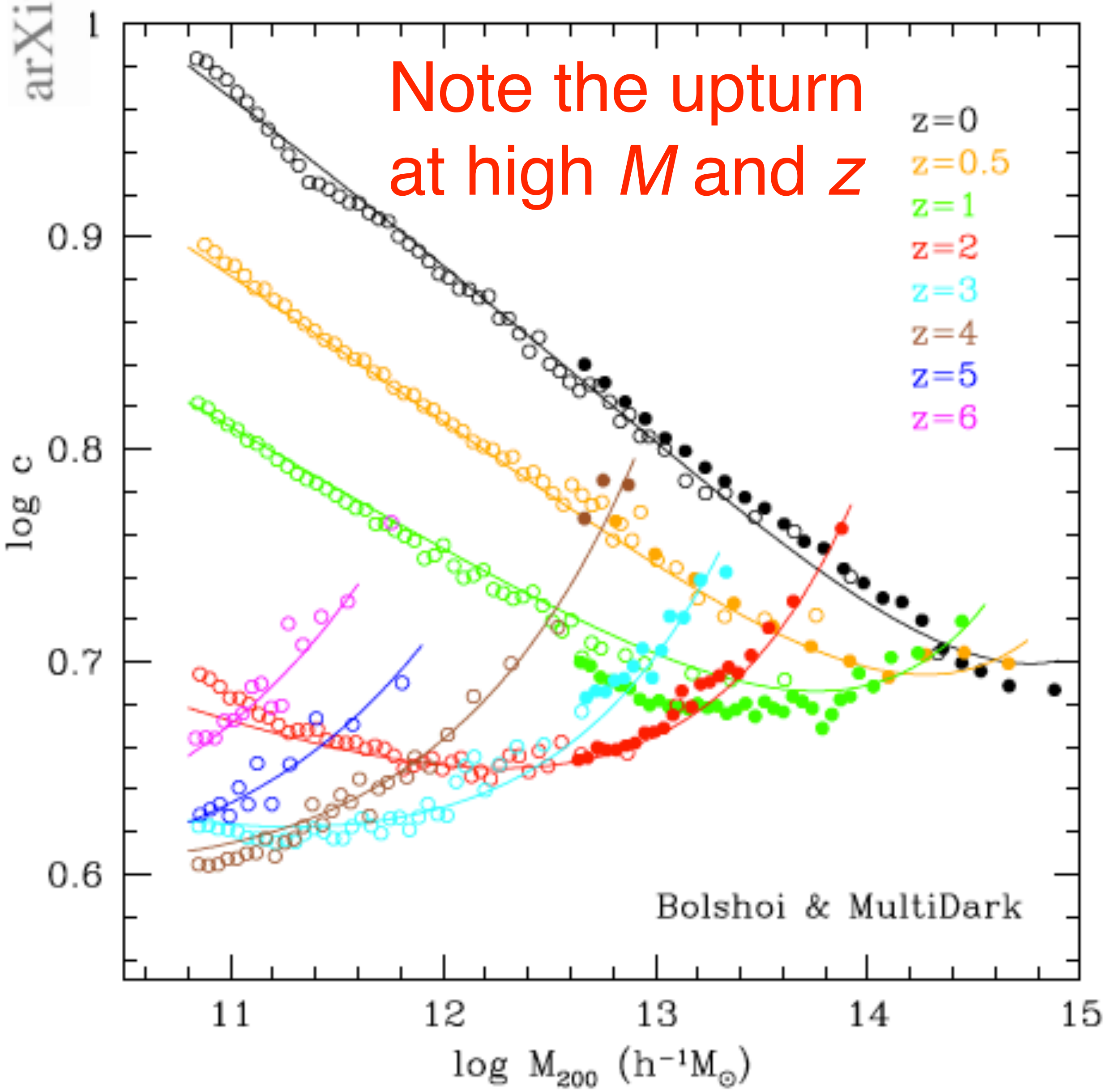
Each panel shows 1/2 of the dark matter particles in cubes of 1 h^{-1} Mpc size. The center of each cube is the exact position of the center of mass of the corresponding FOF halo. The effective radius of each FOF halo in the plots is 150 – 200 h^{-1} kpc. Circles indicate virial radii of distinct halos and subhalos identified by the spherical overdensity algorithm BDM.

Klypin, Trujillo-Gomez, & Primack, 2011 ApJ


1000 Mpc/h
BigBolshoi / MultiDark
8G particles
Same cosmology as Bolshoi: h=0.70, σ_8=0.82, n=0.95, Ω_m=0.27
7 kpc/h resolution, complete to V_{circ} > 170 km/s

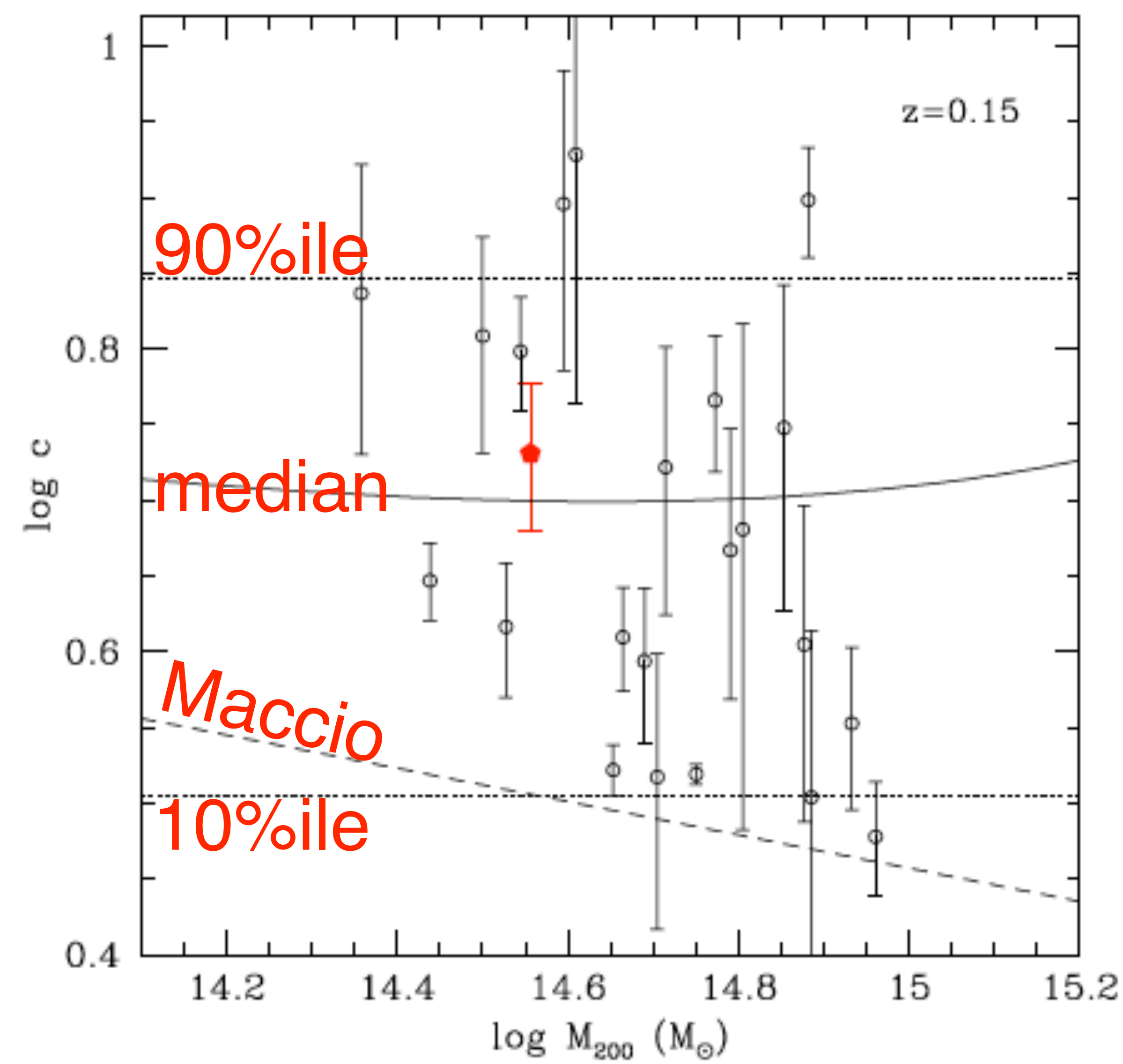
Halo concentrations in the standard CDM cosmology

Francisco Prada, Anatoly A. Klypin, Antonio J. Cuesta, Juan E. Betancort-Rijo, and Joel Primack



Halo mass–concentration relation of distinct halos at different redshifts in the Bolshoi (open symbols) and MultiDark (filled symbols) simulations is compared with an analytical approximation.

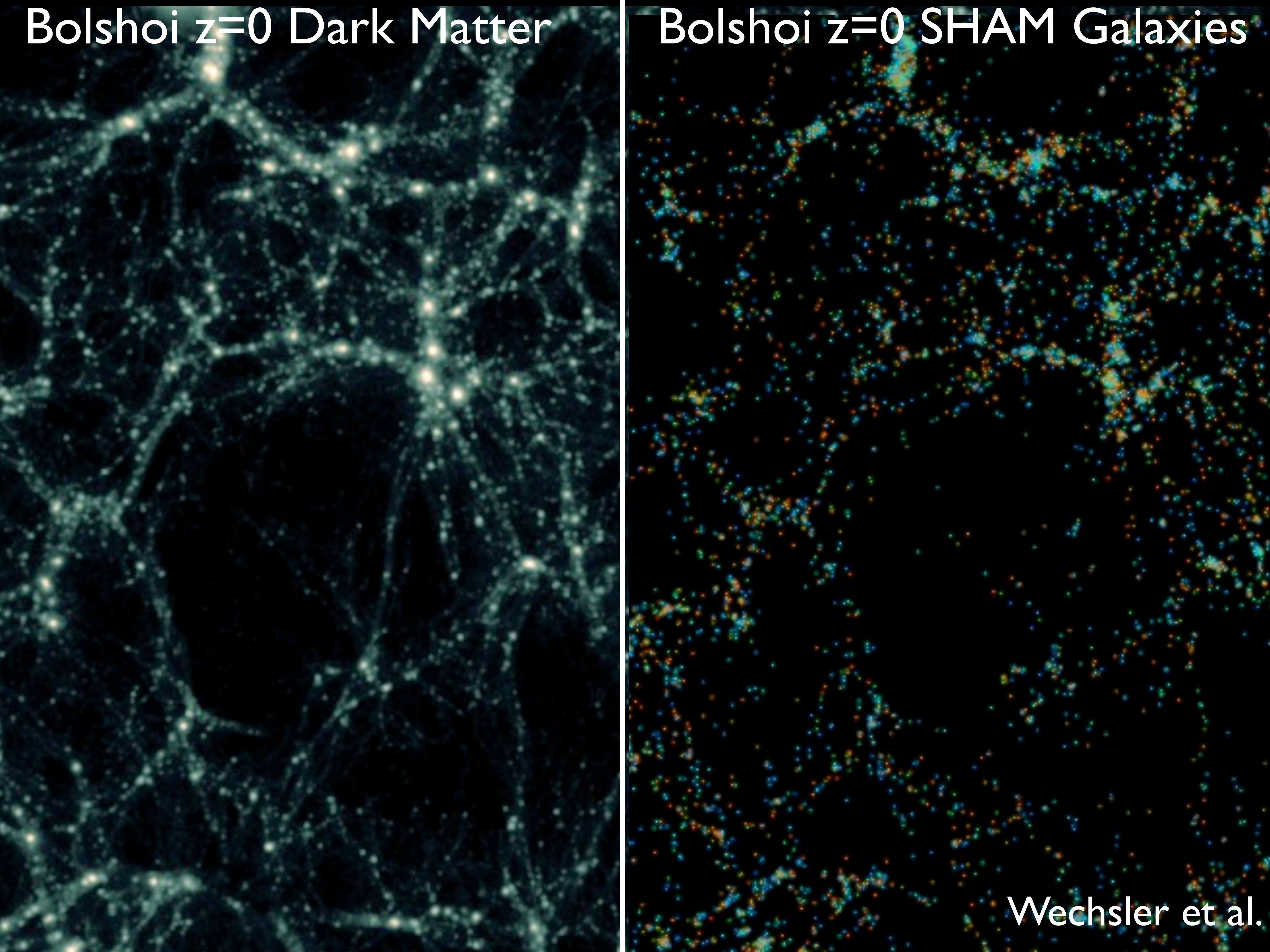
Cluster Concentrations



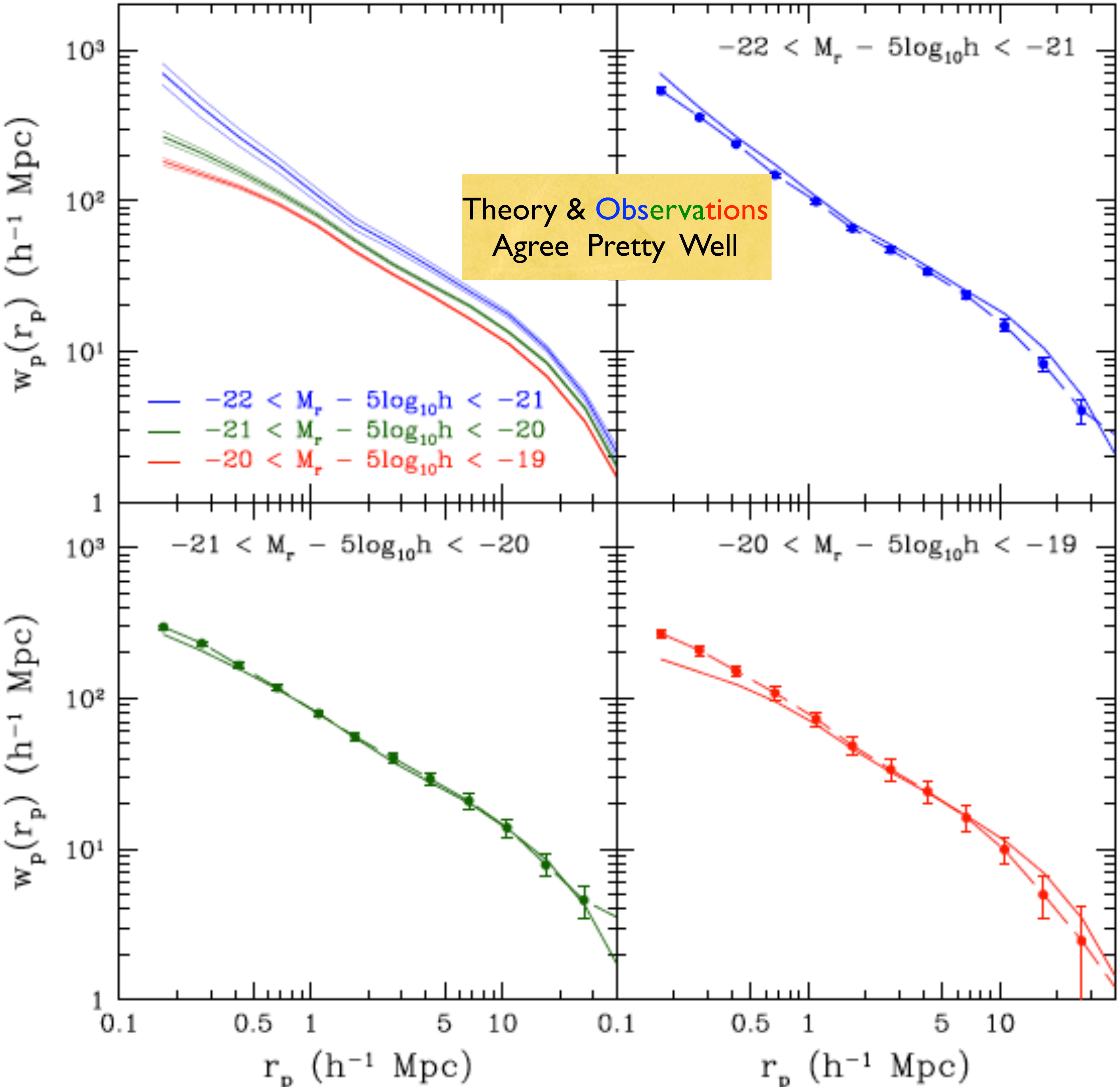
Comparison of observed cluster concentrations (data points with error bars) with the prediction of our model for median halo concentration of cluster-size halos (full curve). Dotted lines show 10% and 90% percentiles. Open circles show results for X-ray luminous galaxy clusters observed with XMMNewton in the redshift range 0.1-0.3 (Ettori et al. 2010). The pentagon presents galaxy kinematic estimate for relaxed clusters by Wojtak & Lokas (2010). The dashed curve shows prediction by Macci`o, Dutton, & van den Bosch (2008), which significantly underestimates the concentrations of clusters.

Merger History of a Big Halo

BOLSHOI
Merger Tree
Peter Behroozi, et al.


Bolshoi z=0 Dark Matter
Bolshoi z=0 SHAM Galaxies
Wechsler et al.

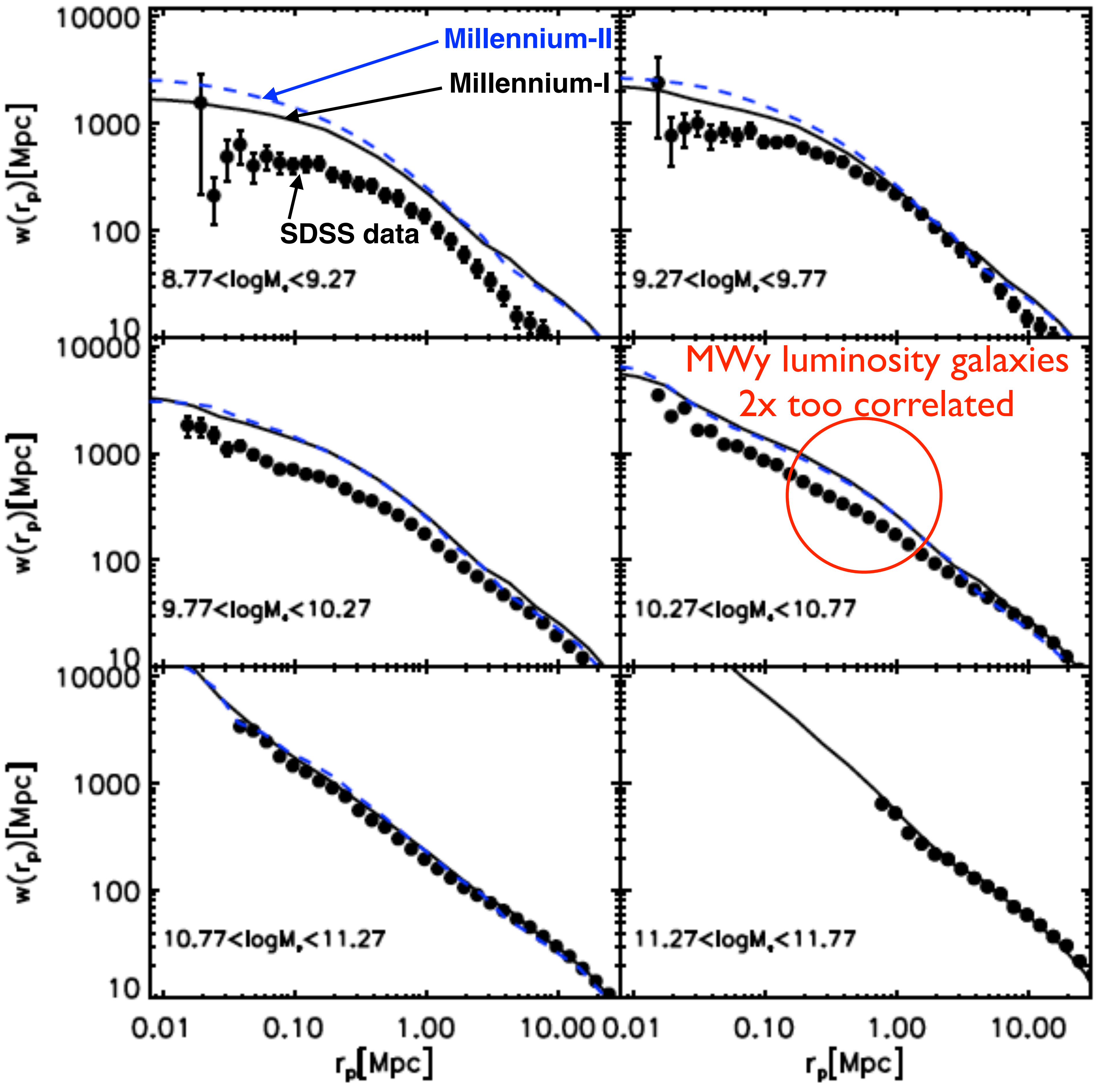
Bolshoi Projected Galaxy Correlation Functions



The correlation function of SDSS galaxies vs. Bolshoi galaxies using halo abundance matching, with scatter using our stochastic abundance matching method. This results in a better than 20% agreement with SDSS. *Top left:* correlation functinon in three magnitude bins, showing Poisson uncertainties as thin lines. *Remaining panels:* correlation function in each luminosity bin compared with SDSS galaxies (points with error bars: Zehavi et al. 2010).

Trujillo-Gomez, Klypin, Primack, & Romanowsky 2011 ApJ

Millennium Projected Galaxy Correlation Functions



Projected correlation functions for galaxies in different stellar mass ranges, in SAM based on Millennium I and II. Black solid and blue dashed curves give results for preferred model applied to the MS and the MS-II, respectively. Symbols with error bars are results for SDSS/DR7 calculated using the same techniques as in Li et al. (2006). The two simulations give convergent results for $M_* > 6\times10^9\ M_{sun}$. At lower mass the MS underestimates the correlations on small scales. The model agrees quite well with the SDSS at all separations for $M_* > 6\times10^{10}\ M_{sun}$. But **at smaller masses the correlations are overestimated substantially, particularly at small separations. The authors attribute this to the too-high $\sigma_8 = 0.90$ used in MS-I & II.**

Guo, White, et al. 2011 MNRAS

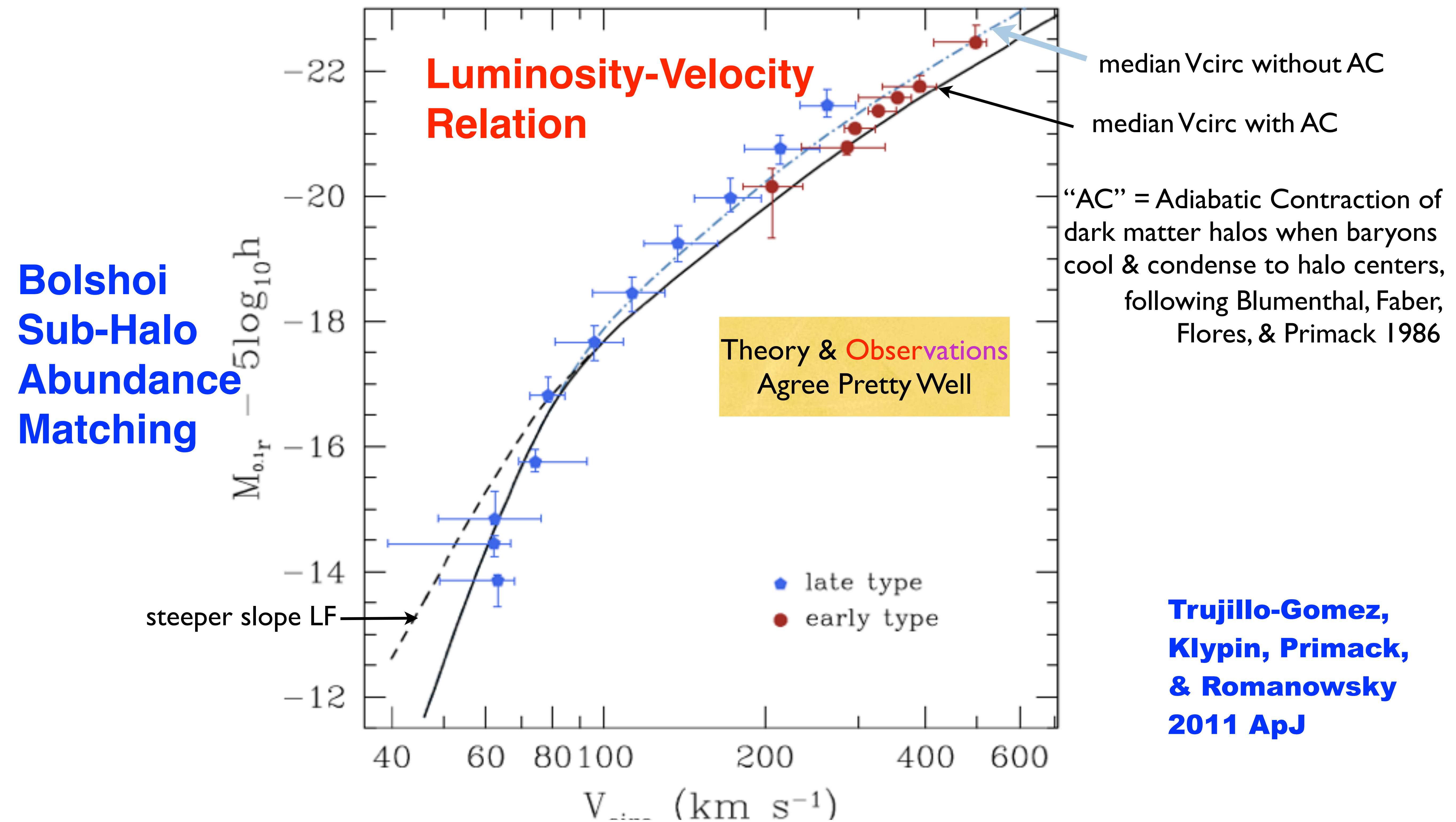
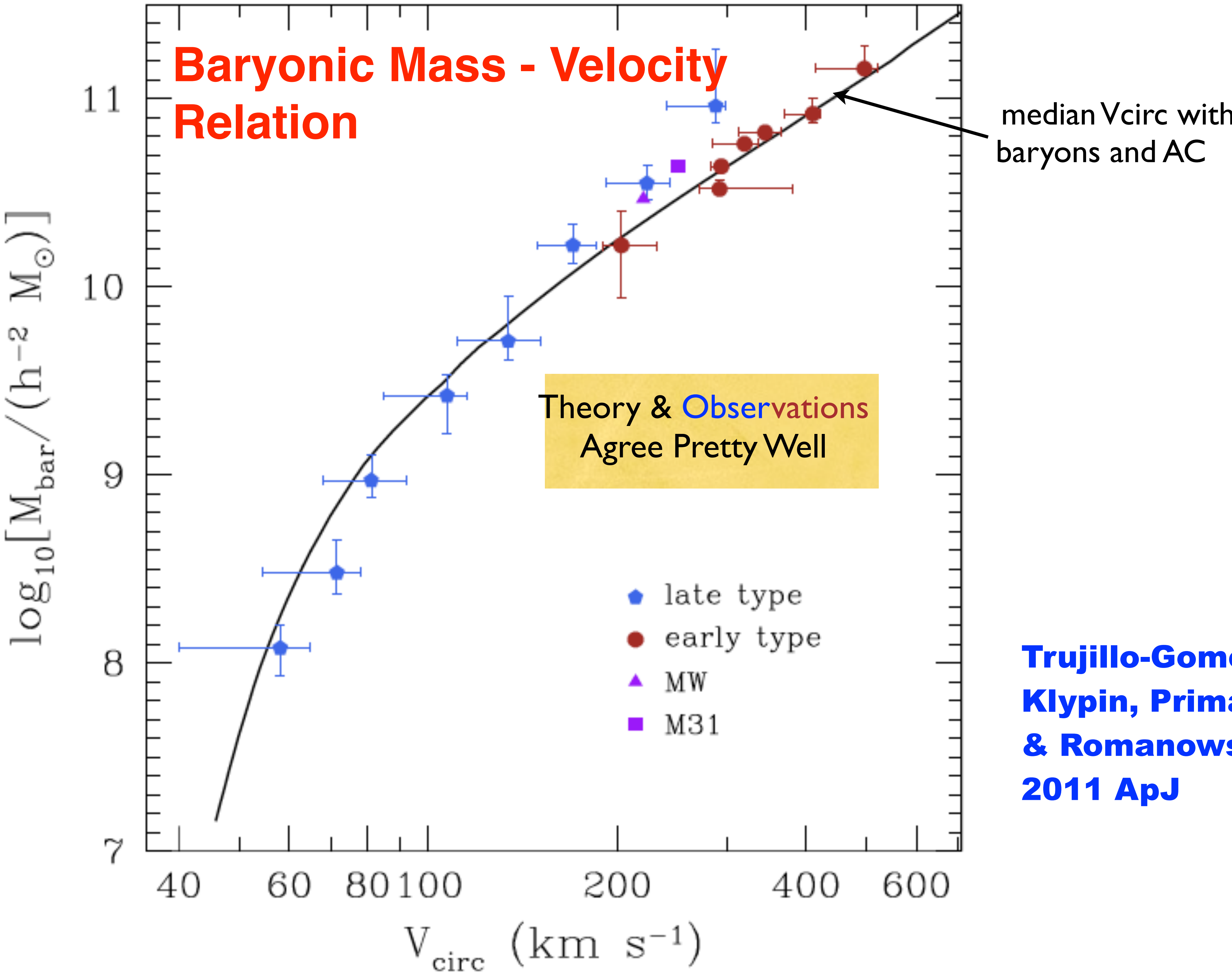


Fig. 4.— Comparison of the observed Luminosity Velocity relation with the predictions of the ΛCDM model. The solid curve shows the median values of $^{0.1}r$-band luminosity vs. circular velocity for the model galaxy sample. The circular velocity for each model galaxy is based on the peak circular velocity of its host halo over its entire history, measured at a distance of 10 kpc from the center including the cold baryonic mass and the standard correction due to adiabatic halo contraction. The dashed curve show results for a steeper ($\alpha = -1.34$) slope of the LF. The dot-dashed curve shows predictions after adding the baryon mass but without adiabatic contraction. Points show representative observational samples.

Bolshoi Sub-Halo Abundance Matching

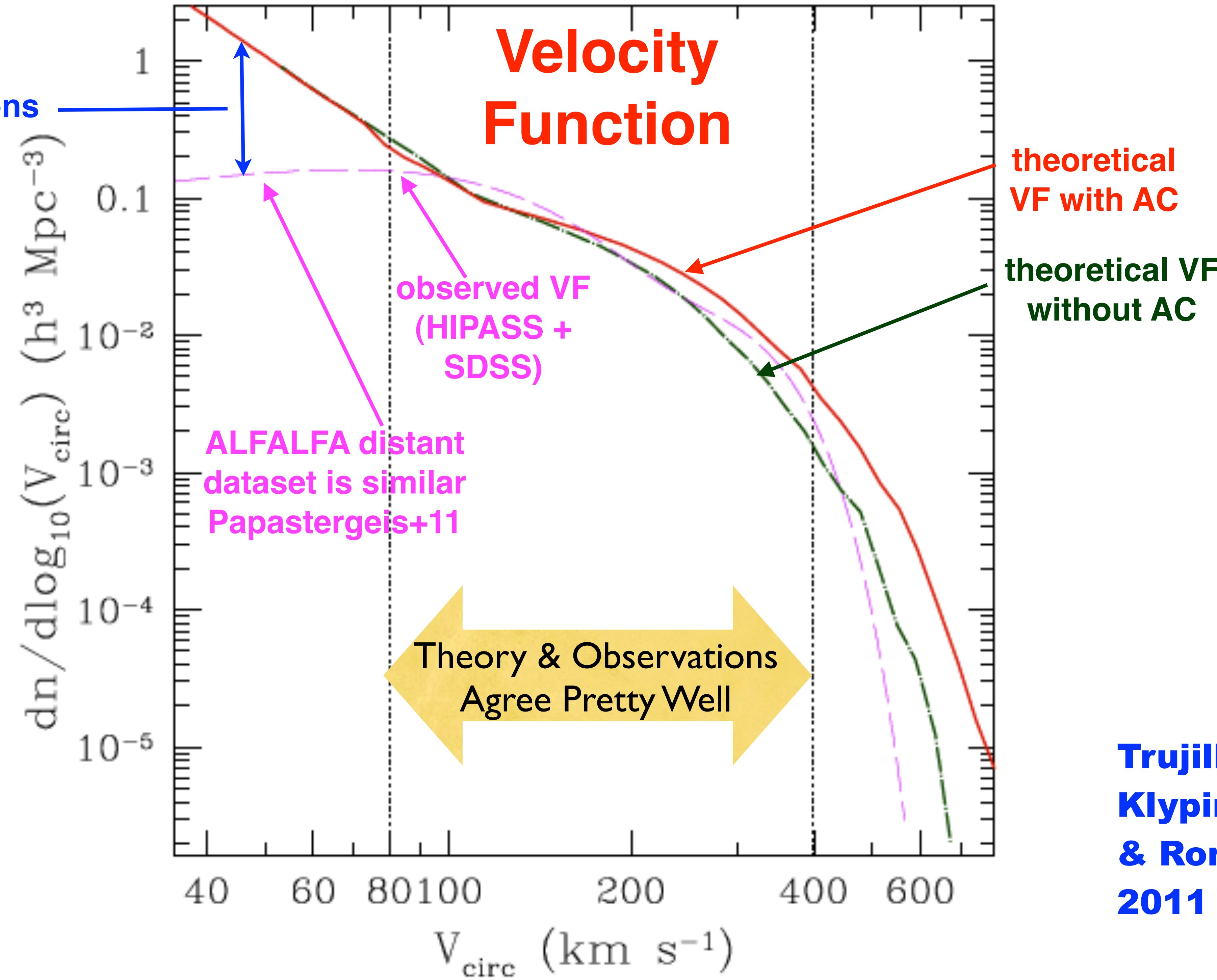


Trujillo-Gomez, Klypin, Primack, & Romanowsky 2011 ApJ

Fig. 10.— Mass in cold baryons as a function of circular velocity. The solid curve shows the median values for the ΛCDM model using halo abundance matching. The cold baryonic mass includes stars and cold gas and the circular velocity is measured at 10 kpc from the center while including the effect of adiabatic contraction. For comparison we show the individual galaxies of several galaxy samples. Intermediate mass galaxies such as the Milky Way and M31 lie very close to our model results.

Bolshoi Sub-Halo Abundance Matching

Discrepancy due to incomplete observations or ΛCDM failure?



Trujillo-Gomez, Klypin, Primack, & Romanowsky 2011 ApJ

Fig. 11.— Comparison of theoretical (dot-dashed and thick solid curves) and observational (dashed curve) circular velocity functions. The dot-dashed line shows the effect of adding the baryons (stellar and cold gas components) to the central region of each DM halo and measuring the circular velocity at 10 kpc. The thick solid line is the distribution obtained when the adiabatic contraction of the DM halos is considered. Because of uncertainties in the AC models, realistic theoretical predictions should lie between the dot-dashed and solid curves. Both the theory and observations are highly uncertain for rare galaxies with $V_{\rm circ} > 400$ km s^{-1}. Two vertical dotted lines divide the VF into three domains: $V_{\rm circ} > 400$ km s^{-1} with large observational and theoretical uncertainties; < 80 km s^{-1} $< V_{\rm circ} <$ 400 km s^{-1} with a reasonable agreement, and $V_{\rm circ} < 80$ km s^{-1}, where the theory significantly overpredicts the number of dwarfs.

Presented at KITP Conference "First Light and Faintest Dwarfs" Feb 2012

Klypin, Karachentsev, Nasonova 2012

Deeper Local Survey -- better agreement with ΛCDM but still more halos than galaxies below 50 km/s

Local Volume: D <10Mpc

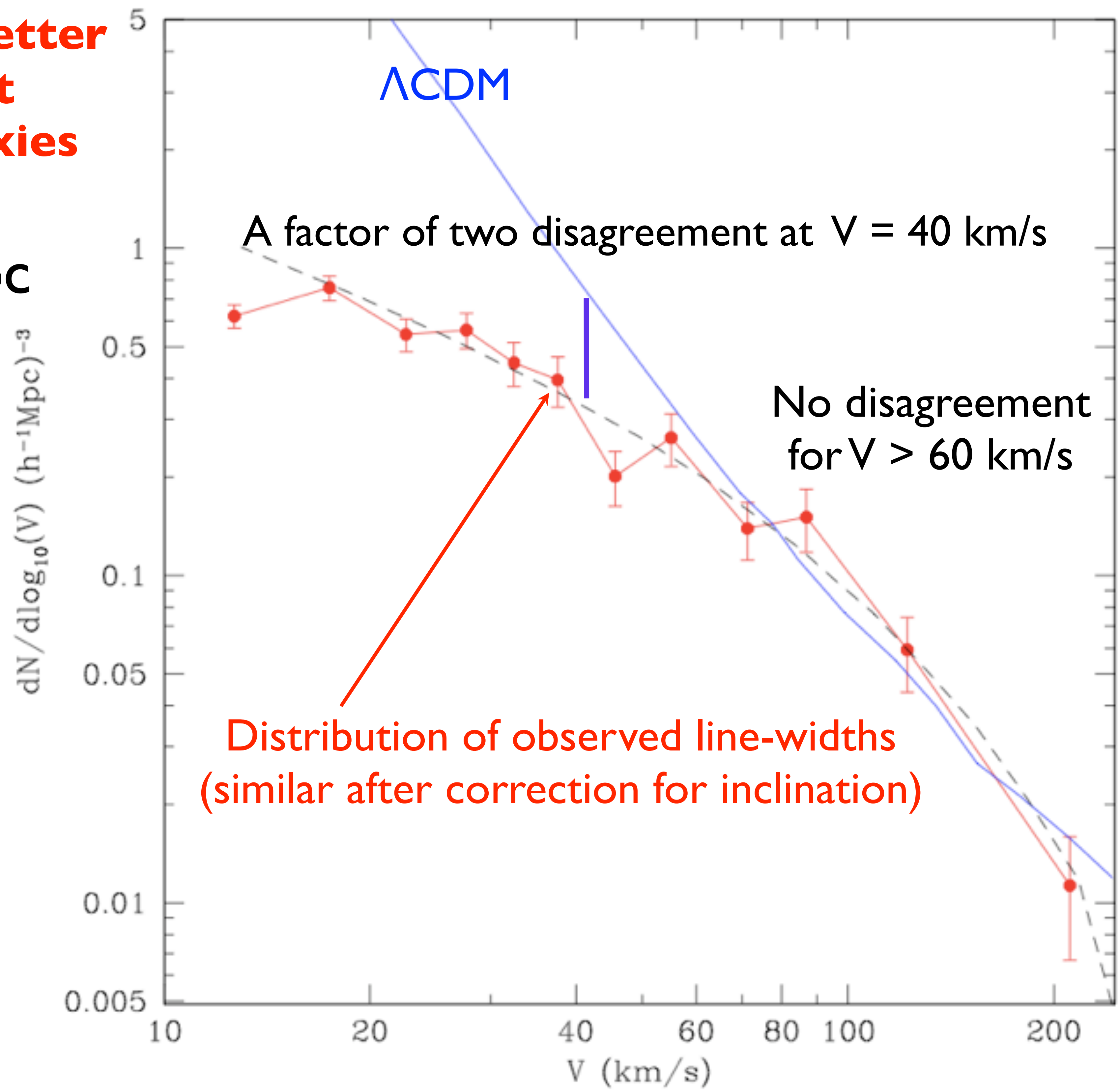
Total sample: 813 galaxies
Within 10Mpc: 686
M_B<-13 N=304
M_B<-10 N=611

80-90% are spirals or dIrr (T>0)

Accuracy of distances are 8-10%

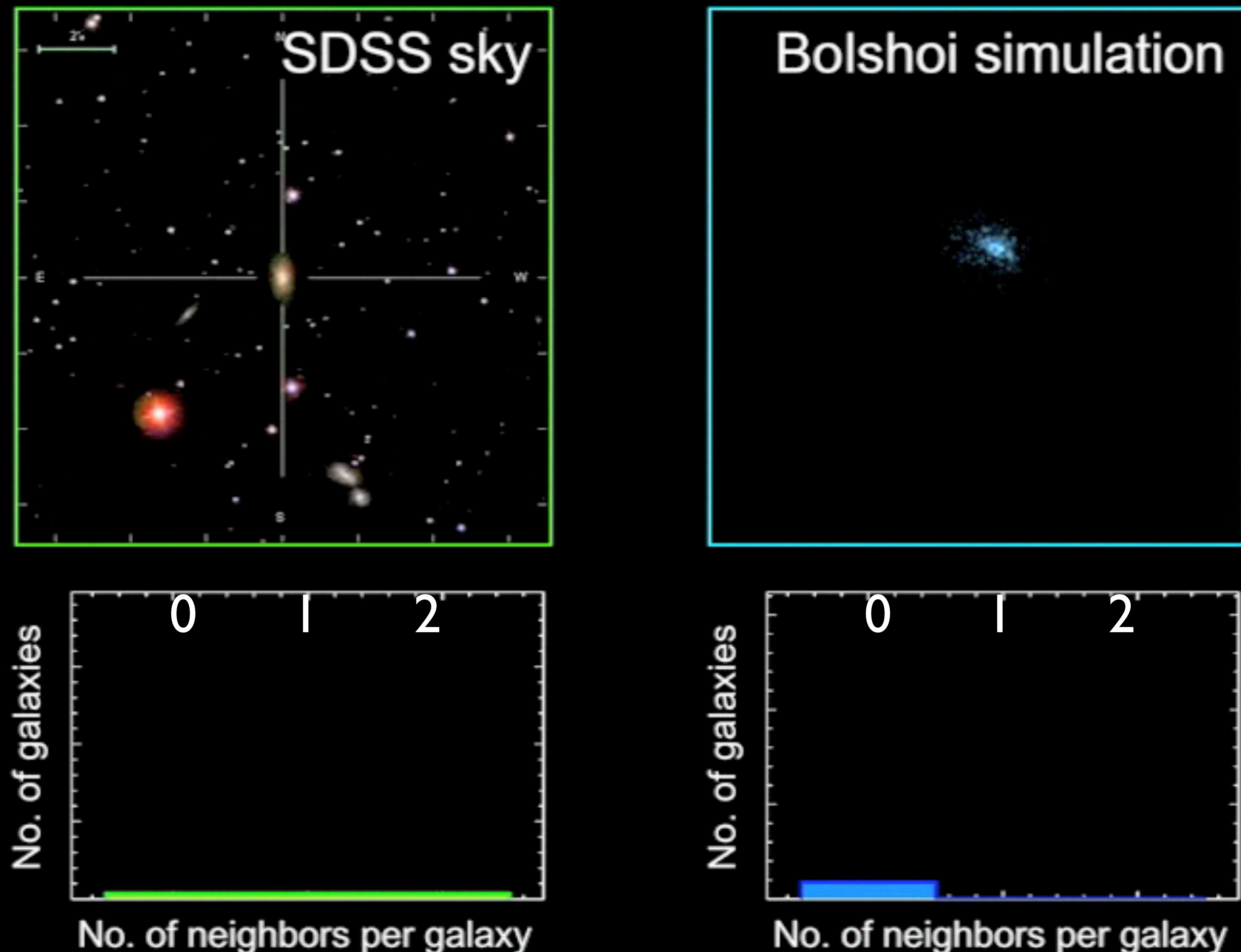
80% with D<10Mpc have HI linewidth

Vrot = 150x10^(-(20.5+M_B)/8.5)km/s



The Milky Way has two large satellite galaxies, the small and large Magellanic Clouds

The Bolshoi simulation + halo abundance matching predicts the likelihood of this

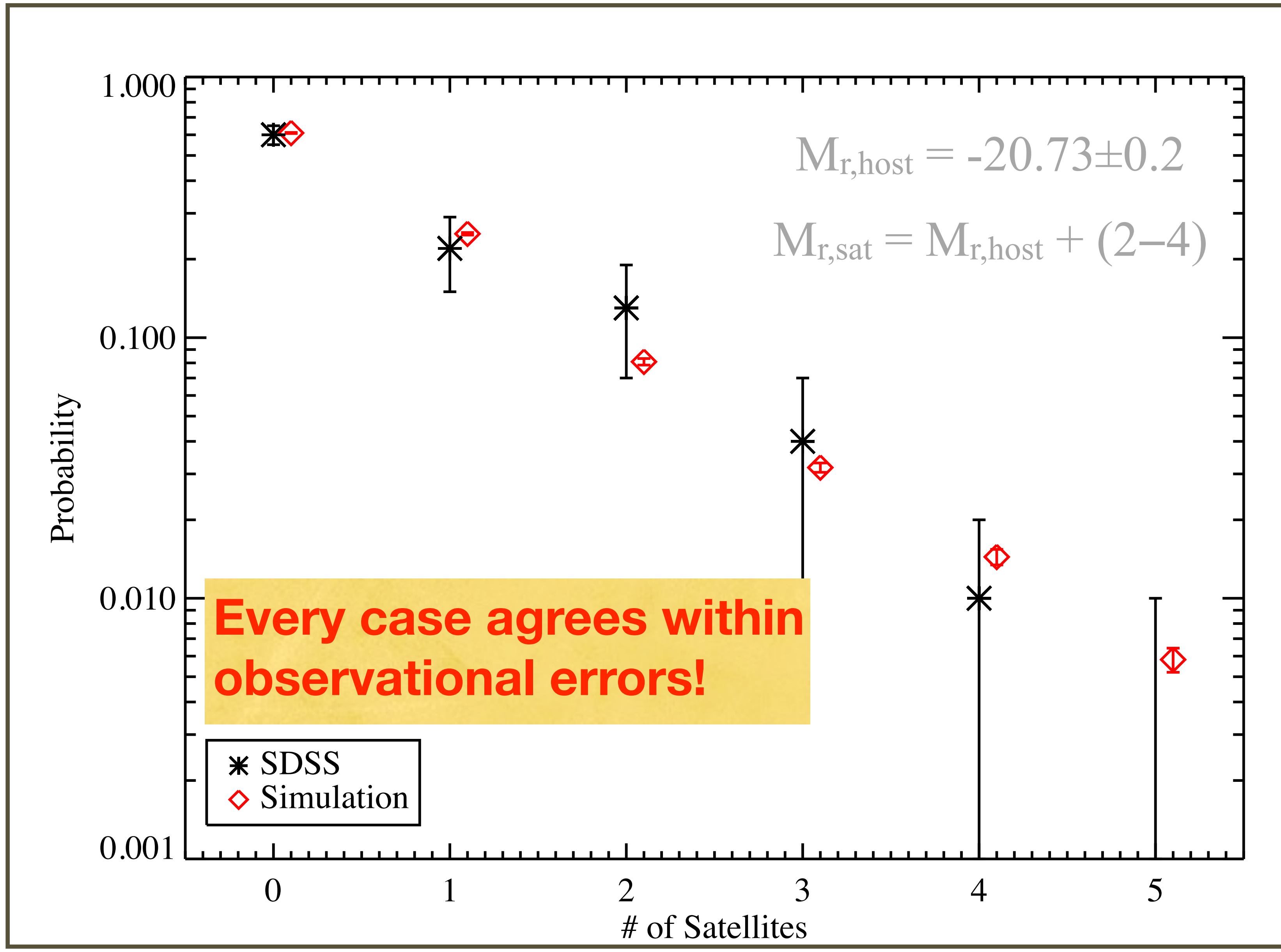


SDSS sky
N
E
W
S
Bolshoi simulation
0
1
2
No. of galaxies
No. of neighbors per galaxy
0
1
2
No. of galaxies
No. of neighbors per galaxy

- Apply the same absolute magnitude and isolation cuts to Bolshoi+SHAM galaxies as to SDSS:
  - Identify all objects with absolute $^{0.1}M_r$ = -20.73±0.2 and observed m_r < 17.6
  - Probe out to z = 0.15, a volume of roughly 500 (Mpc/h)^3
  - leaves us with 3,200 objects.
- Comparison of Bolshoi with SDSS observations is in close agreement, well within observed statistical error bars.

# of Subs	Prob (obs)	Prob (sim)
0	60%	61%
1	22%	25%
2	13%	8.1%
3	4%	3.2%
4	1%	1.4%
5	0%	0.58%

Statistics of MW bright satellites: SDSS data vs. Bolshoi simulation



Busha et al. 2011 ApJ
Liu et al. 2011 ApJ

Risa Wechsler

Similarly good agreement with SDSS for brighter satellites with spectroscopic redshifts compared with Millennium-II using abundance matching -- Tollerud, Boylan-Kolchin, et al. 2011 ApJ

SMALL-SCALE STRUCTURE IN THE SDSS AND ΛCDM: ISOLATED $\sim L_*$ GALAXIES WITH BRIGHT SATELLITES

ERIK J. TOLLERUD[1], MICHAEL BOYLAN-KOLCHIN[1,2], ELIZABETH J. BARTON[1], JAMES S. BULLOCK [1], CHRISTOPHER Q. TRINH[3,1]

ApJ 2011

Similarly good agreement with SDSS for brighter satellites with spectroscopic redshifts compared with Millennium-II using abundance matching.

We use a volume-limited spectroscopic sample of isolated galaxies in the Sloan Digital Sky Survey (SDSS) to investigate the frequency and radial distribution of luminous ($M_r \lesssim -18.3$) satellites like the Large Magellanic Cloud (LMC) around $\sim L_*$ Milky Way analogs and compare our results object-by-object to ΛCDM predictions based on abundance matching in simulations. We show that 12% of Milky Way-like galaxies host an LMC-like satellite within 75 kpc (projected), and 42% within 250 kpc (projected). This implies $\sim$ 10% have a satellite within the distance of the LMC, and $\sim$ 40% of L_* galaxies host a bright satellite within the virialized extent of their dark matter halos. Remarkably, the simulation reproduces the observed frequency, radial dependence, velocity distribution, and luminosity function of observed secondaries exceptionally well, suggesting that ΛCDM provides an accurate reproduction of the observed Universe to galaxies as faint as $L \sim 10^9 L_\odot$ on $\sim$ 50 kpc scales. When stacked, the observed projected pairwise velocity dispersion of these satellites is $\sigma \simeq$ 160 km s^{-1}, in agreement with abundance-matching expectations for their host halo masses. Finally, bright satellites around L_* primaries are significantly *redder* than typical galaxies in their luminosity range, indicating that environmental quenching is operating within galaxy-size dark matter halos that typically contain only a single bright satellite. This redness trend is in stark contrast to the Milky Way's LMC, which is unusually blue even for a field galaxy. We suggest that the LMC's discrepant color might be further evidence that it is undergoing a triggered star-formation event upon first infall.

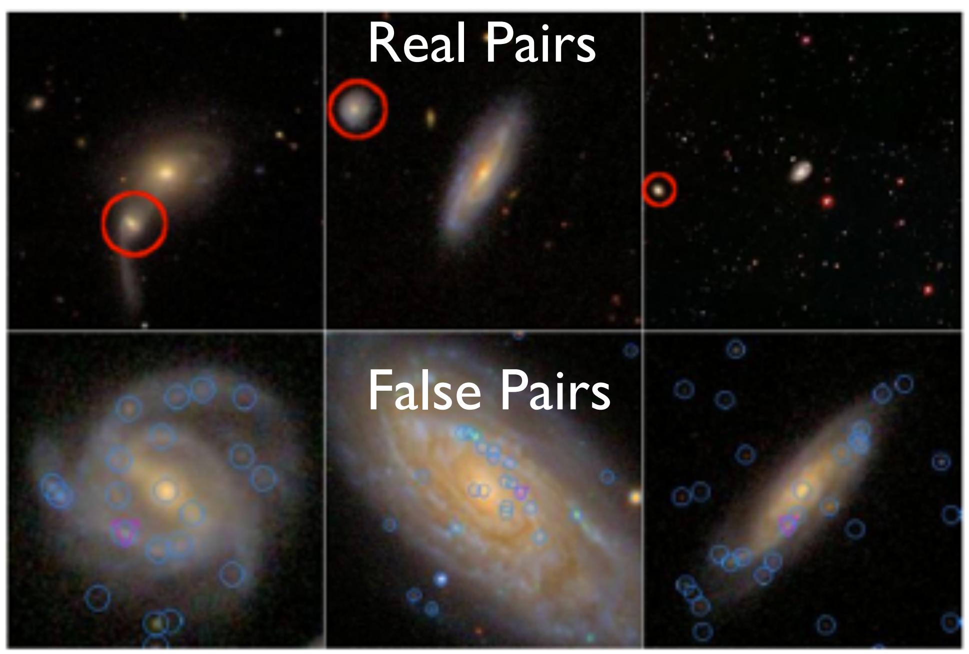


FIG. 1.— Examples of SDSS primary/secondary pairs in the clean sample (upper) and false pairs (lower). Secondaries identified by our criteria (see text) are marked with red circles (upper panels) or magenta triangles (lower panels). The upper three are all in the clean sample (have redshifts close to the primary) and span a range of projected separations. For the lower three images, blue circles are SDSS pipeline photometric objects, clearly showing the identification of HII regions as photometric objects. For these same lower three, the secondaries are clearly HII regions in the primary (or satellites that are indistinguishable from HII regions). We visually identify and remove all pairs of this kind from our sample.

Good agreement between simulated and observed pairwise velocities

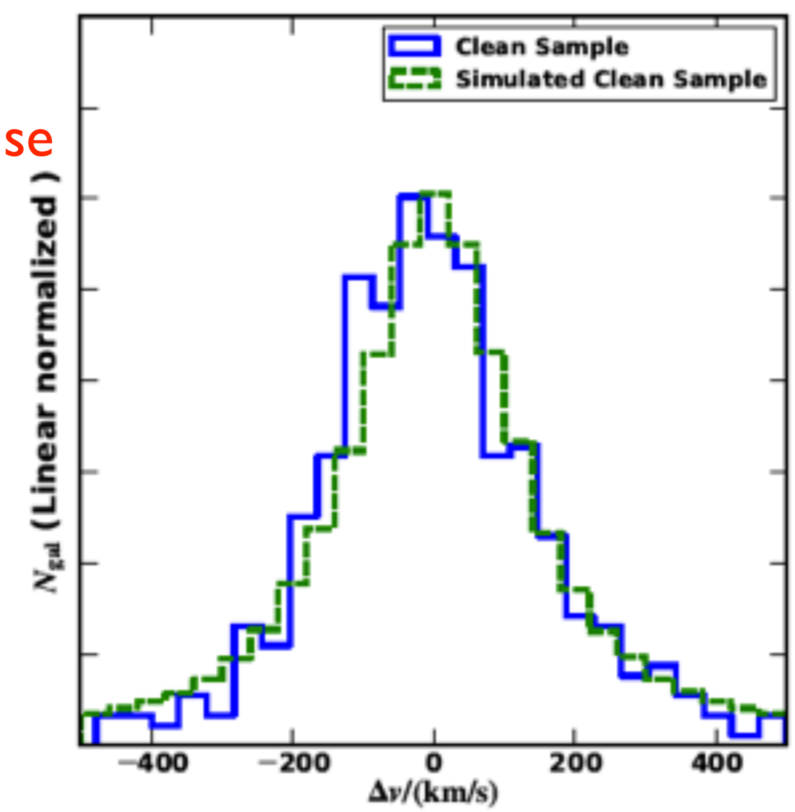


FIG. 6.— Distribution of $\Delta v \equiv c(z_{pri} - z_{sec})$ for the clean sample (solid blue histogram), the clean-like sample from MS-II (dashed green). The KS test yields $p_{KS} = 33\%$. The pairwise velocity dispersion in the observed sample is $\sigma = 161$ km s^{-1}.

ΛCDM vs. Downsizing

ΛCDM:

hierarchical formation
(small things form first)

"Downsizing":

massive galaxies are old, star formation moves to smaller galaxies

small structures	early	large galaxies
↓		↓
large structures	late	small galaxies

ΛCDM vs. Downsizing

ΛCDM:

hierarchical formation
(small things form first)

=

mass assembly

simulations (DM)

↓

present-day structure

"Downsizing":

massive galaxies are old, star formation moves to smaller galaxies

=

star formation history

semi-analytic models

↓

current stellar population

How are these processes related?

Formation of galaxies and large-scale structure with cold dark matter

Blumenthal, Feber, Primack, & Rees -- Nature 311, 517 (1984)

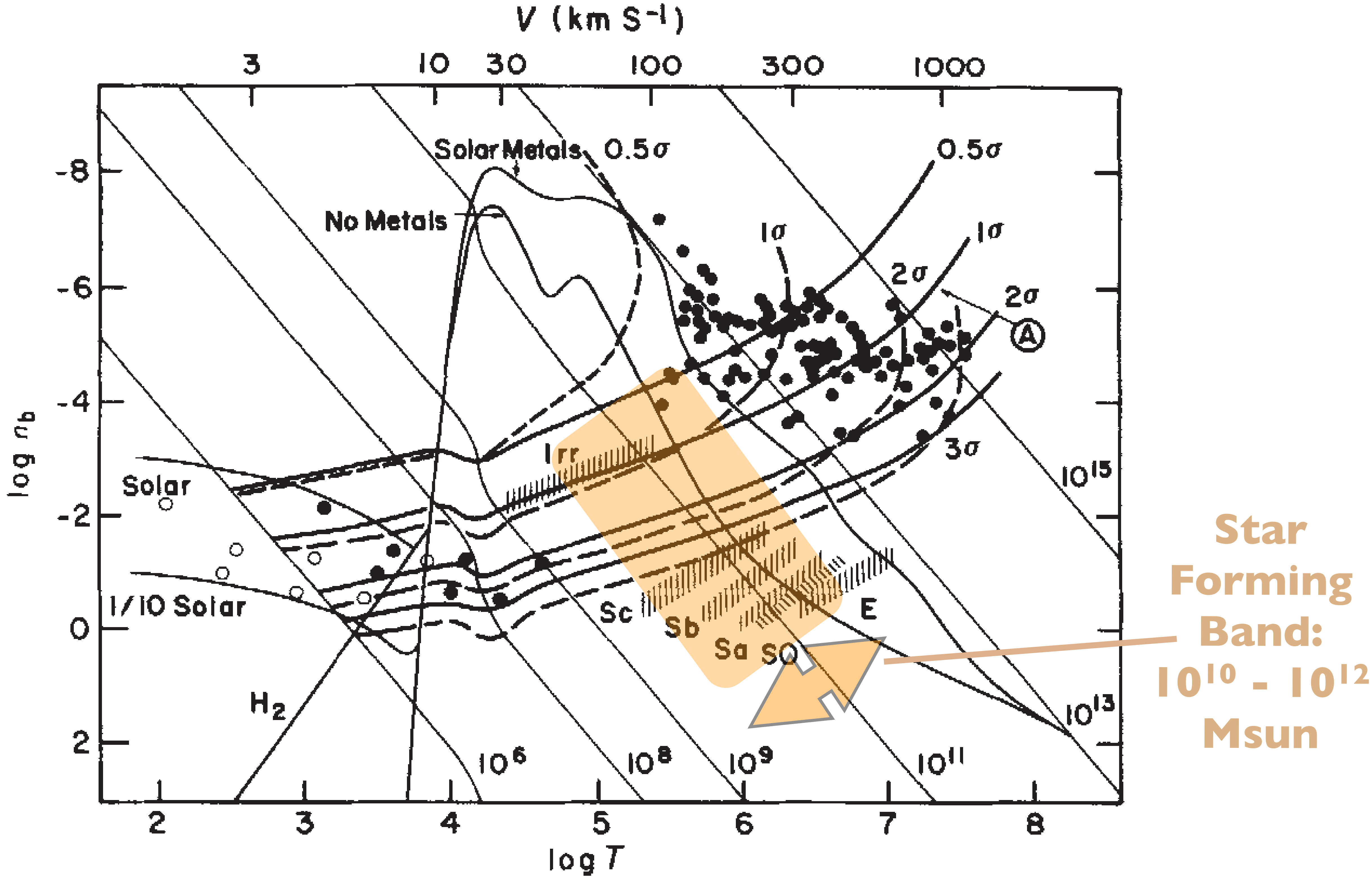
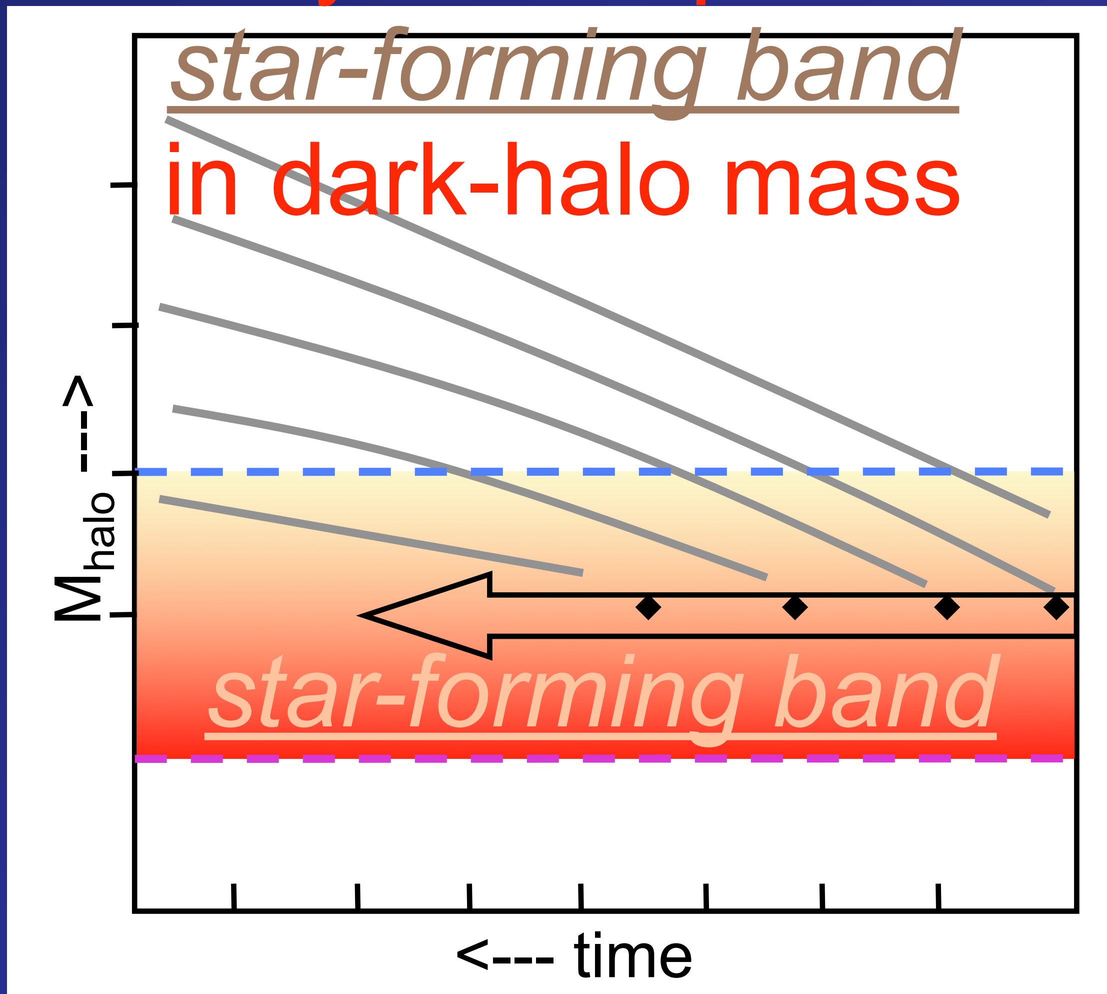


Fig. 3 Baryon density n_b versus three-dimensional, r.m.s. velocity dispersion V and virial temperature T for structures of various size in the Universe. The quantity T is $\mu V^2/3k$, where μ is mean molecular weight ($\approx$0.6 for ionized, primordial H+He) and k is Boltzmann's constant.

Key assumption:



Implications and Predictions of the Model

1) Each halo has a unique dark-matter growth path and associated stellar mass growth path.

2) Stellar mass follows halo mass until M_{halo} crosses M_{crit}.

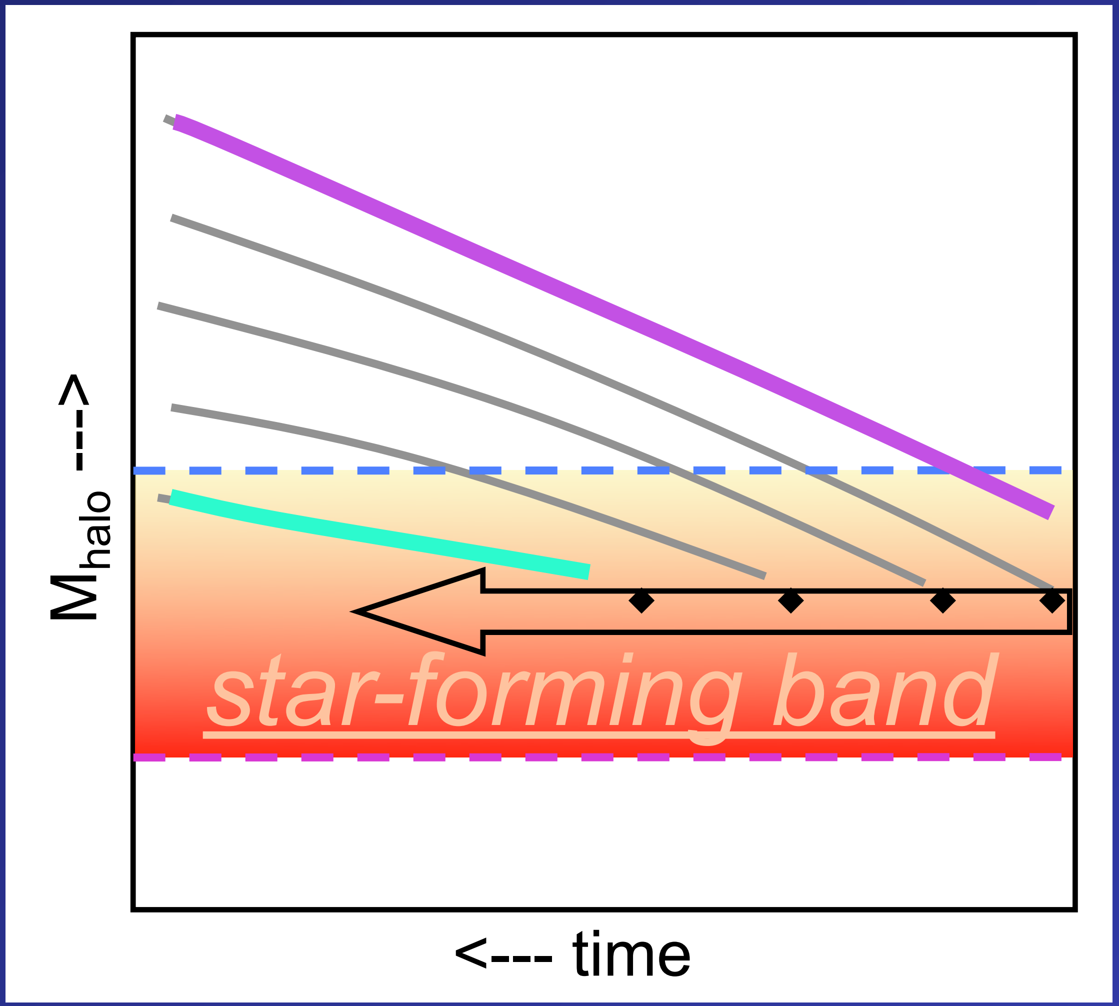
SAMs: $M_{star} < 0.05\ M_{halo}$

3) A *mass sequence* comes from the fact that different halo masses enter the star-forming band at different times. A galaxy's position is determined by its *entry redshift* into the band. More massive galaxies enter earlier. Thus:

z_{entry} <--> M_{halo} <--> M_{star}

Sandy Faber

Implications and Predictions of the Model



Massive galaxies:

- Started forming stars early.
- Shut down early.
- Are red today.
- Populate dark halos that are much more massive than their stellar mass.

Small galaxies:

- Started forming stars late.
- Are still making stars today.
- Are blue today.
- Populate dark halos that match their stellar mass.

"Downsizing"

Star formation is a wave that started in the largest galaxies and swept down to smaller masses later (Cowie et al. 1996).

Sandy Faber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16 September 2011 Last updated at 13:47 ET

Dwarf galaxies suggest dark matter theory may be wrong

By Leila Battison
Science reporter, Bradford

Thanks to Piero Madau!

small scale issues

Angular momentum

The Eris simulation shows that ΛCDM simulations are increasingly able to form realistic spiral galaxies, as resolution improves and feedback becomes more realistic.

Cusps

WDM doesn't resolve cusp issues. New observations and simulations suggest that observed velocity structure of LSB, dSpiral, dSph galaxies may be consistent with cuspy ΛCDM halos. But the "too big to fail" problem needs solution.

Satellites and Subhalos

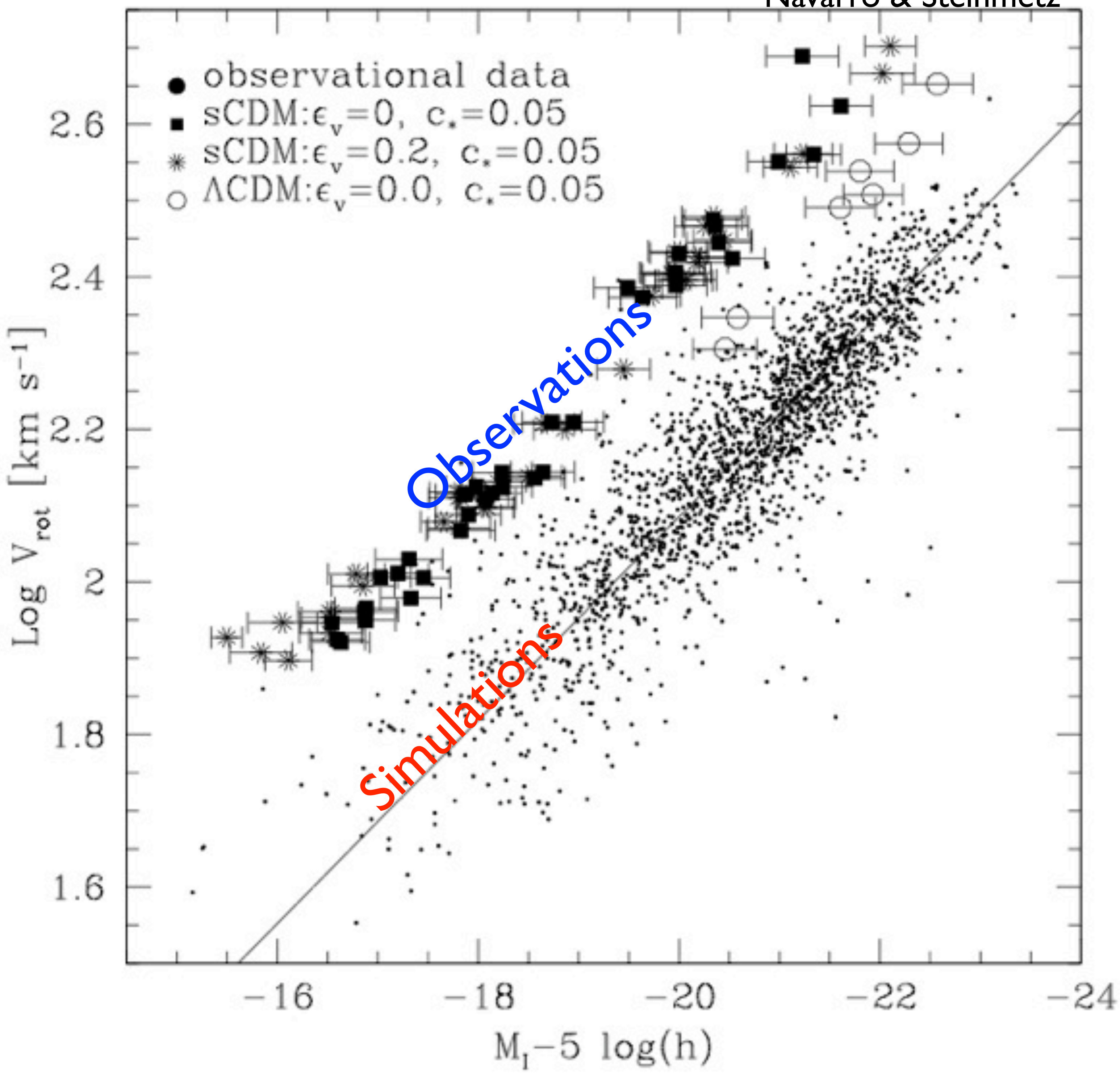
The discovery of many faint Local Group dwarf galaxies is consistent with ΛCDM predictions. Satellites, reionization, lensing flux anomalies, gaps in stellar streams, and Lyα forest data imply that WDM must be Tepid or Cooler.

Can ΛCDM Simulations Form Realistic Galaxies?

The Angular Momentum Problem

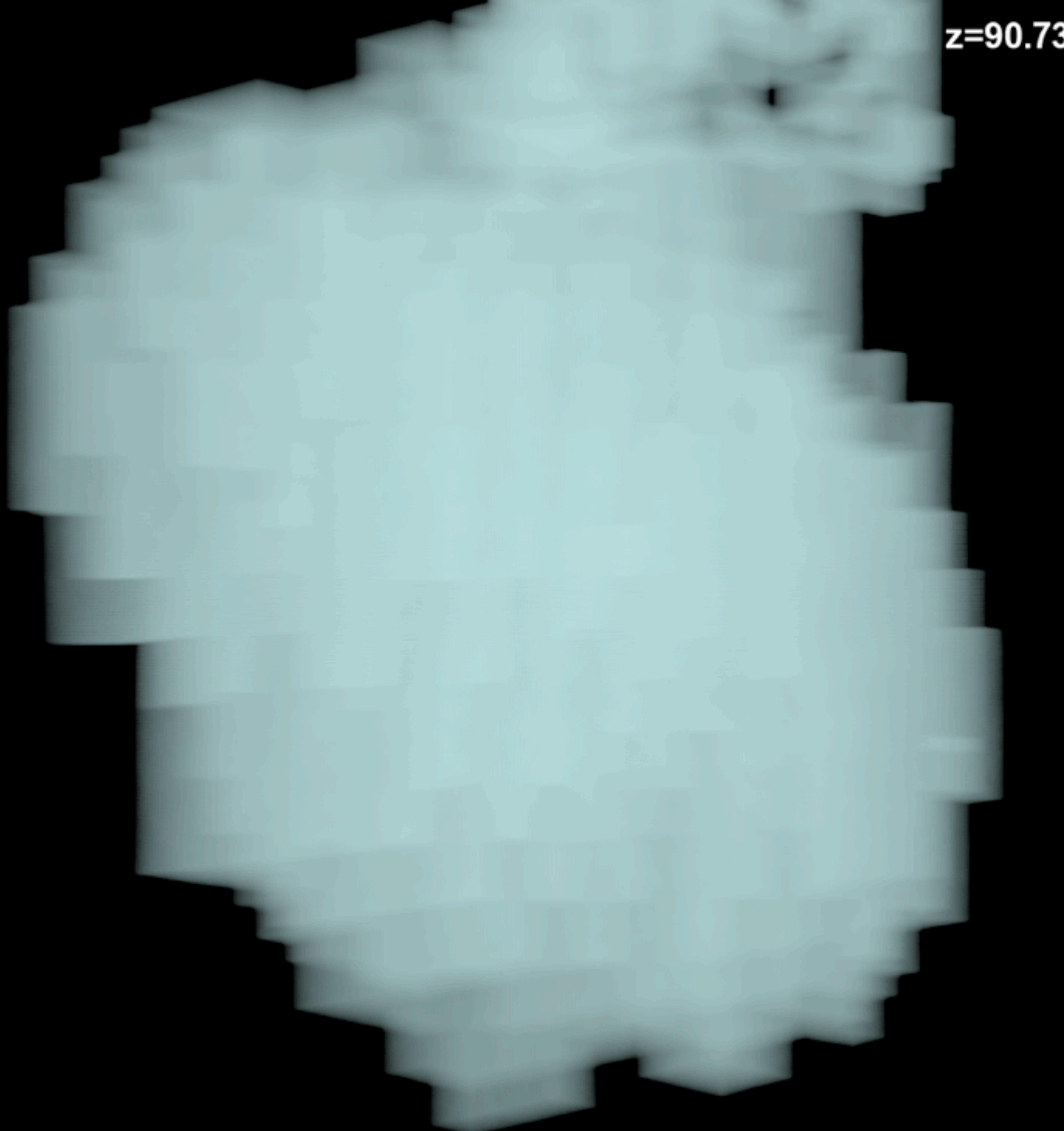
Cooling was too effective particularly in low-mass halos at early times.

Navarro & Steinmetz



"Agreement between model and observations appears to demand substantial <u>revision to the CDM scenario or to the manner in which baryons are thought to assemble</u> and evolve into galaxies in hierarchical universes."

Navarro & Steinmetz 2000 ApJ

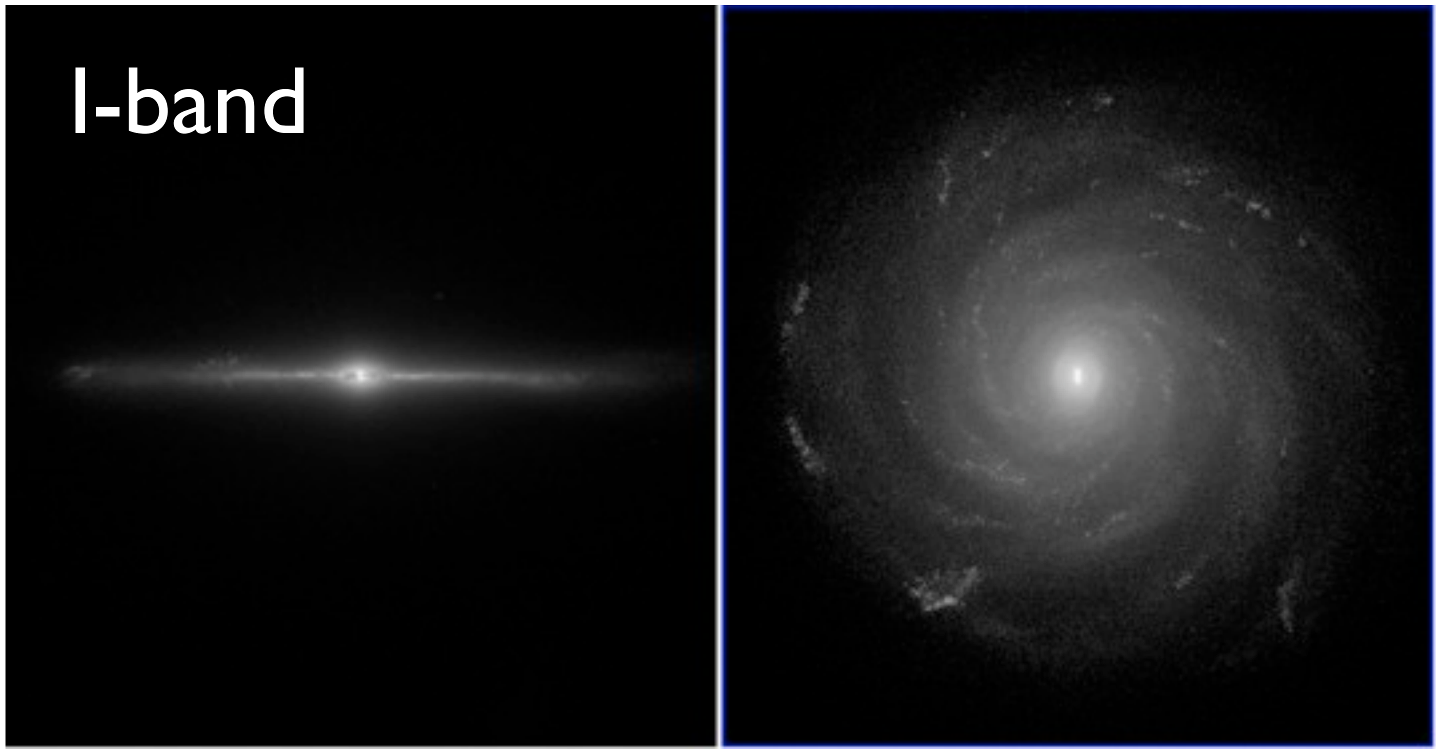


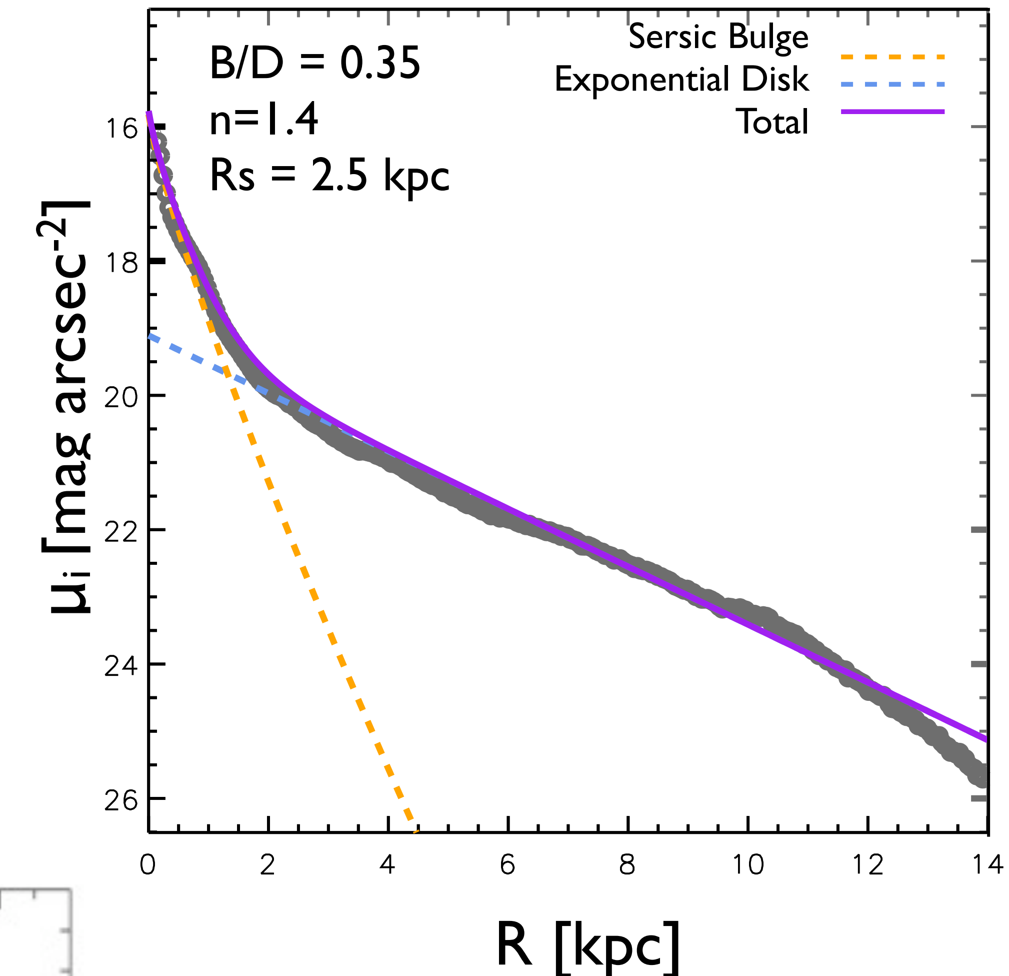
Eris

Simulation

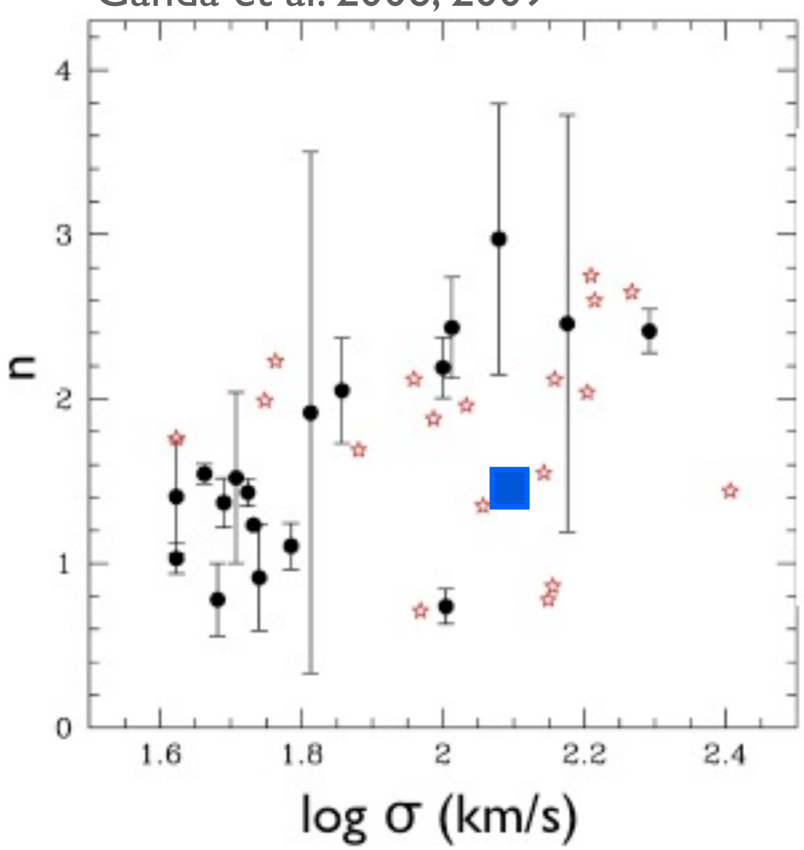
Guedes et al.

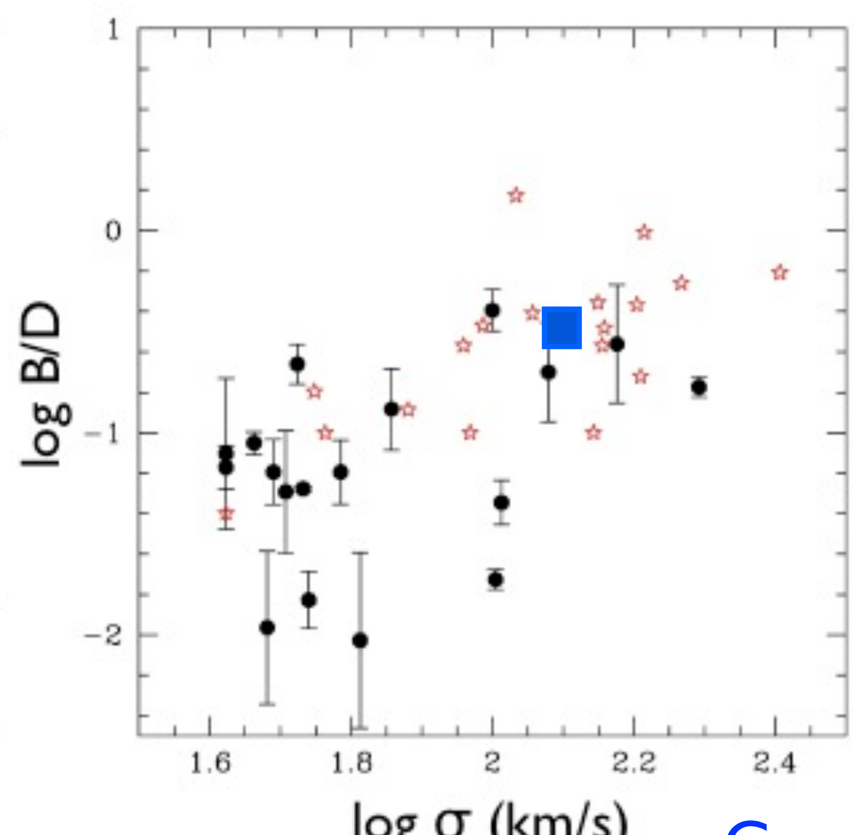
Structural Properties: Eris Bulge-to-Disk Ratio





Ganda et al. 2006, 2009





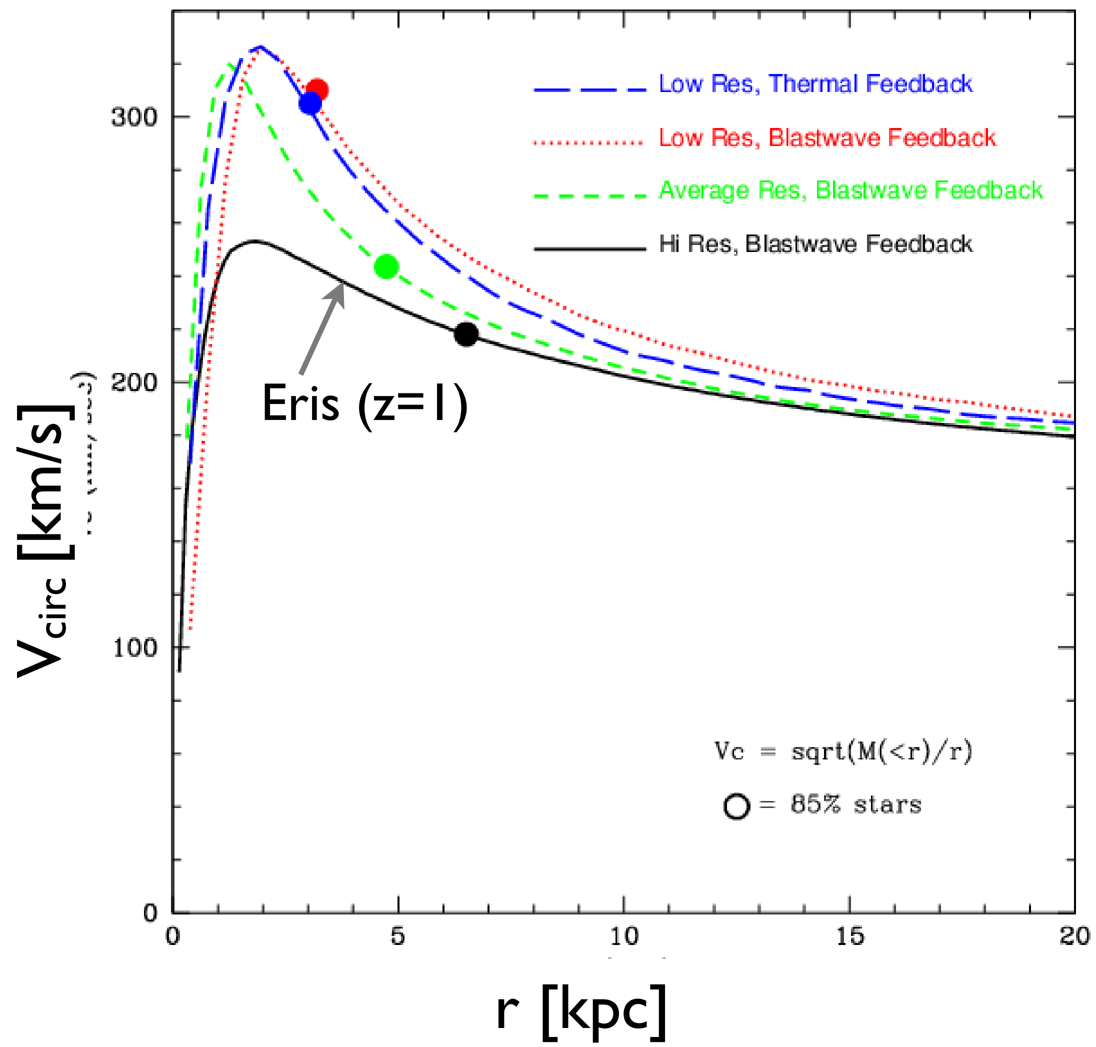
Photometric decomposition in i-band using Galfit (Peng et al. 2002)

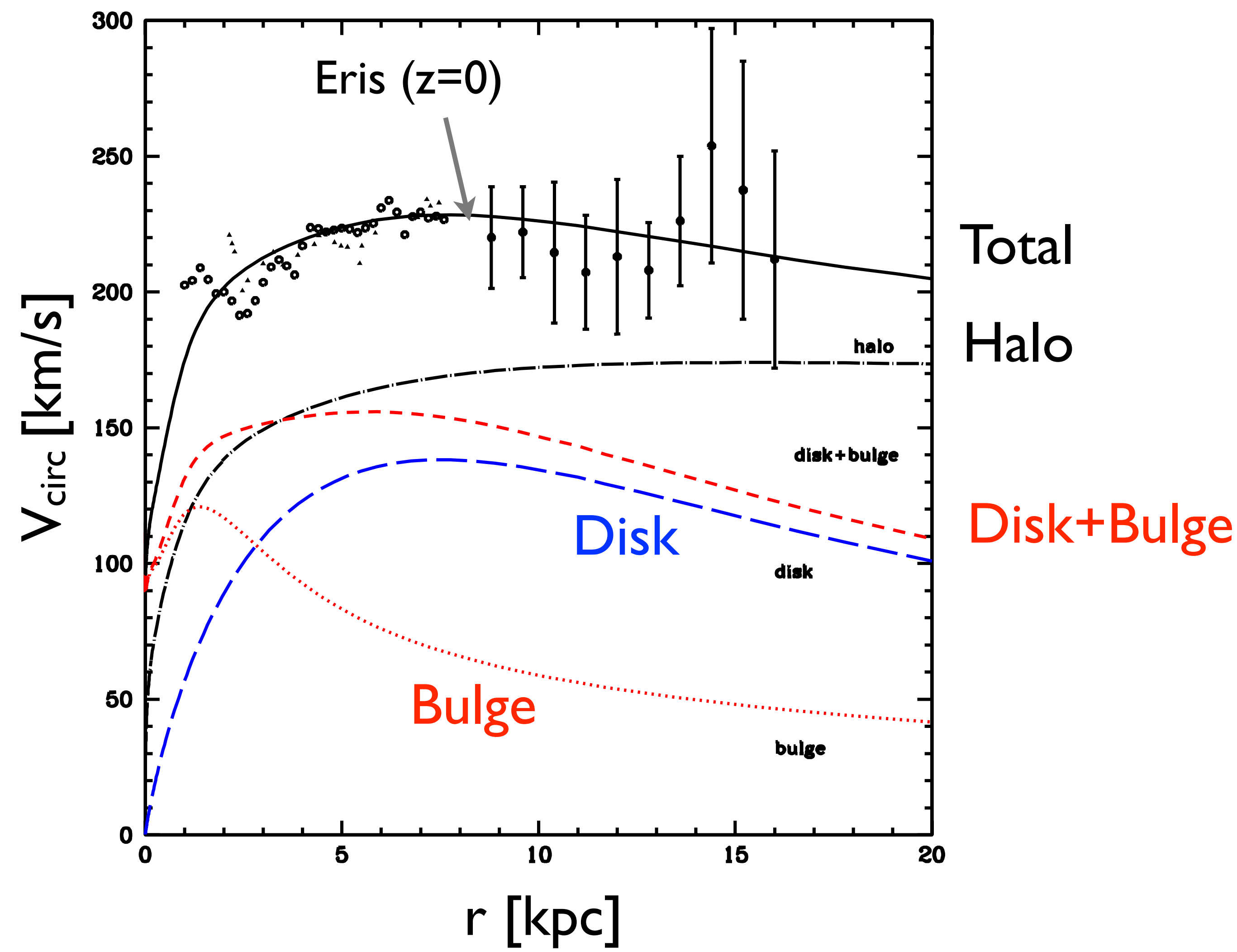
- ● Late-type spirals
- ☆ Early-type spirals
- ■ Eris

Guedes, Callegari, Madau, Mayer 2011 ApJ

No Angular Momentum Problem in the Eris Simulation

Simulations tend to produce too many stars at the center, which translates into steeply rising rotation curves.





Solution:
* Mimic star formation as occurs in real galaxies, i.e. localized, on high-density peaks only.
* Feedback from SN becomes more efficient in removing gas from high-density regions.
These outflows remove preferentially low angular momentum material, suppressing the formation of large bulges.

Guedes, Callegari, Madau, Mayer 2011 ApJ

Cusps

WDM doesn't resolve cusp issues. New observations and simulations suggest that observed velocity structure of LSB and dSpiral galaxies may be consistent with cuspy ΛCDM halos. But the "too big to fail" problem needs solution.

New Developments

- New observations undermine some previous evidence for dark matter cores in dwarf galaxies
- The properties of density cores of dwarf spiral galaxies are inconsistent with expectations from WDM
- New simulations show that gas blowout during evolution of dwarf spiral galaxies can remove cusps
- But the biggest subhalos in MWy size dark matter simulations may be too dense to host the observed satellites

Beware of darkness: A cuspy dark matter halo from stellar kinematics where gas shows a core

NGC 2976 presented in ApJ, Vol. 745, 92, 2012; 10 more galaxies coming in future papers

Joshua J. Adams[1], Joshua D. Simon[1], Karl Gebhardt[2], Guillermo A. Blanc[1], Maximilian H. Fabricius[3], Gary J. Hill[4], Jeremy D. Murphy[2], Remco C.E. van den Bosch[5], Glenn van de Ven[5]

We here present measurements and anisotropic Jeans models for late-type dwarfs obtained from stellar kinematics. Until recently, DM mass profiles in such systems have been obtained exclusively from atomic or ionized gas. The nearby member of the M81 group, NGC 2976 (SAc), has been measured in ionized gas to have a DM core with a strong constraint on the DM power law index of $0.01<\alpha<0.17$ (Simon et al. 2003), where $\alpha=1$ corresponds to the center of an NFW profile. **In our first work on NGC 2976, we confirm that the simplest models from gas kinematics reveal a cored DM halo but find that the stellar kinematic are most consistent with an NFW profile. We advocate the stellar kinematics as more robust due to the tracer's collisionless nature while the gas is subject to more uncertainties from radial motion, warped disks, and pressure support.** We are making an ongoing study by which the type, strength, and conditions of feedback can be constrained from new measurements and comparison to simulations.

Joshua Adams poster at KITP Conference “First Light and Faintest Dwarfs” February 2012

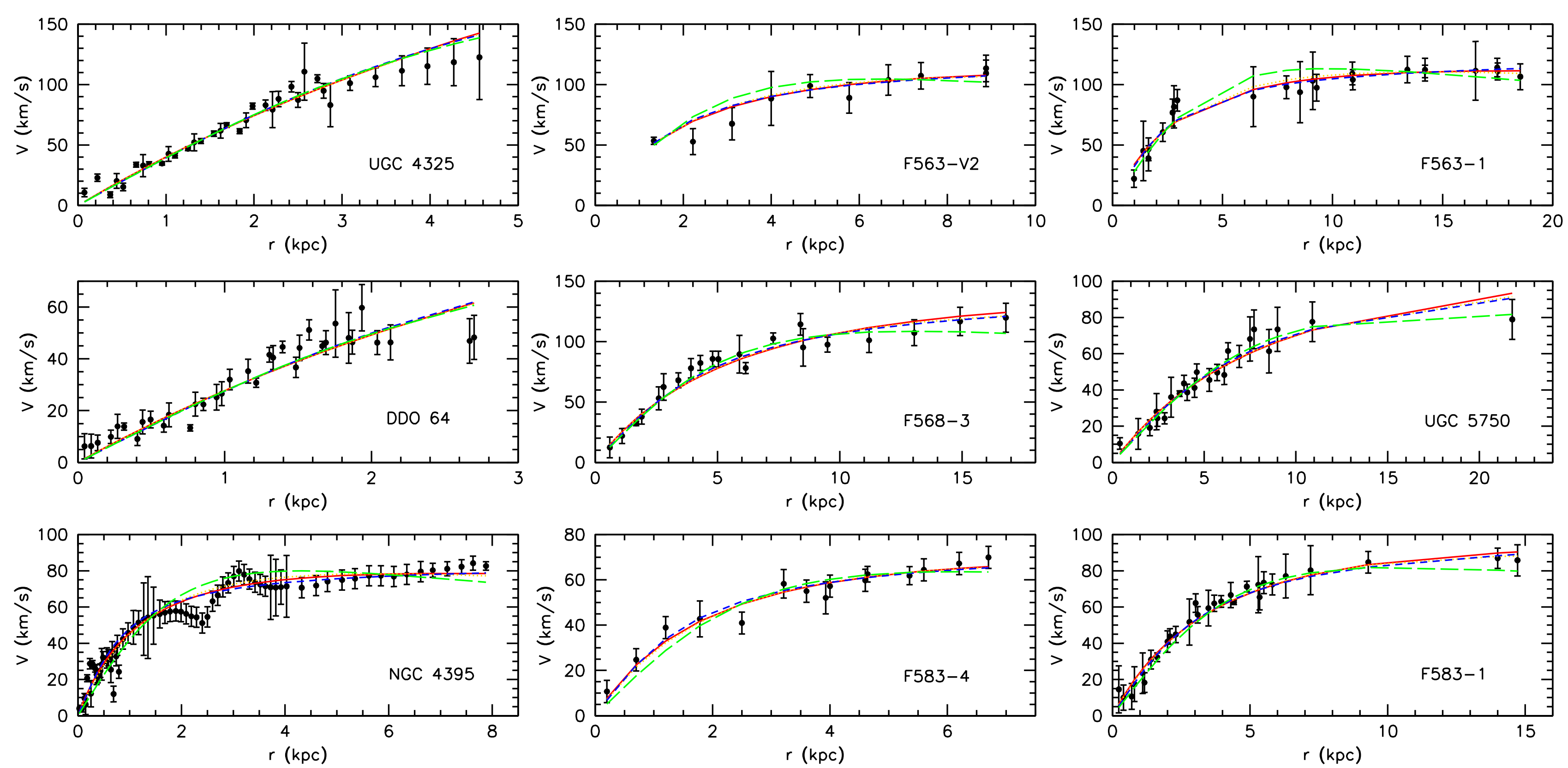
The Case Against Warm or Self-Interacting Dark Matter as Explanations for Cores in Low Surface Brightness Galaxies

2010, ApJ, 710L, 161

Rachel Kuzio de Naray, Gregory D. Martinez, James S. Bullock, Manoj Kaplinghat

Warm dark matter (WDM) and self-interacting dark matter (SIDM) are often motivated by the inferred cores in the dark matter halos of low surface brightness (LSB) galaxies. We test thermal WDM, non-thermal WDM, and SIDM using high-resolution rotation curves of nine LSB galaxies. If the core size is set by WDM particle properties, then even the smallest cores we infer would require primordial phase space density values that are orders of magnitude smaller than lower limits obtained from the Lyman alpha forest power spectra. We also find that the dark matter halo core densities vary by a factor of about 30 while showing no systematic trend with the maximum rotation velocity of the galaxy. This strongly argues against the core size being directly set by large self-interactions (scattering or annihilation) of dark matter. We therefore conclude that the inferred cores do not provide motivation to prefer WDM or SIDM over other dark matter models.

We fit these dark matter models to the data and determine the halo core radii and central densities. While the minimum core size in WDM models is predicted to decrease with halo mass, we find that the inferred core radii increase with halo mass and also cannot be explained with a single value of the primordial phase space density.

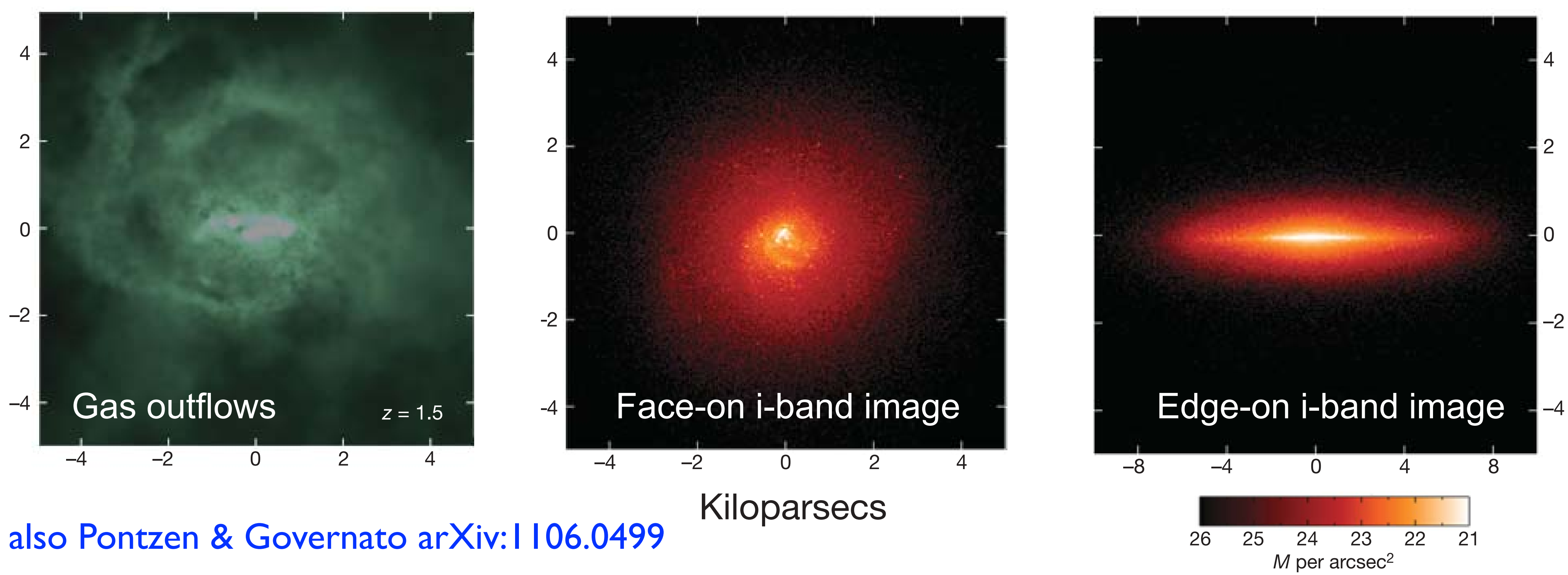


New simulations show that several episodes of gas blowout during evolution of dwarf spiral galaxies can remove cusps

Bulgeless dwarf galaxies and dark matter cores from supernova-driven outflows

F. Governato, C. Brook, L. Mayer, A. Brooks, G. Rhee, J. Wadsley, P. Jonsson, B. Willman, G. Stinson, T. Quinn & P. Madau **Nature** 463, 203 (Jan 2010)

Most observed dwarf galaxies consist of a rotating stellar disk embedded in a massive dark-matter halo with a near-constant-density core. Models based on CDM, however, invariably form galaxies with dense spheroidal stellar bulges and steep central dark-matter profiles, because low-angular-momentum baryons and dark matter sink to the centers of galaxies through accretion and repeated mergers. Here we report hydrodynamical simulations in which the inhomogeneous interstellar medium is resolved. **Strong outflows from supernovae remove low-angular-momentum gas, which inhibits the formation of bulges and decreases the dark-matter density to less than half of what it would otherwise be within the central kiloparsec. The analogues of dwarf galaxies—bulgeless and with shallow central dark-matter profiles—arise naturally in these simulations.** Simulations using the same implementation of star formation and feedback reproduce some global scaling properties of observed galaxies across a range of masses and redshifts.



See also Pontzen & Governato arXiv:1106.0499

Cuspy No More: How Outflows Affect the Central Dark Matter and Baryon Distribution in ΛCDM Galaxies.

F.Governato[1⋆] A.Zolotov[2], A.Pontzen[3], C.Christensen[4], S.H.Oh[5,6], A. M.Brooks[7], T.Quinn[1], S.Shen[8], J.Wadsley[9]

MNRAS in press 2012

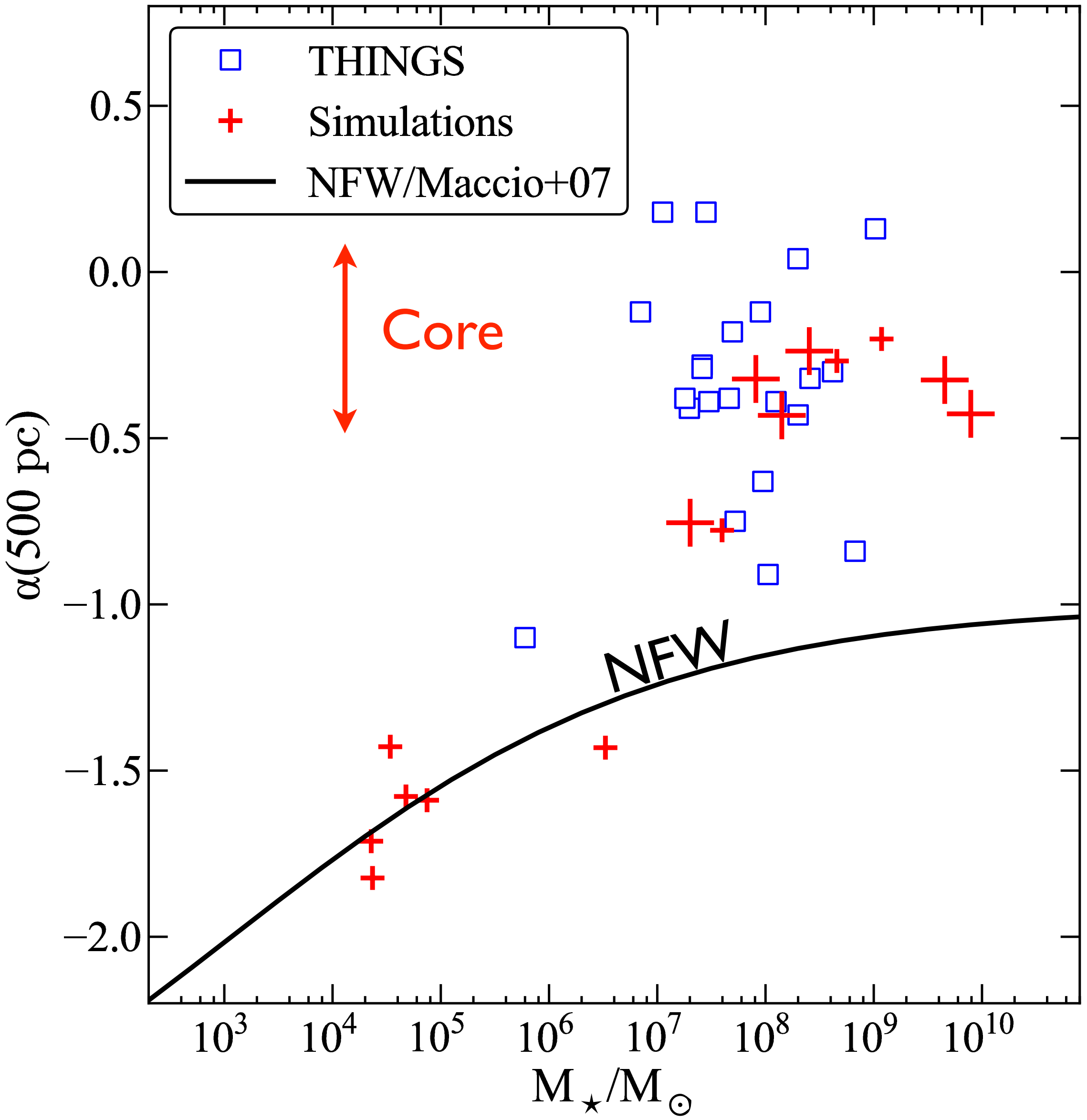


Figure 1. *The slope of the dark matter density profile* α vs stellar mass measured at 500 pc and z=0 for all the resolved halos in our sample. The Solid 'DM-only' line is the slope predicted for the same CDM cosmological model assuming i) the NFW concentration parameter trend given by Macció et al (2007) and ii) the same stellar mass vs halo mass relation as measured in our simulations to convert from halo masses. Large Crosses: haloes resolved with more than 0.5×10^6DM particles within R_{vir}. Small crosses: more than 5×10^4 DM particles. The small squares represent 22 observational data points measured from galaxies from the THINGS and LITTLE THINGS surveys.

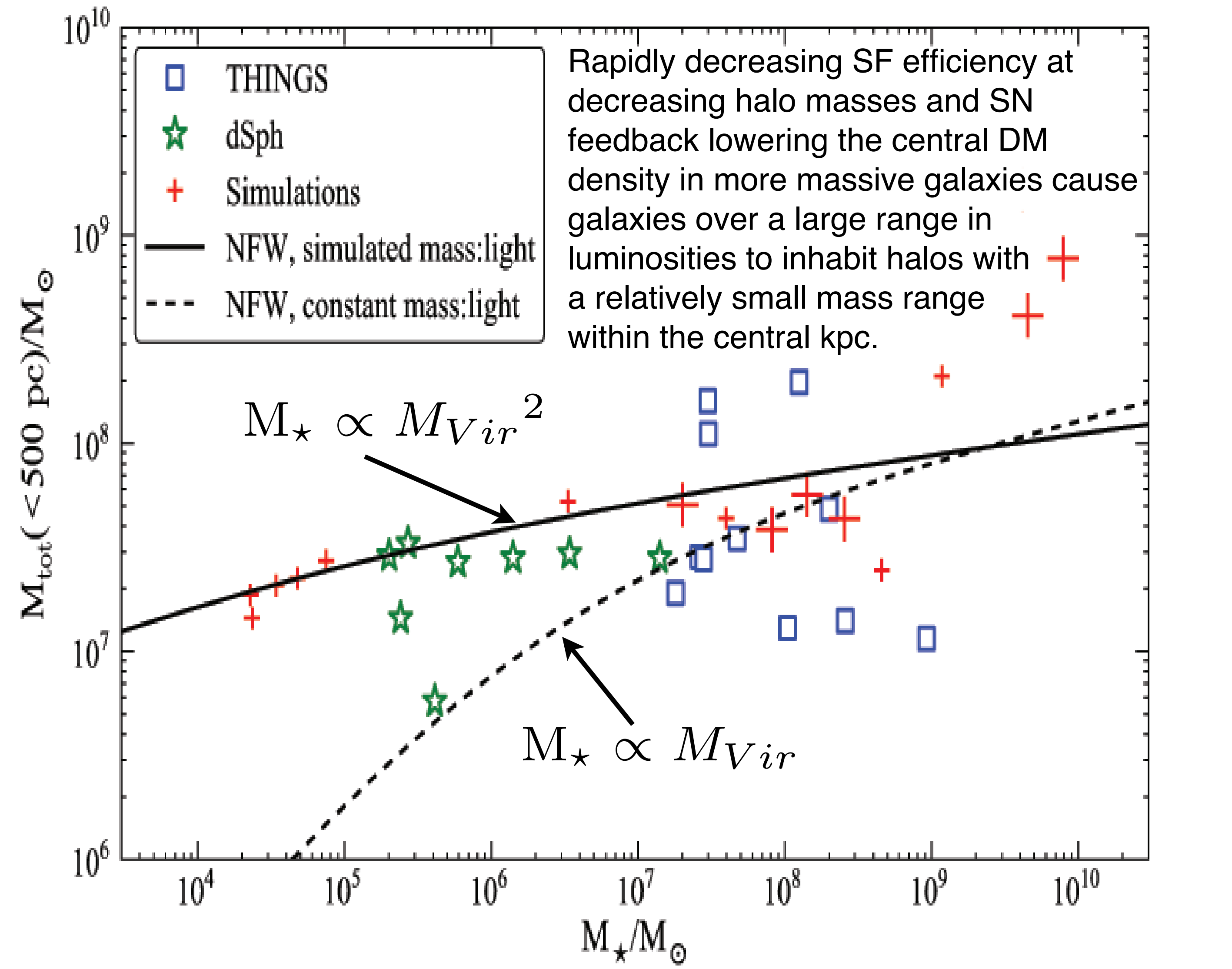


Figure 4. The total mass (baryons and DM) within the central 500 pc as a function of stellar mass: Large and small crosses: simulations. Open squares: galaxies from THINGS (Oh et al. in prep). Stars: dSph from Walker (priv. comm.). Theoretical predictions reproduce the observed flat trend from 10^5 to 10^9 M⊙. This is largely due to the large drop in SF efficiency at small halo masses, that stretches the range of galaxy luminosities over a relatively smaller halo mass range. The solid and dashed lines assume different stellar mass - total halo mass relations. A close fit to the simulations as $M^* \sim M_{Vir}^2$ (solid) and one showing $M^* \sim M_{Vir}$ (dashed). Only when the star formation efficiency is a steep function of halo mass it is possible to reproduce the observed trend, as discussed in §4. More massive galaxies above the solid line have a small bulge component.

A METHOD FOR MEASURING (SLOPES OF) THE MASS PROFILES OF DWARF SPHEROIDAL GALAXIES

MATTHEW G. WALKER[1,2,3] & JORGE PEÑARRUBIA[2] 2011 ApJ

Using separately higher metal stars at lower radiu plus lower metal stars farther out gives dm radial slope inconsistent with NFW at high confidence for Sculptor and Fornax dwarf spheroidal MWy satellites.

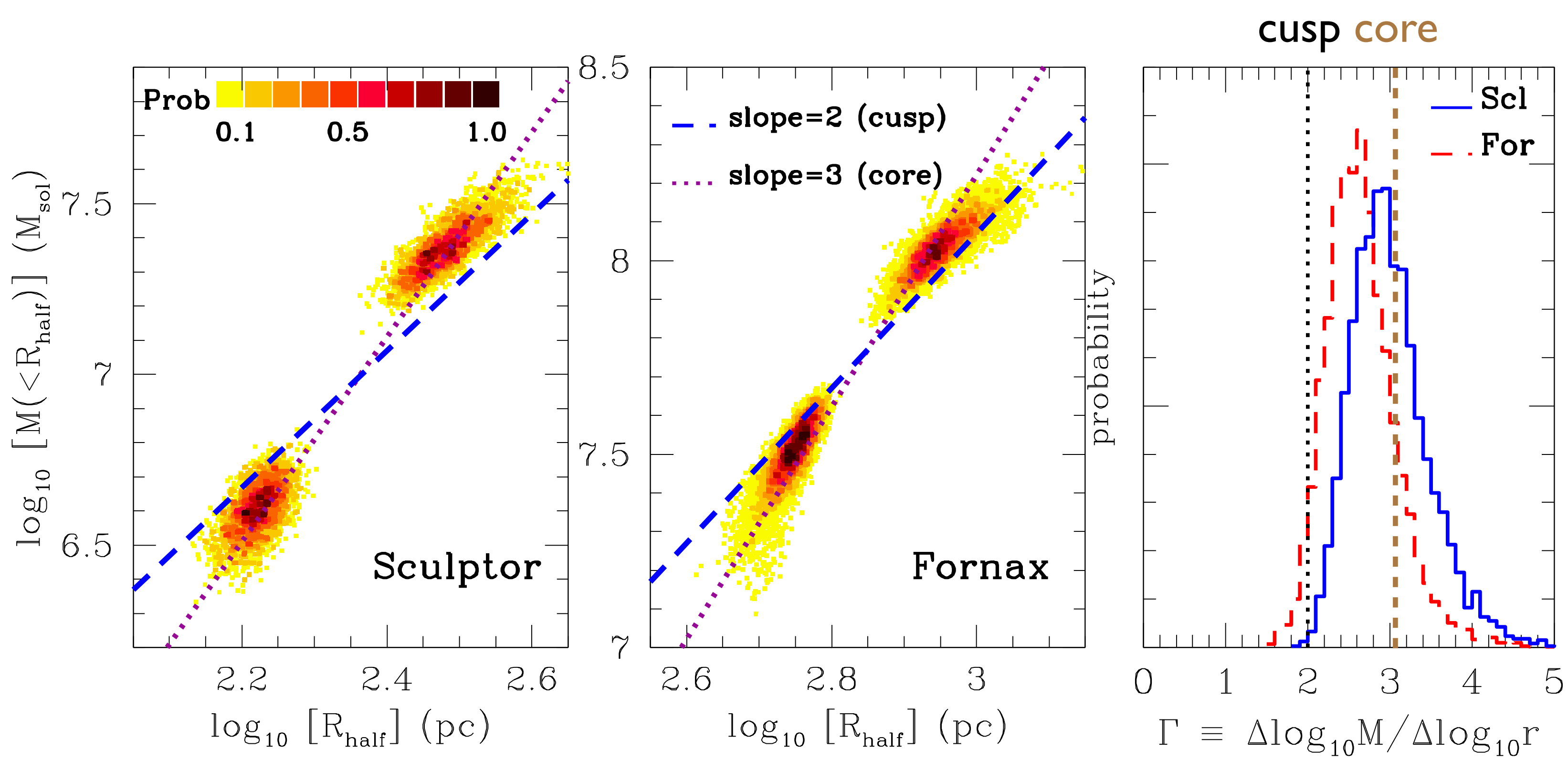


FIG. 10.— *Left, center:* Constraints on halflight radii and masses enclosed therein, for two independent stellar subcomponents in the Fornax and Sculptor dSphs. Plotted points come directly from our final MCMC chains, and color indicates relative likelihood (normalized by the maximum-likelihood value). Overplotted are straight lines indicating the central (and therefore maximum) slopes of cored ($\lim r \to 0\ d \log M/d \log r] = 3$) and cusped ($\lim r \to 0\ d \log M/d \log r] = 2$) dark matter halos. *Right:* Posterior PDFs for the slope Γ obtained for Fornax and Sculptor. The vertical dotted line marks the maximum (i.e., central) value of an NFW profile (i.e., cusp with $\gamma_{DM} = 1$, $\lim r \to 0\ [d \log M/d \log r] = 2$). These measurements rule out NFW and/or steeper cusps ($\gamma DM \geq 1$) with significance $s \geq 96\%$ (Fornax) and $s \geq 99\%$ (Sculptor).

Similar results for Sculptor in Amorisco & Evans 2012 MNRAS. Jardel & Gebhardt present a Schwarzschild model fit to the Fornax dwarf, again favoring core rather than cusp.

Satellites and Subhalos

The discovery of many faint Local Group dwarf galaxies is consistent with ΛCDM predictions. Satellites, reionization, lensing flux anomalies, stellar streams, and Lyα forest data imply that WDM must be Tepid or Cooler.

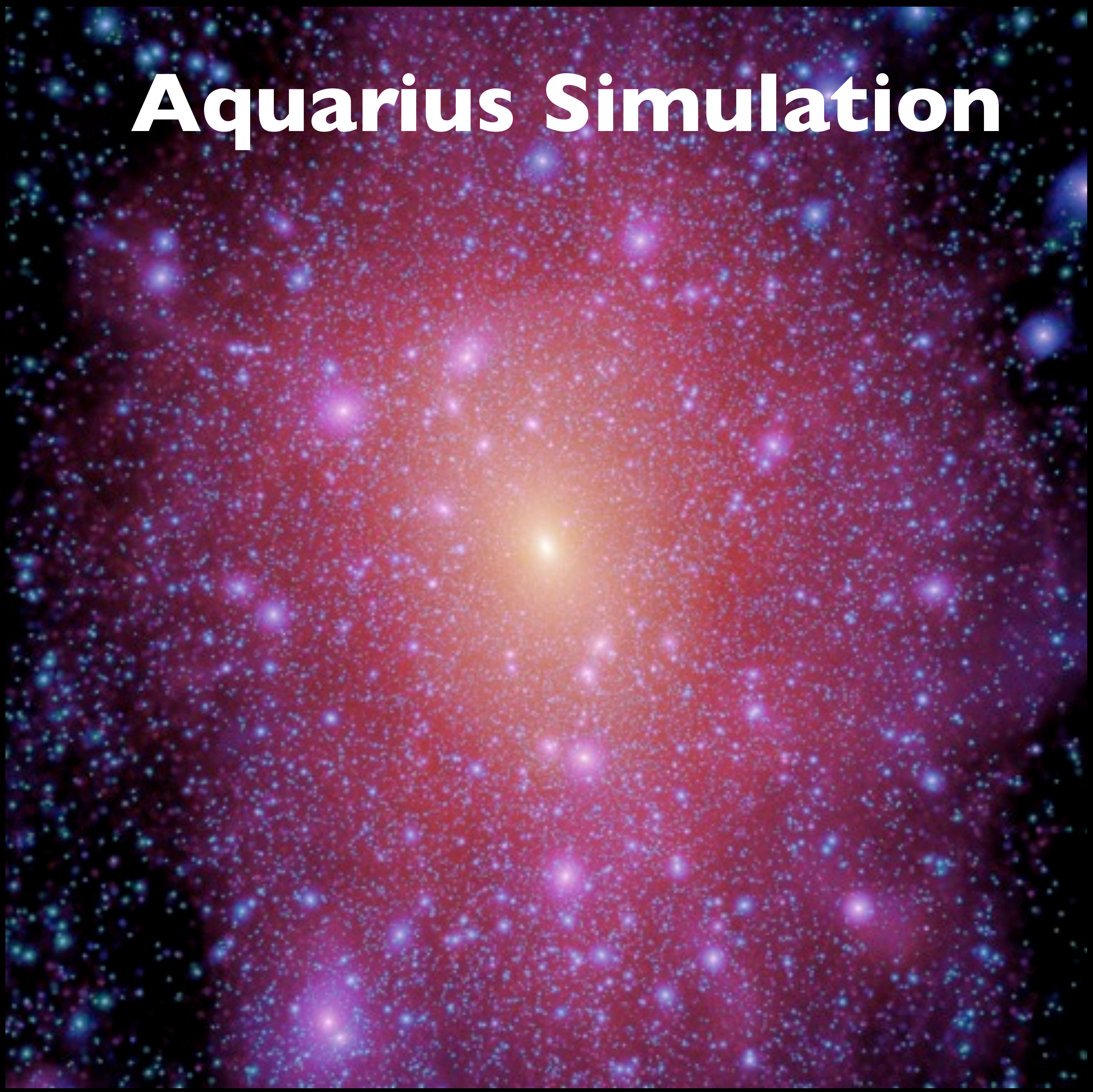
New Developments

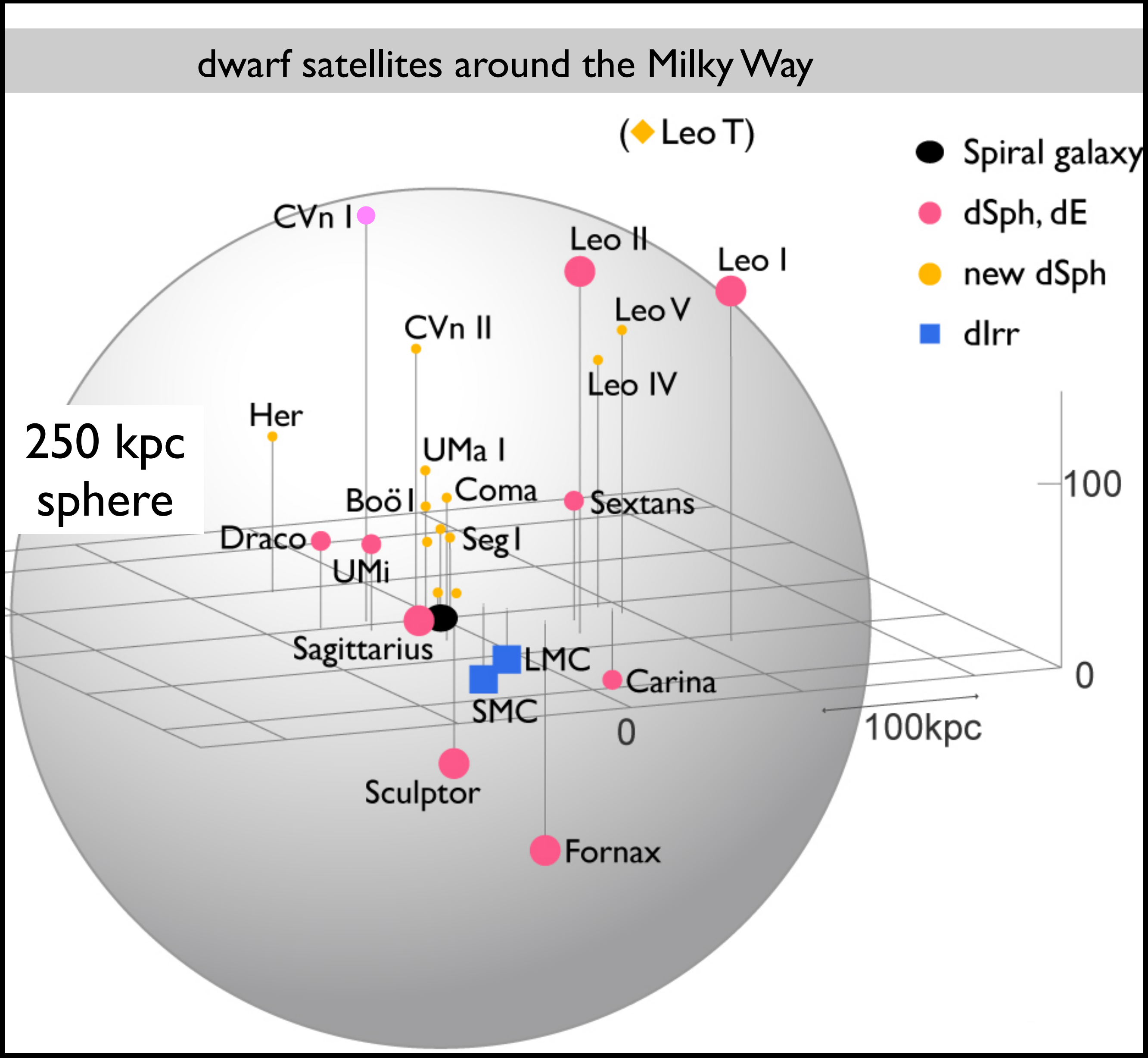
- The "too big to fail" problem appears to be the most serious current challenge for ΛCDM, and may indicate the need for a more complex theory of dark matter.
- High resolution ΛCDM simulation substructure is consistent with quad-lens radio quasar flux and galaxy-galaxy lensing anomalies and indications of substructure by stellar stream gaps.
- ΛCDM predicts that there is a population of low-luminosity stealth galaxies around the Milky Way. Will new surveys with bigger telescopes find them?

The “too big to fail” problem

ΛCDM subhalos vs. Milky Way satellites

“Missing satellites”: Klypin et al. 1999, Moore et al. 1999





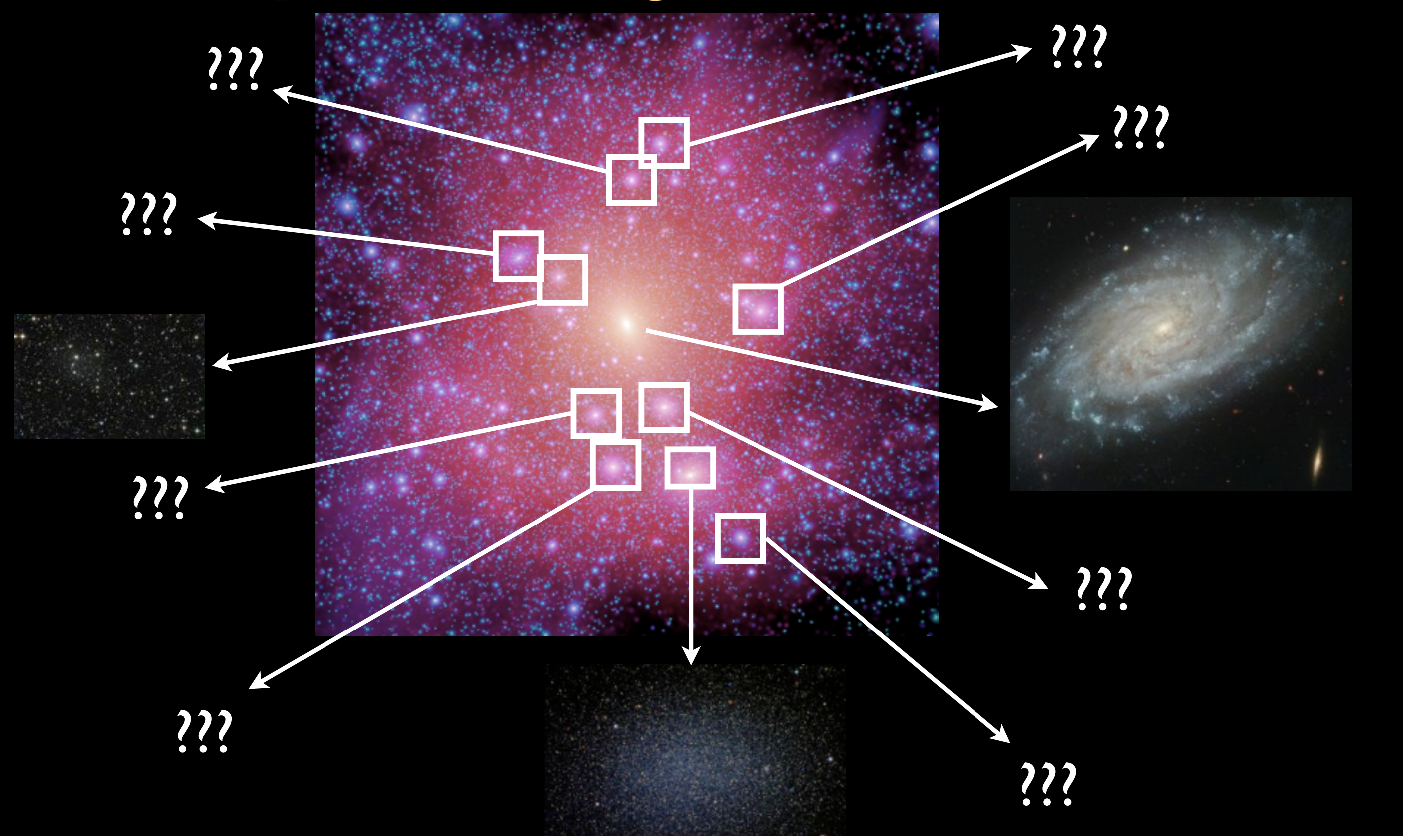
$>10^5$ identified subhalos

12 bright satellites $(L_V > 10^5 L_\odot)$

V. Springel / Virgo Consortium

S. Okamoto

Of the ~10 biggest subhalos, ~8 cannot host *any* known bright MW satellite

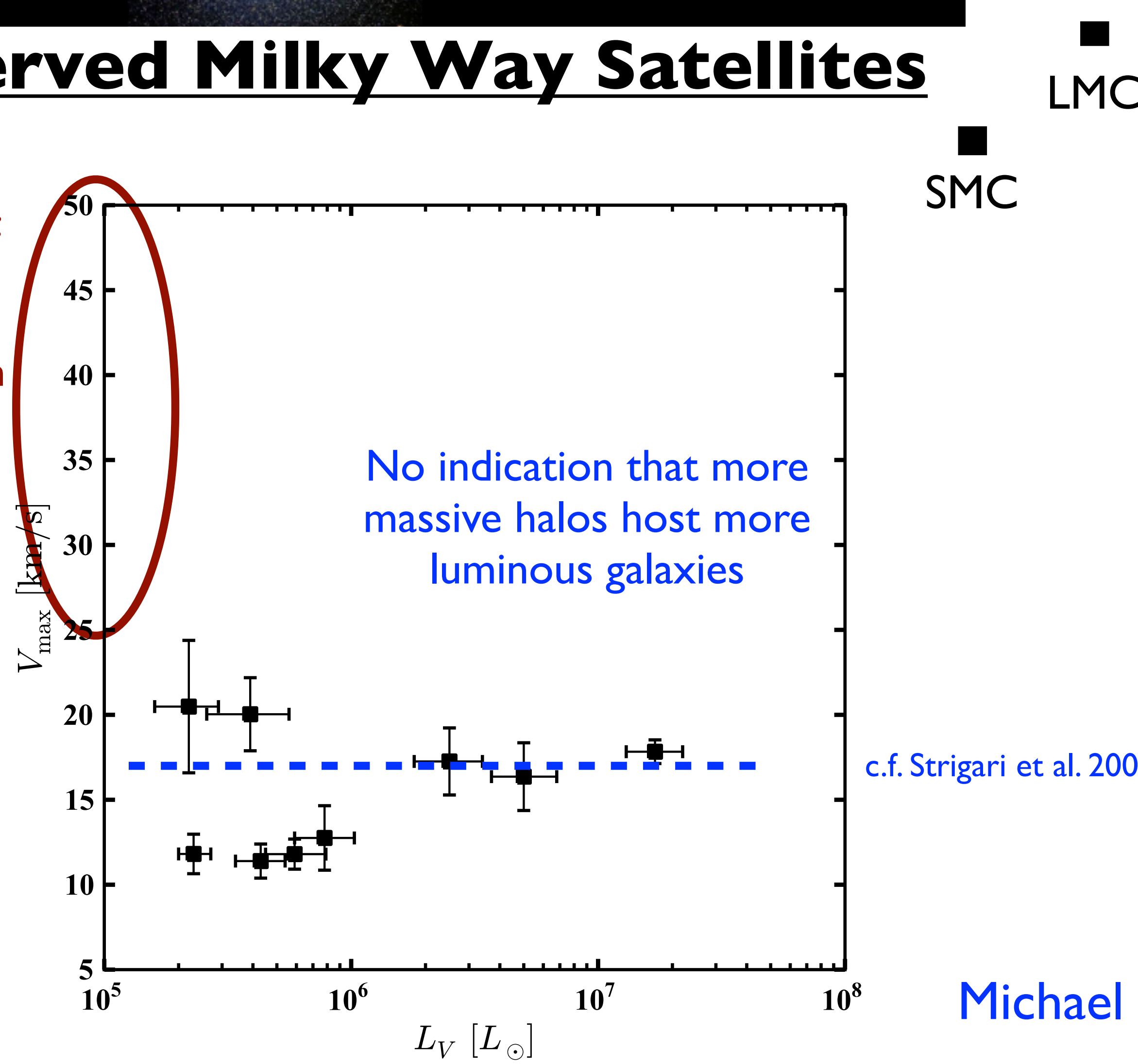


Observed Milky Way Satellites

"massive failures": highest resolution LCDM simulations predict ~10 subhalos in this range in the MW, but we don't see ***any*** such galaxies [except Sagittarius (?)]

All of the bright MW dSphs are consistent with $V_{\rm max} \lesssim 25\,{\rm km/s}$

(see also Strigari, Frenk, & White 2010)



c.f. Strigari et al. 2008

Possible Solutions to "too big to fail"

Baryons strongly modify the structure of subhalos?

The Milky Way is anomalous?

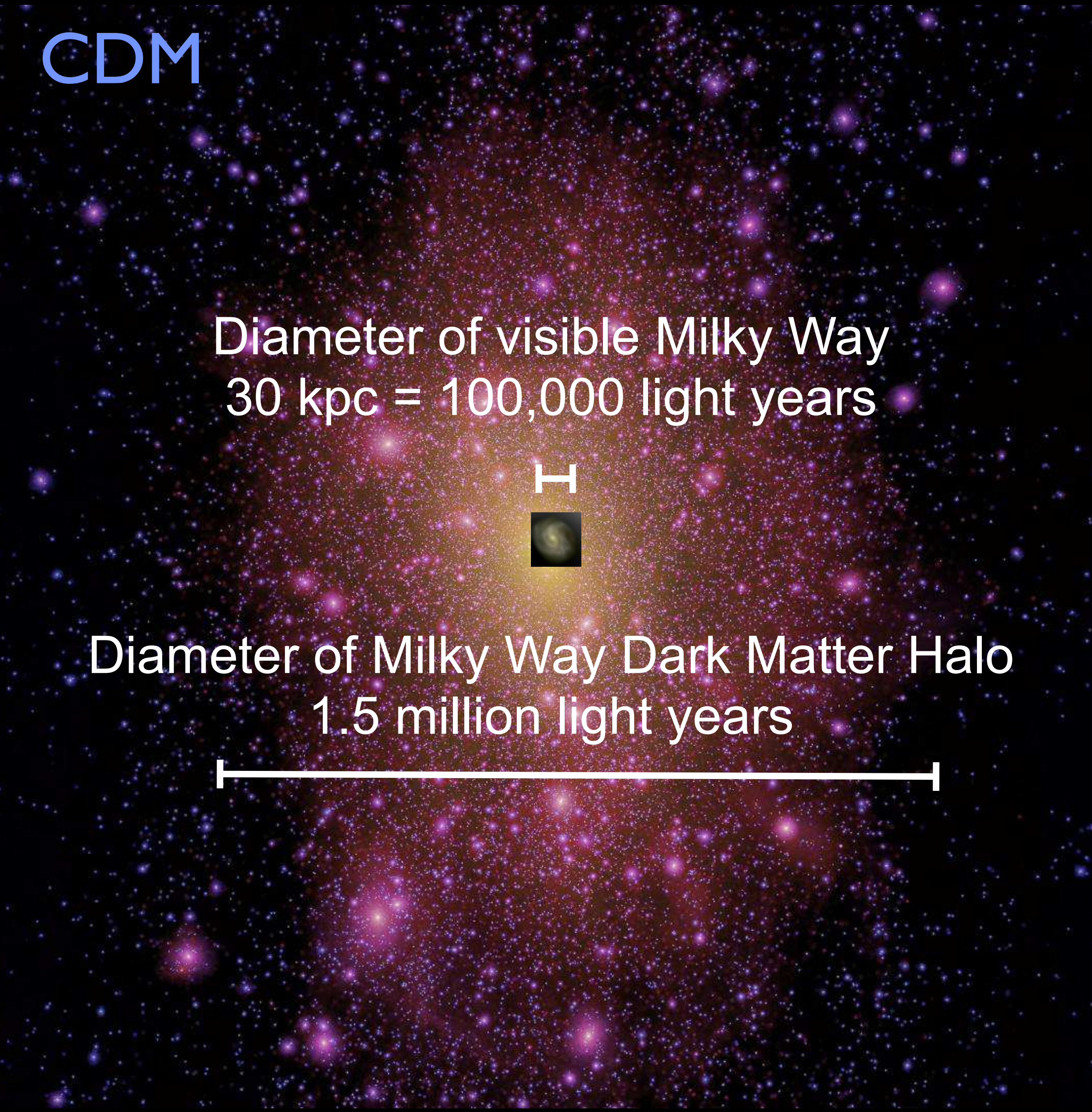
The Milky Way has a low mass dark matter halo?

Galaxy formation is stochastic at low masses?

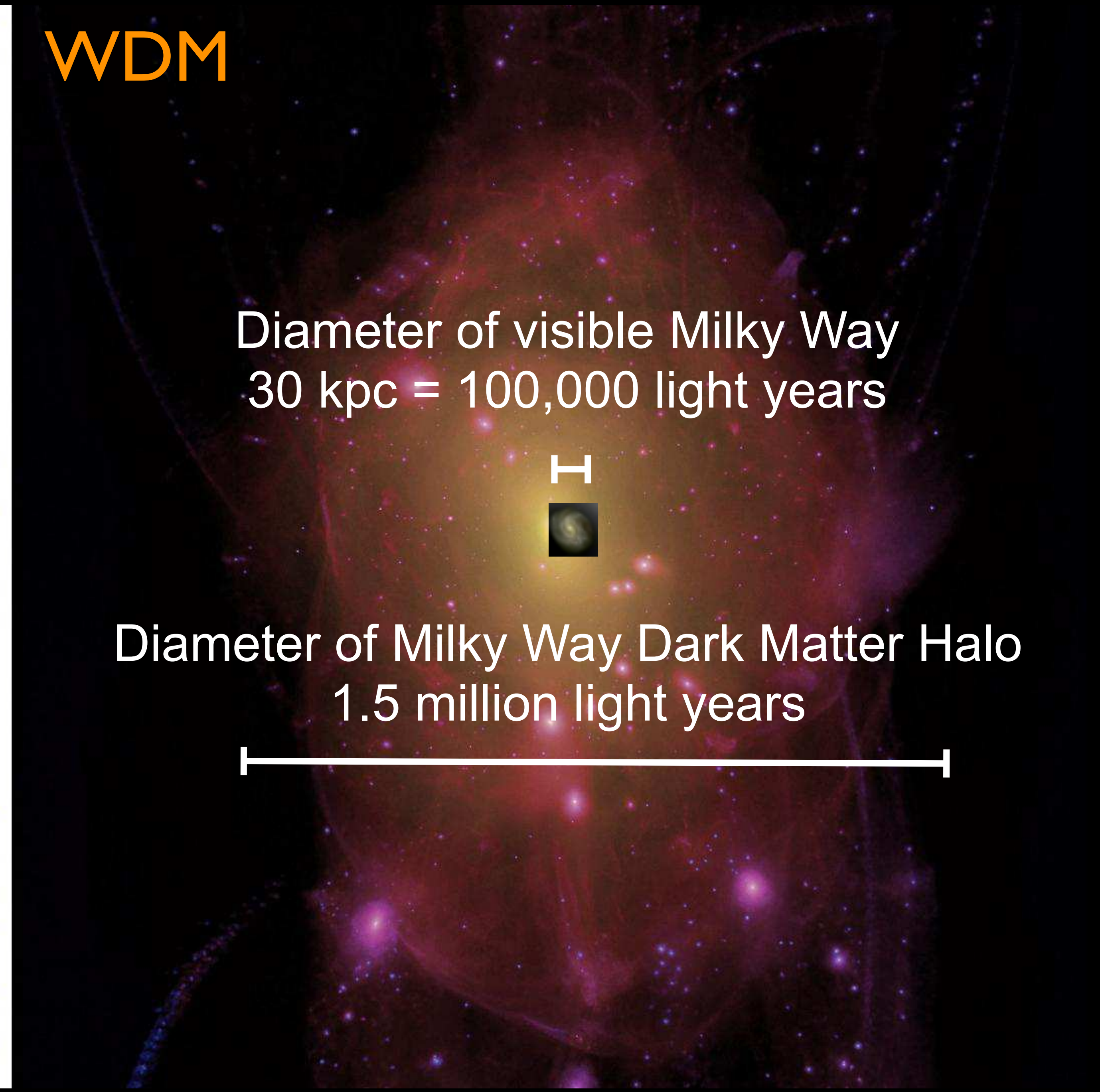
Dark matter is not just CDM -- maybe WDM or even self-interacting?

(Or maybe existing high-resolution CDM simulations are being misinterpreted?)

Michael Boylan-Kolchin, Bullock, Kaplinghat 2011, 2012



Aquarius simulation. Springel et al. 2008



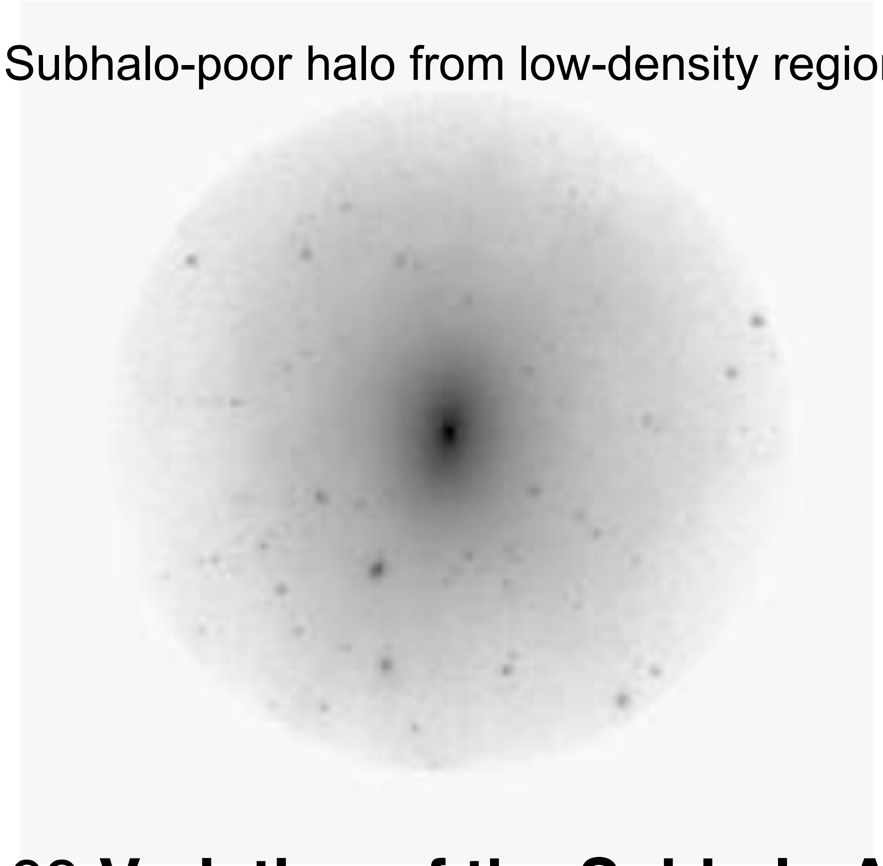
Lovell, Eke, Frenk, et al. arXiv:1104.2929

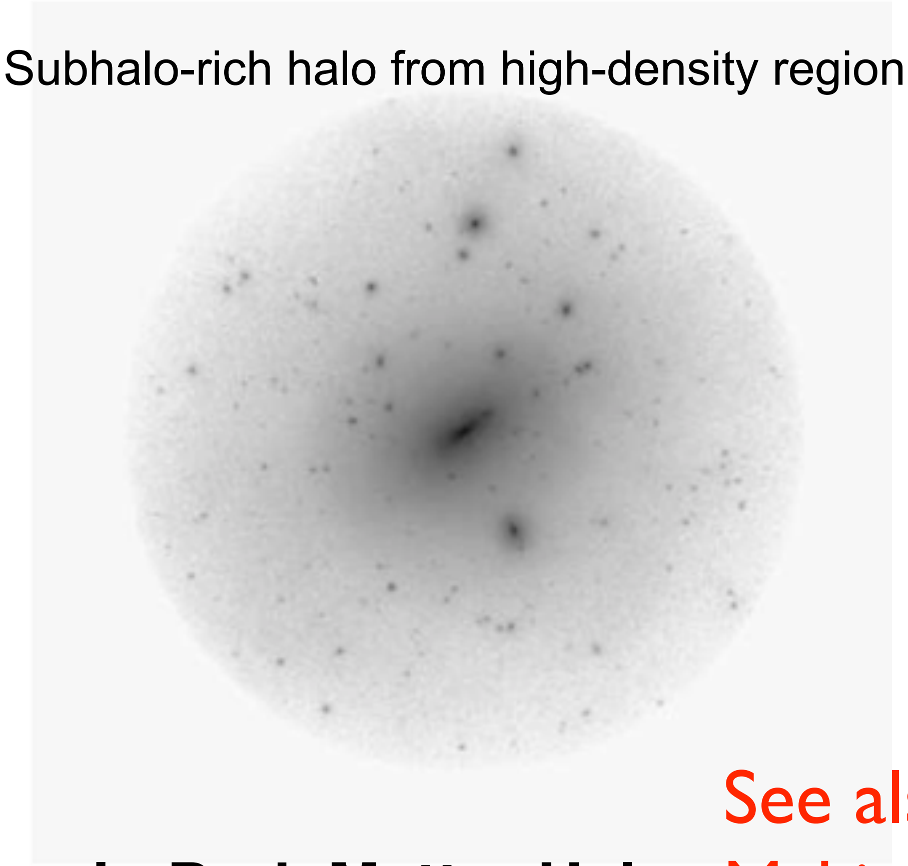
WDM simulation at right has no "too big to fail" subhalos, but it doesn't lead to the right systematics to fit dwarf galaxy properties as Kuzio de Naray et al. showed. It also won't have the subhalos needed to explain radio flux anomalies and gaps in stellar streams.

Possible Solution: Milky-Way-size halos in low-density regions have fewer DM satellites, according to new simulations.

Environmental effect on the subhalo abundance -- a solution to the missing dwarf problem, Tomoaki Ishiyama, Toshiyuki Fukushige, Junichiro Makino PASJ, 60, 13 (2009a)

We have performed simulation of a single large volume and measured the abundance of subhalos in all massive halos. We found that the variation of the subhalo abundance is very large, and those with largest number of subhalos correspond to simulated halos in previous studies. The subhalo abundance depends strongly on the local density of the background. Halos in high-density regions contain large number of subhalos. Our galaxy is in the low-density region. For our simulated halos in low-density regions, the number of subhalos is within a factor of three to that of our galaxy. We argue that the ``missing dwarf problem'' is not a real problem but caused by the biased selection of the initial conditions in previous studies, which were not appropriate for field galaxies.





Ishiyama+09 **Variation of the Subhalo Abundance in Dark Matter Halos**

Halos formed earlier have smaller number of subhalos at present. ApJ, 696, 2115 (2009)

See also Ishiyama, Makino, + CosmoGrid Sim arXiv:1101:2020

New miniBolshoi simulation will provide statistics

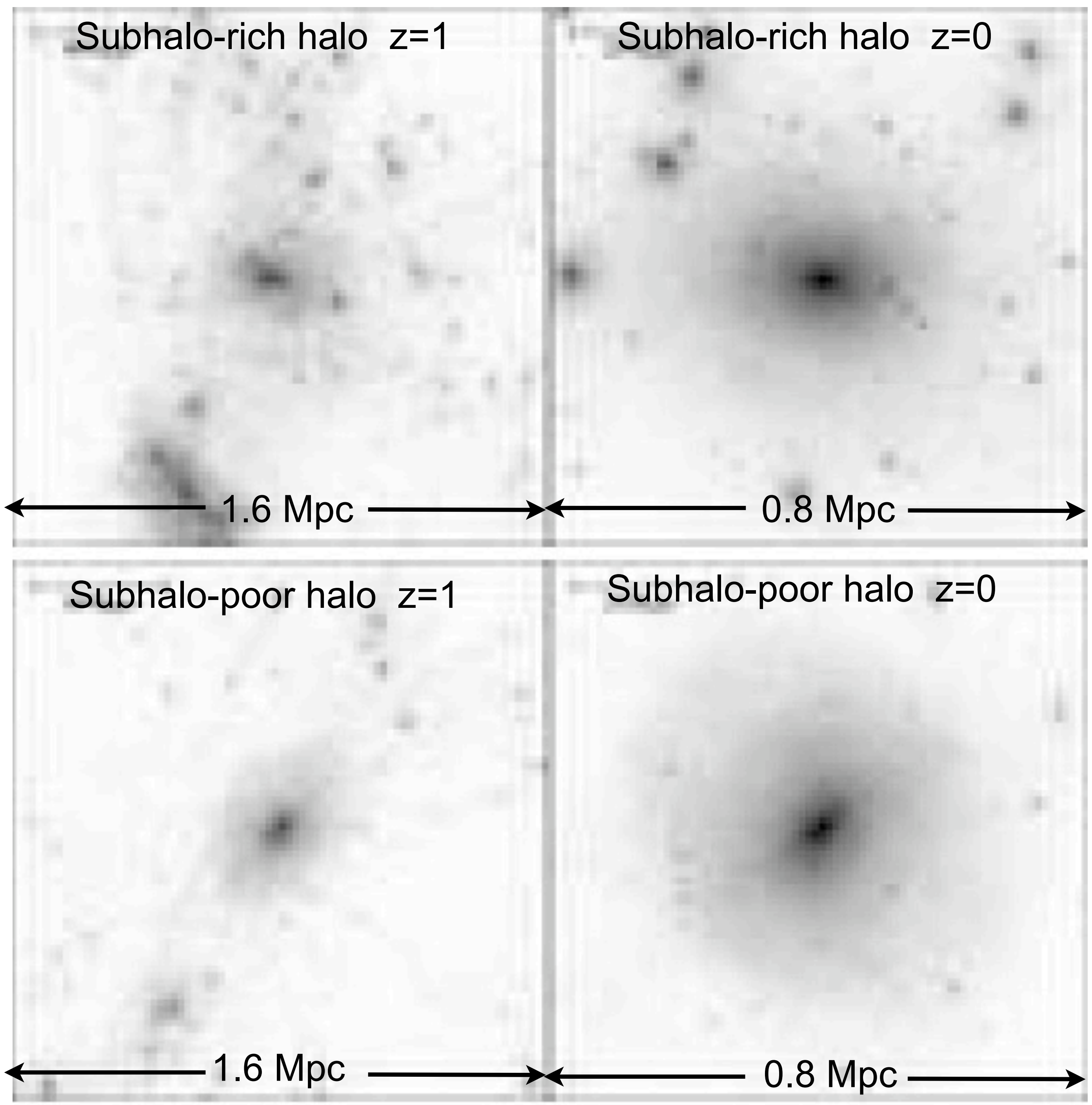
Milky-Way-size halos have large variation in number of DM satellites, according to new simulations.

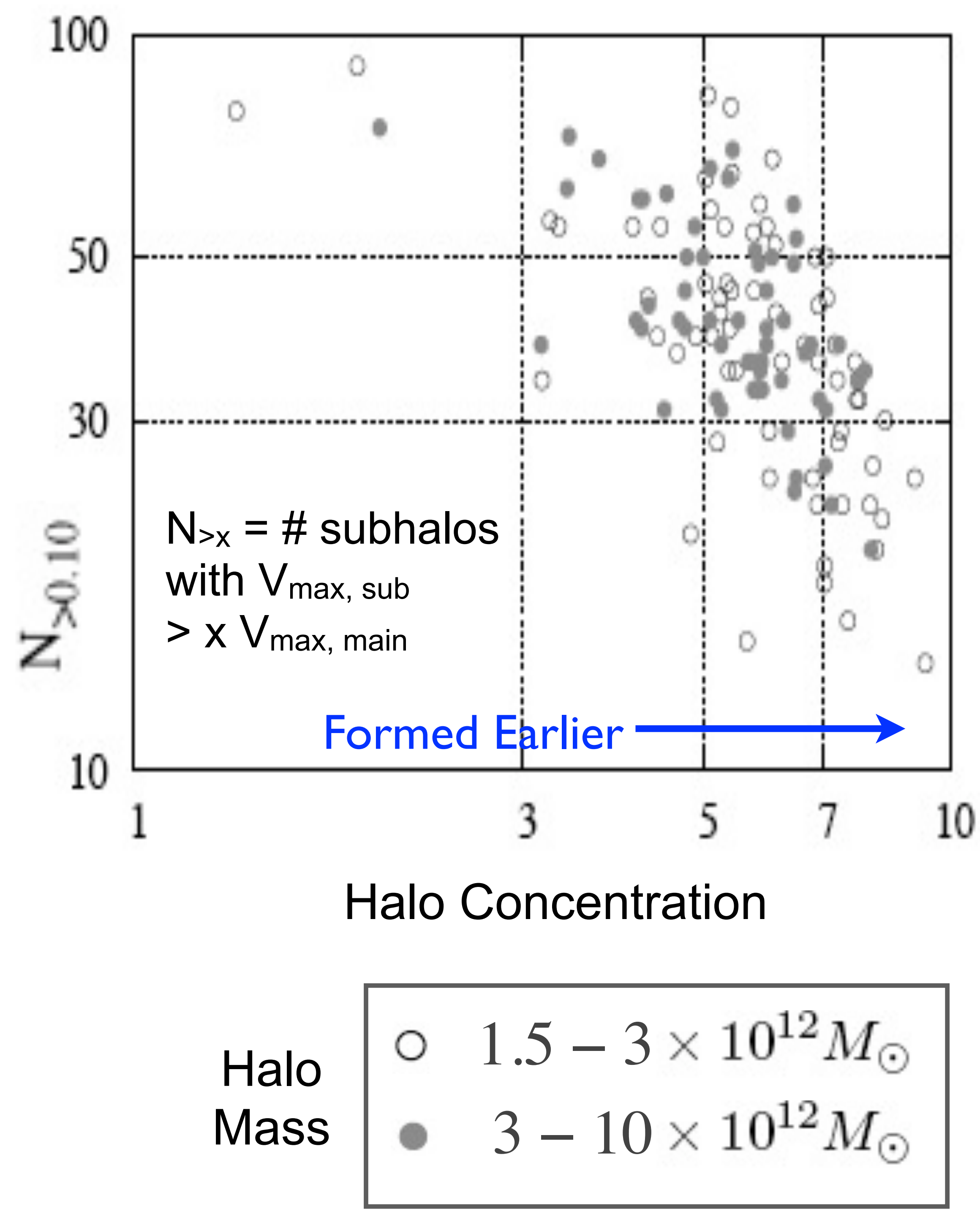
Ishiyama, Fukushige, and Makino

Variation of the Subhalo Abundance in Dark Matter Halos ApJ, 696, 2115 (2009b)

Galaxy halos formed earlier have higher concentration and smaller number of subhalos at present .

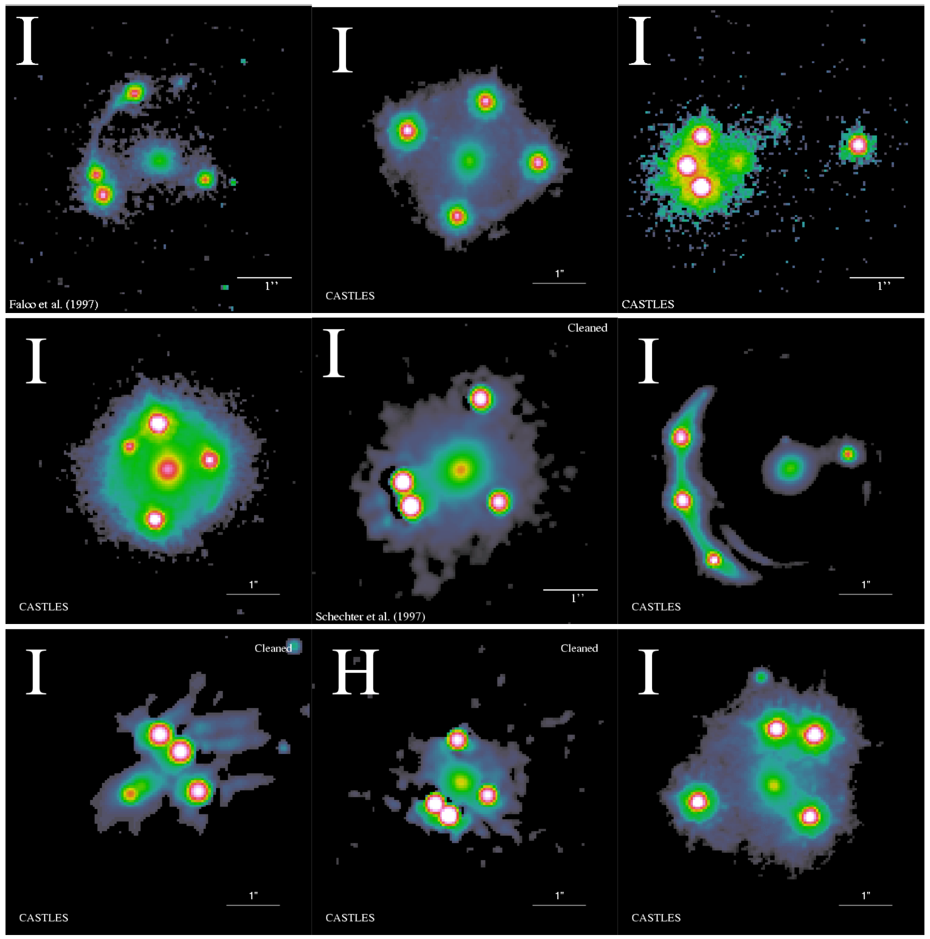
Mass resolution 1x106 Msun (3x better than 2009a). Force resolution 700 pc (2x better than 2009a).





Radio flux-ratio anomalies

Quasar lenses



(CASTLES project, http://www.cfa.harvard.edu/castles)

Radio flux-ratio anomalies ⇒

Strong evidence for dark matter clumps ~ 10^6 - 10^8 Msun as expected in ΛCDM

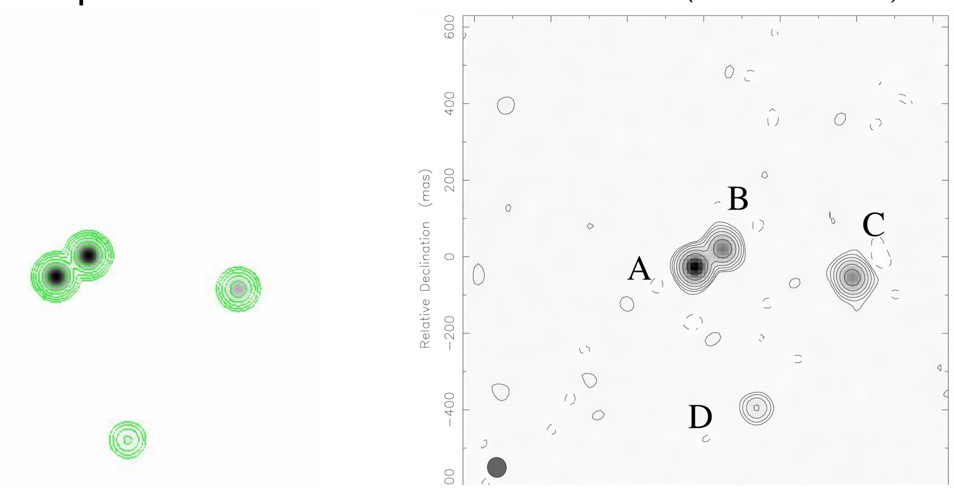
Flux ratio anomalies are generic

"Easy" to explain image positions (even to ~0.1% precision)

- ellipsoidal galaxy
- tidal forces from environment

But hard to explain flux ratios!

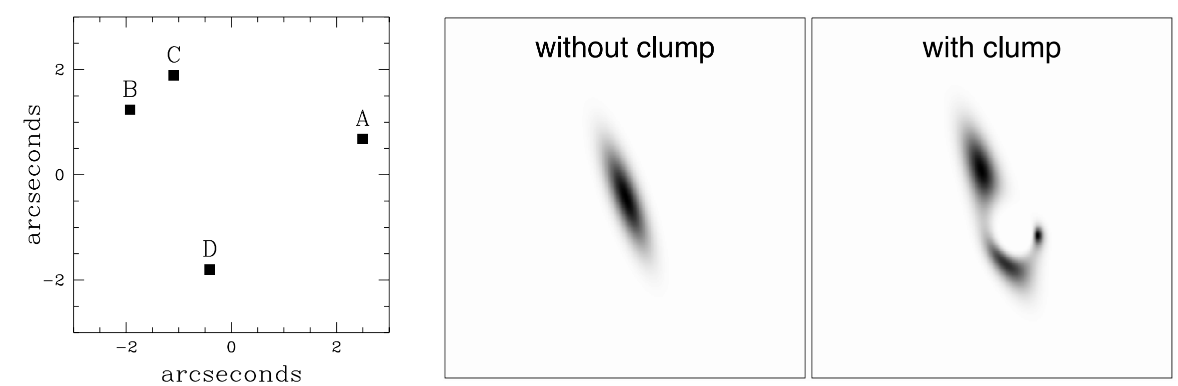
expected — observed *(Marlow et al. 1999)*



Substructure and lensing

Q) What happens if lens galaxies contain mass clumps?

A) The clumps distort the images on small scales.



(cf. Mao & Schneider 1998; Metcalf & Madau 2001; Chiba 2002)

Chuck Keeton

The Aquarius simulations have not quite enough substructure to explain quad-lens radio quasar flux anomalies -- but perhaps including baryons in simulations will help.

Effects of dark matter substructures on gravitational lensing: results from the Aquarius simulations

D. D. Xu, Shude Mao, Jie Wang, V. Springel, Liang Gao, S. D. M. White, Carlos S. Frenk, Adrian Jenkins, Guoliang Li and Julio F. Navarro MNRAS 398, 1235–1253 (2009)

We conclude that line-of-sight structures can be as important as intrinsic substructures in causing flux-ratio anomalies. ... This alleviates the discrepancy between models and current data, but a larger observational sample is required for a stronger test of the theory.

Effects of Line-of-Sight Structures on Lensing Flux-ratio Anomalies in a ΛCDM Universe

D. D. Xu, Shude Mao, Andrew Cooper, Liang Gao, Carlos S. Frenk, Raul Angulo, John Helly MNRAS (2012)

We investigate the statistics of flux anomalies in gravitationally lensed QSOs as a function of dark matter halo properties such as substructure content and halo ellipticity. ... The constraints that we are able to measure here with current data are roughly consistent with ΛCDM N-body simulations.

Constraints on Small-Scale Structures of Dark Matter from Flux Anomalies in Quasar Gravitational Lenses

R. Benton Metcalf, Adam Amara MNRAS 419, 3414 (2012)

Substructure in lens galaxies: first constraints on the mass function

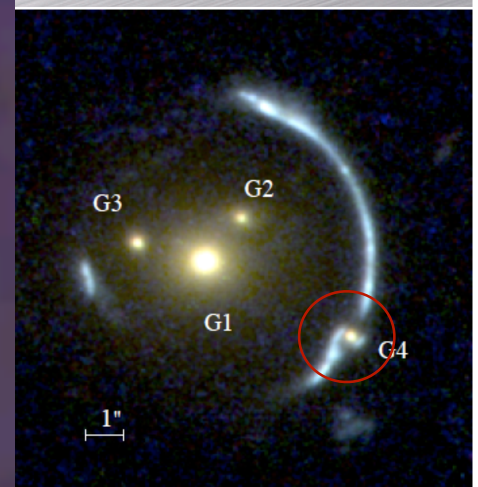
Simona Vegetti (MIT)

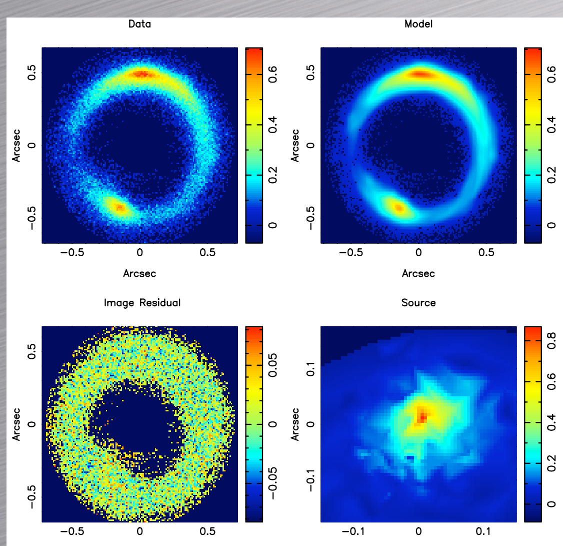
Gravitational detection of a low-mass dark satellite galaxy at cosmological distance, 2012 Nature

Talk at KITP conference "First Light and Faintest Dwarfs"

How do we recognise the effect of substructure?

Extended galaxy





Substructure as a truncated pseudo Jaffe

$$M_{sub} = (1.9 \pm 0.1) \times 10^8 M_\odot$$

$$M(<0.6) = (1.15 \pm 0.06) \times 10^8 M_\odot$$

$$M(<0.3) = (7.24 \pm 0.6) \times 10^7 M_\odot$$

Substructure as SIS

$$M(<0.3) = 3.4 \times 10^7 M_\odot$$

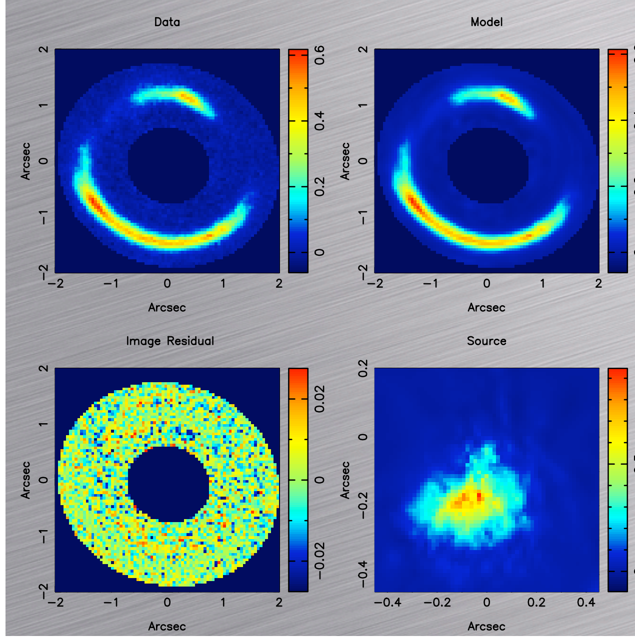
$$\sigma_v \approx 16\ km\ s^{-1}$$

$$V_{max} \approx 27\ km\ s^{-1}$$

$\Delta \log E = 65.0$ $\quad 12\ \sigma\ detection$

J0946+1066 - Double ring

Power-Law smooth model + Power-Law substructure



$$M_{\rm sub} = (3.51 \pm 0.15) \times 10^9 M_\odot$$

$$r_t = 1.1\ kpc$$

$$\Delta \log \mathcal{E} = -128.0$$

equivalent to a ~16σ detection

$$M_{\rm 3D}(<0.3) = 5.83 \times 10^8 M_\odot$$

Our results are consistent with the predictions from cold dark matter simulations at the 95 per cent confidence level, and therefore agree with the view that galaxies formed hierarchically in a Universe composed of cold dark matter. Vegetti et al. 2012 Nature 481, 341.

Conclusions

- Surface brightness anomalies can be used to find low mass galaxies at high z
- Simulations show that with HST quality data, 10 systems are sufficient to constrain the mass function
- Using high resolution adaptive optics data and the gravitational imaging technique we discovered an analogue of the Fornax satellite at redshift about 1
- The first constraints on the mass function are consistent with prediction from CDM (large errors ....)

CLUMPY STREAMS FROM CLUMPY HALOS: DETECTING MISSING SATELLITES WITH COLD STELLAR STRUCTURES

JOO HEON YOON[1*], KATHRYN V. JOHNSTON[1], AND DAVID W. HOGG[2]

ApJ Accepted

2011 ApJ 731, 58

Dynamically cold stellar streams are ideal probes of the gravitational field of the Milky Way. This paper re-examines the question of how such streams might be used to test for the presence of "missing satellites" — the many thousands of dark-matter subhalos with masses $10^5 - 10^7 M_\odot$ which are seen to orbit within Galactic-scale dark-matter halos in simulations of structure formation in ΛCDM cosmologies. Analytical estimates of the frequency and energy scales of stream encounters indicate that these missing satellites should have a negligible effect on hot debris structures, such as the tails from the Sagittarius dwarf galaxy. However, long cold streams, such as the structure known as GD-1 or those from the globular cluster Palomar 5 (Pal 5) are expected to suffer many tens of direct impacts from missing satellites during their lifetimes. Numerical experiments confirm that these impacts create gaps in the debris' orbital energy distribution, which will evolve into degree- and sub-degree-scale fluctuations in surface density over the age of the debris. Maps of Pal 5's own stream contain surface density fluctuations on these scales. The presence and frequency of these inhomogeneities suggests the existence of a population of missing satellites in numbers predicted in the standard ΛCDM cosmologies.

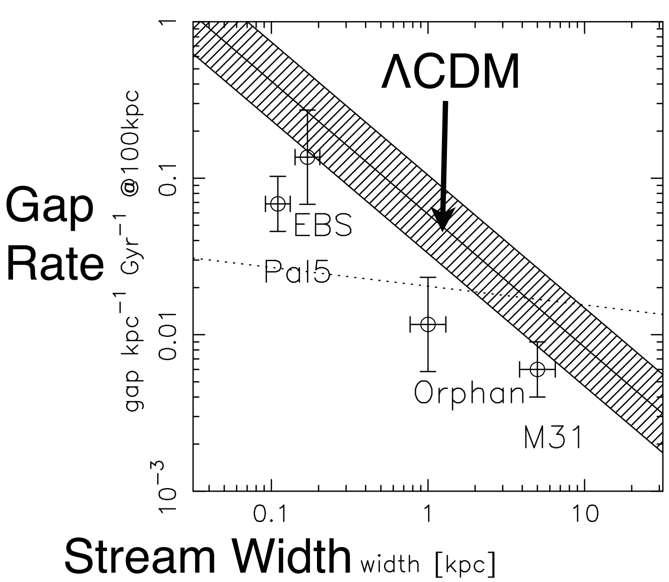


FIG. 11.— The estimated gap rate vs stream width relation for M31 NW, Pal 5, the EBS and the CDM halo prediction. All data have been normalized to 100 kpc. The width of the theoretical relation is evaluated from the dispersion in the length-height relation of Fig. 8. Predictions for an arbitrary alternative mass functions, $N(M) \propto M^{-1.6}$ normalized to have 33 halos above $10^9\,M_\odot$ is shown with a dotted line.

DARK MATTER SUB-HALO COUNTS VIA STAR STREAM CROSSINGS

R. G. CARLBERG[1] arXiv:1201.1347

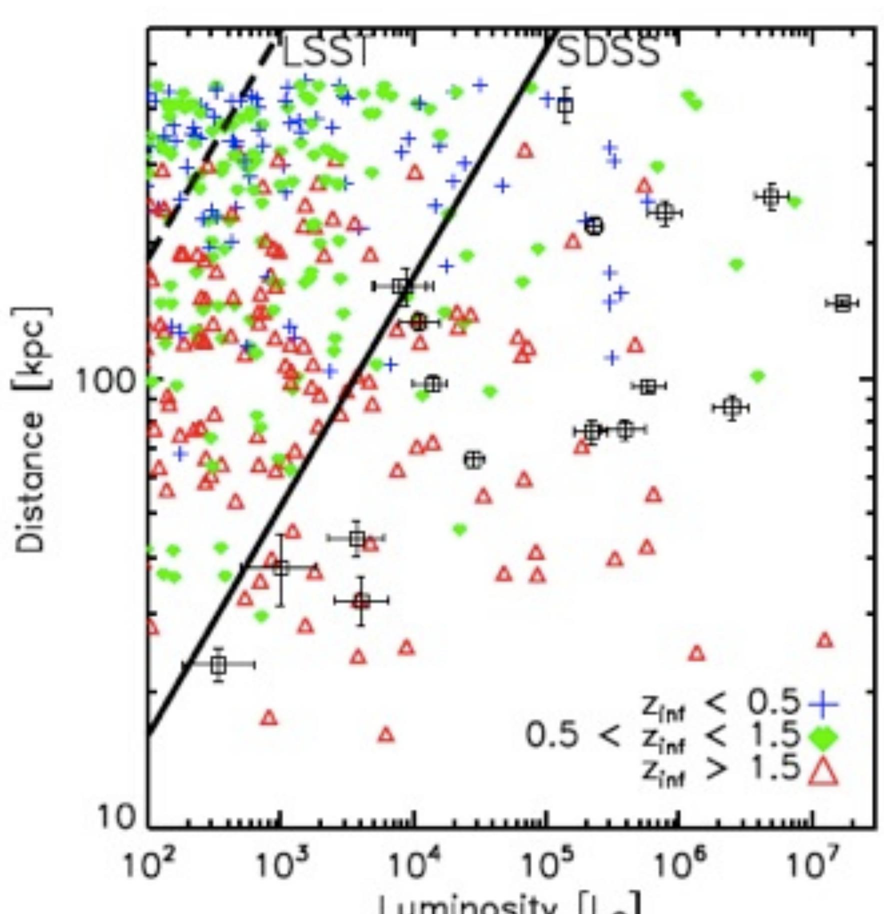
Comparison of the CDM based prediction of the gap rate-width relation with published data for four streams shows generally good agreement within the fairly large measurement errors. **The result is a statistical argument that the vast predicted population of sub-halos is indeed present in the halos of galaxies like M31 and the Milky Way.** The data do tend to be somewhat below the prediction at most points. This could be the result of many factors, such as the total population of sub-halos is expected to vary significantly from galaxy to galaxy, allowing for the stream age would lower the predicted number of gaps for the Orphan stream and possibly others as well, and most importantly these are idealized stream models.

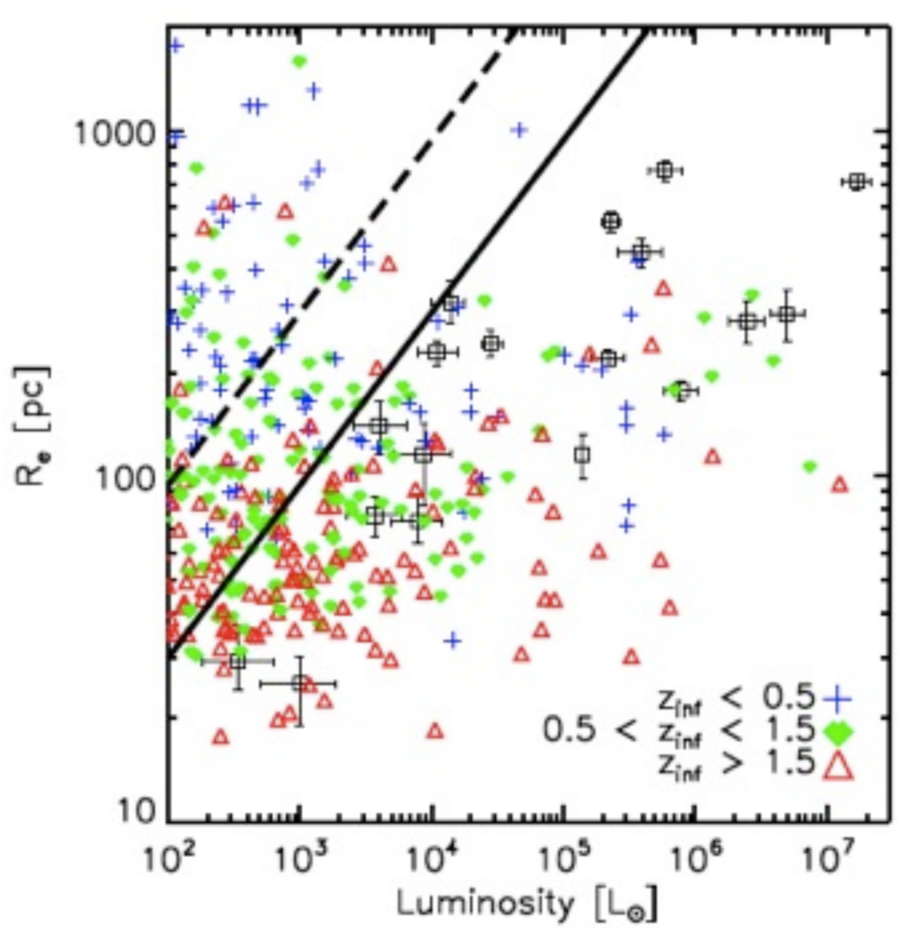
ΛCDM predicts that there is a population of low-luminosity stealth galaxies around the Milky Way. 2010 ApJ

STEALTH GALAXIES IN THE HALO OF THE MILKY WAY

James S. Bullock, Kyle R. Stewart, Manoj Kaplinghat, and Erik J. Tollerud

We predict that there is a population of low-luminosity dwarf galaxies with luminosities and stellar velocity dispersions that are similar to those of known ultrafaint dwarf galaxies but they have more extended stellar distributions (half light radii greater than about 100 pc) because they inhabit dark subhalos that are slightly less massive than their higher surface brightness counterparts. One implication is that **the inferred common mass scale for Milky Way dwarfs may be an artifact of selection bias.** A complete census of these objects will require deeper sky surveys, 30m-class follow-up telescopes, and more refined methods to identify extended, self-bound groupings of stars in the halo.

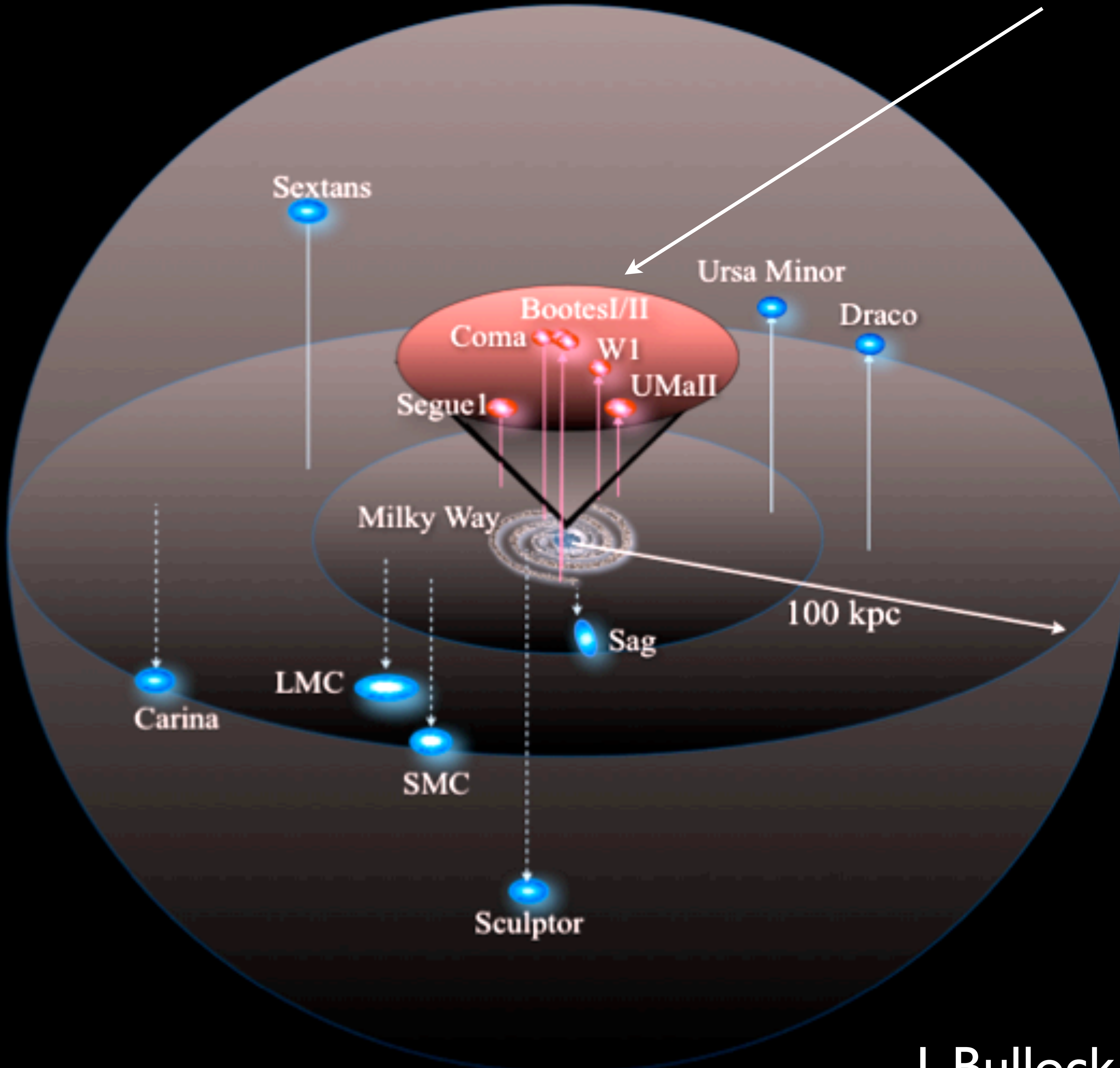




Satellite Infall redshift

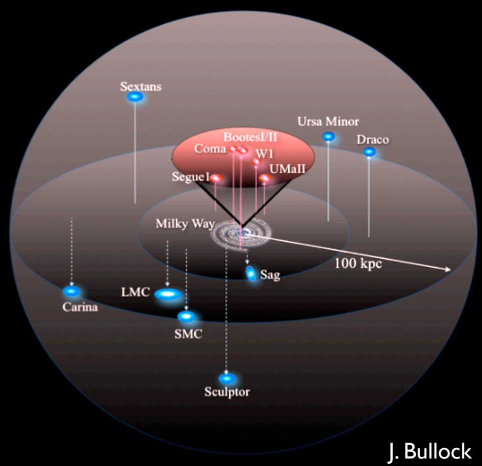
$z_{inf} < 0.5$
$0.5 < z_{inf} < 1.5$
$z_{inf} > 1.5$

SDSS satellite search

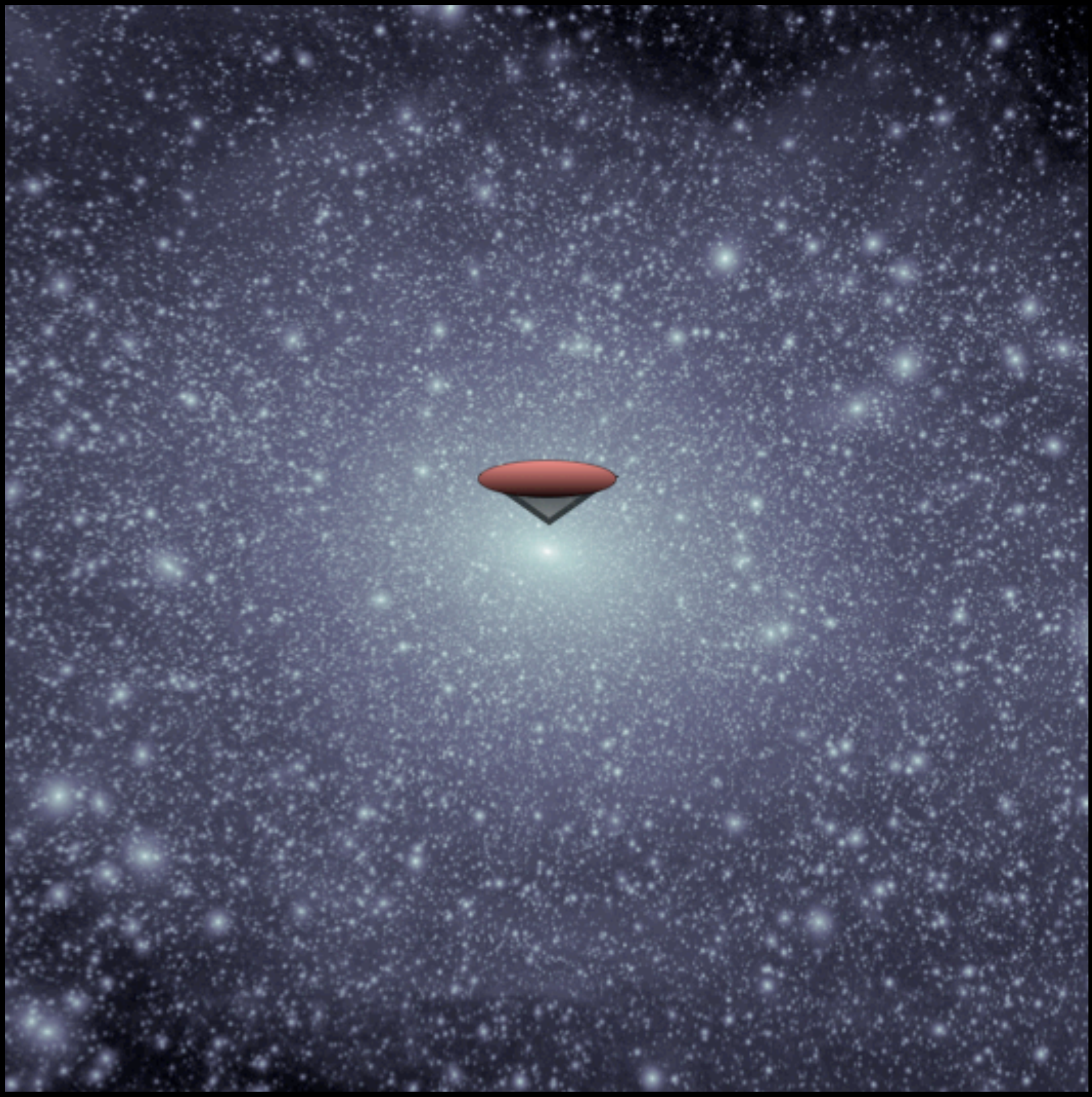


J. Bullock

The search for faint Milky Way satellites has just begun



J. Bullock



The Dark Energy Survey will cover a larger region of the Southern Sky, and LSST will go much deeper yet

small scale issues

Angular momentum

The Eris simulation shows that ΛCDM simulations are increasingly able to form realistic spiral galaxies, as resolution improves and feedback becomes more realistic.

Cusps

WDM doesn't resolve cusp issues. New observations and simulations suggest that observed velocity structure of LSB, dSpiral, dSph galaxies may be consistent with cuspy ΛCDM halos. But the "too big to fail" problem needs solution.

Satellites and Subhalos

The discovery of many faint Local Group dwarf galaxies is consistent with ΛCDM predictions. Satellites, reionization, lensing flux anomalies, gaps in stellar streams, and Lyα forest data imply that WDM must be Tepid or Cooler.

Conclusions

- CMB and large-scale structure predictions of ΛCDM with WMAP5/7 cosmological parameters are in excellent agreement with observations. There are no known discrepancies.
- On galaxy and smaller scales, many of the supposed former challenges to ΛCDM are now at least partially resolved. The "angular momentum catastrophe" in galaxy formation appears to be resolved with better resolution and more realistic feedback. Cusps can be removed by starbursts blowing out central gas.
- Lensing flux anomalies and gaps in cold stellar streams appear to require the sort of substructure seen in ΛCDM simulations. However, the biggest subhalos in ΛCDM MWy-type dark matter halos do not host observed satellites. This "too big to fail" problem appears to be the most serious current challenge for ΛCDM, and may indicate the need for a more complex theory of dark matter -- or perhaps just better understanding of DM simulations and/or of baryonic physics.

UNIVERSITY OF CALIFORNIA
HIGH-PERFORMANCE ASTROCOMPUTING CENTER

The purpose of ...h-Performance AstroComputing Center (UC-HiPACC) is to r... University of California world class resources in... ead the letter from the Director

Applications Open to Journalists for "From Planets to Cosmos" June 24-27, 2012

Applications have just opened for UC-HiPACC's Science/Engineering Journalism Boot Camp "Computational Astronomy: From Planets to Cosmos" June 24-27, 2012 for 12 to 16 practicing science journalists--the first such intensive backgrounder on astronomy to be offered on the West Coast. Two full days of sessions led by top astrophysics faculty from across the UC system and affiliated DOE labs will be held on the campus of the University of California, Santa Cruz, followed by a day visiting the computational/visualization facilities at NASA Ames Research Center and California Academy of Sciences. For details, including confirmed faculty and online application

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Journalism Boot Camp
2012 Summer School
2011 Summer School
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Press Room

Recent press releases about computational astronomy across the HiPACC consortium

March 1, 2012 - SDSC, UC HiPACC to Host Summer School on Astroinformatics...

Upcoming Events

2012 Summer School

In the News

- "A Box of Universe: Watch the cosmos evolve in a cube one billion light-years wide" about the Bolshoi simulation; by Brian Hayes, American Scientist (January-February 2012)......view article
- A dwarf galaxy that collided twice with our own Milky Way galaxy may have triggered formation of the Milky Way's spiral arms; by Chris Purcell et al., Nature (September 15,

http://hipacc.ucsc.edu

UC-HiPACC & SDSC present:

AstroInformatics

The 2012 International Summer School on AstroComputing

Apply by March 16, 2012

July 9 - 20, 2012

San Diego Supercomputer Center
University of California, San Diego

http://hipacc.ucsc.edu/ISSAC2012.html

SDSC's Gordon Supercomputer. Photo: Alan Decker.

The data available to astronomers is growing exponentially. Large new instruments and new surveys are generating ever larger data sets, which are all publicly available. Supercomputer simulations are used by an increasingly wider community of astronomers. Many new observations are compared to and interpreted through the latest simulations. The Virtual Astronomical Observatory is creating a set of data-oriented services available to everyone. In this world, it is increasingly important to know how to deal with this data avalanche effectively, and perform the data analysis efficiently. The summer school will address this analysis challenge. The topics of the lectures will include how to bring observations and simulations to a common framework, how to query largedata-bases, how to do new types of on-line analyses and overall, how to deal with the large data challenge. The school will be hosted at the San Diego Supercomputer Center, whose data-intensive computing facilities, Including the new Gordon supercomputer with a third of a petabyte of flash storage, are among the best in the world. Special access to these resources will be provided by SDSC.

Director: Alex Szalay (Johns Hopkins University)

Students must **apply** by filling in the online form at http://hipacc.ucsc.edu/ISSAC2012_Application.php

Applications are due March 16, 2012, although it may be possible to consider late applications. We aim to tell students who apply on time whether they are admitted by April 2, 2012. Upon acceptance all students who plan to attend will pay a registration fee of $300. Week day lunches, coffee breaks, the school banquet, and a special excursion will be provided for attendees.

This is the third UC-HiPACC International Summer School on Astro-Computation. The 2010 school at UCSC was on galaxy simulations and the 2011 school at Berkeley and LBNL was on computational explosive astrophysics. A key feature of the UC-HiPACC summer schools has been the access by all students to accounts on a powerful supercomputer on which the lecturers have put relevant codes and sample inputs and outputs, and the inclusion in the school of workshops each afternoon in which the students can learn how to use these tools. For the 2012 summer school on AstroInformatics, all students will have accounts on the new Gordon data-centric supercomputer at SDSC, and many relevant astronomical datasets and simulation outputs will be put on Gordon's massive FLASH memory for the use of the students.

Speakers and Topics will include:

http://hipacc.ucsc.edu

Main lecturers (3 lectures each and lead afternoon workshop):

Tamas Budavari (Johns Hopkins University) - multidimensional indexing, GPU programming, analytics
Andy Connolly (University of Washington) - time-domain information analysis, especially for the Large Synoptic Survey Telescope (LSST)
Darren Croton (Swinburne University, Melbourne, Australia) - from simulations and semi-analytic models to mock catalogs for galaxy surveys
Gerard Lemson (Max Planck Institute for Astrophysics, Garching, Germany) - the Millennium simulation databases and the evolution of virtual astronomical observatories
Risa Wechsler (Stanford University and KIPAC) - analyzing supercomputer cosmological simulations and comparing with observational surveys
Rick White (Space Telescope Science Institute, STScI) - creation and use of large astronomical databases, including the Multi-mission Archive at STScI (MAST)

Additional lecturers (1 or 2 lectures each, possibly lead afternoon workshop):

Mike Norman (SDSC / UCSD) - analyzing outputs from large ENZO cosmological simulations, the future of AstroInformatics
Peter Nugent (LBNL / UC Berkeley) - time-domain astronomy, especially the Palomar Transient Factory
Joel Primack (UCSC) - comparison of observed and simulated galaxies; evolution of high-performance astro-computing
Alex Szalay (Johns Hopkins University) - the Sloan Digital Sky Survey and the evolution of AstroInformatics
Matt Turk (Columbia University) - visualizing and analyzing observational data and simulation outputs

For further information and answers to questions please contact Coral Connor in the UC-HiPACC office (email: hipacc@ucsc.edu , phone: 831 459-1531).